

Cosmology

Master NPAC

Lesson 1 (Introductory lesson) :
History and Observational facts
Measuring distances in the Universe
Universe dynamics
What is the Universe made of ?

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2016-11-07

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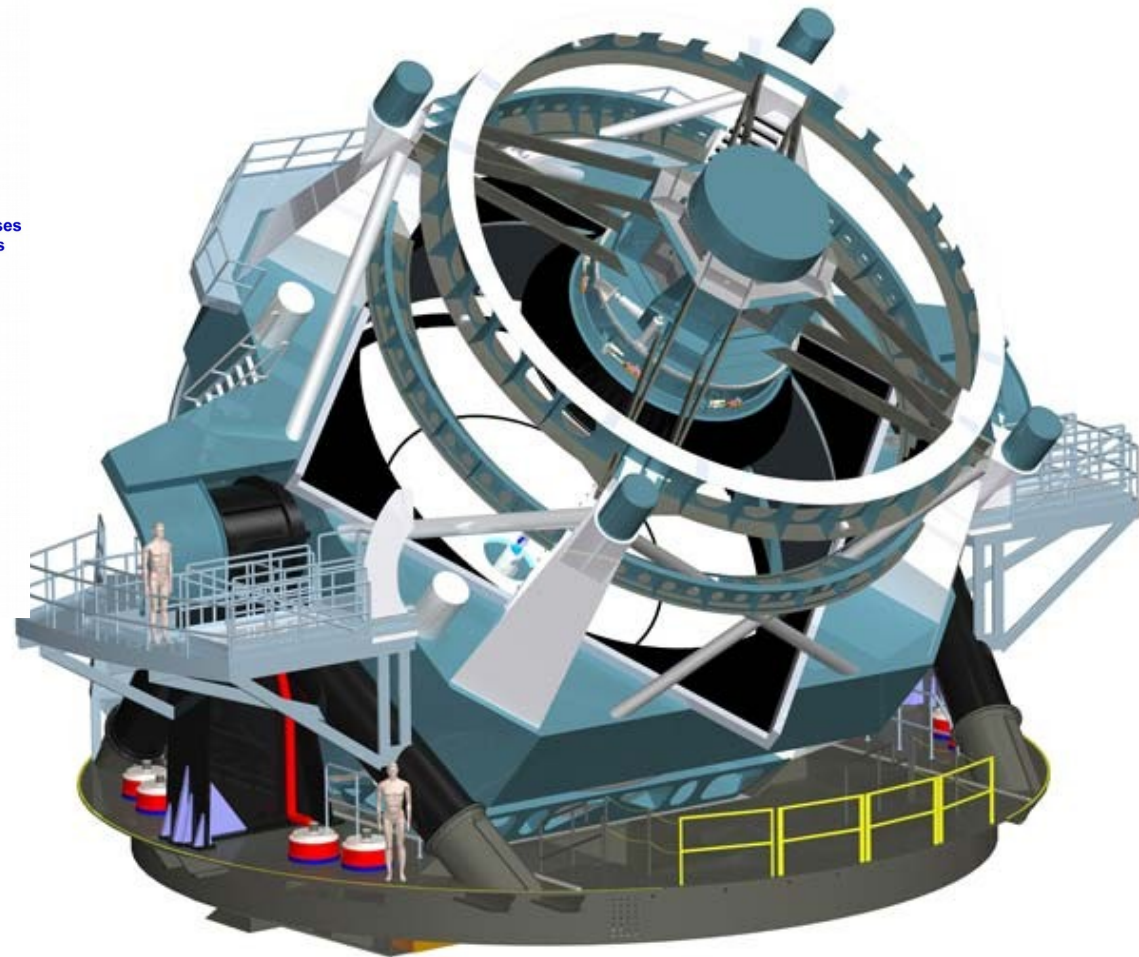
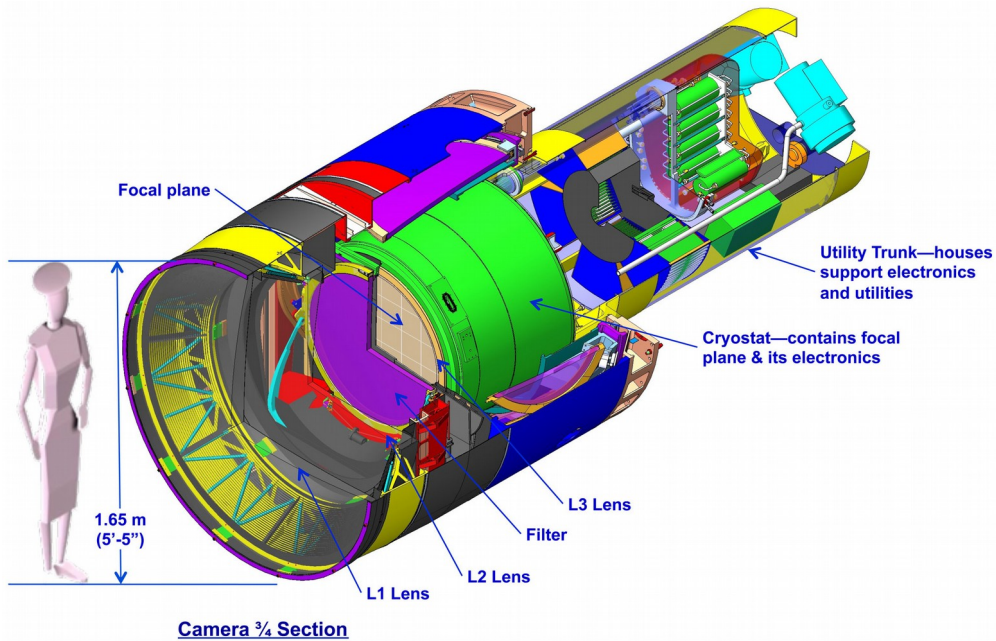
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Jussieu campus, 1213-1-12.

Observational cosmology team, LPNHE, Paris
(14 faculty/CNRS, 4 postdocs, 3 PhD students)

Cosmology with SNIa, BAO, instrumentation, calibration.

Research: Instrumentation for Cosmology (LSST), Instrumental Calibration (DICE), SN observations, Baryonic Acoustic Oscillations (DESI)

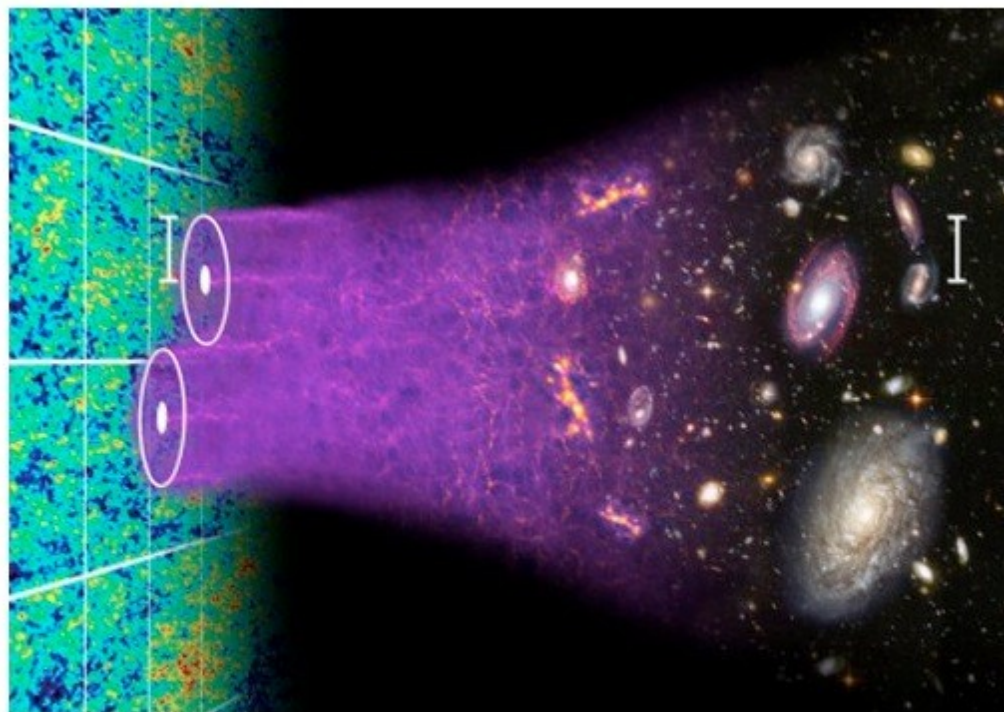
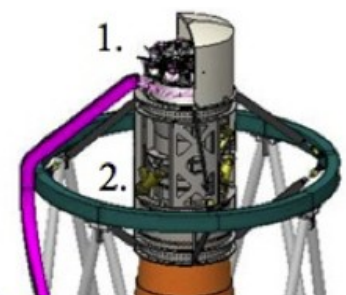
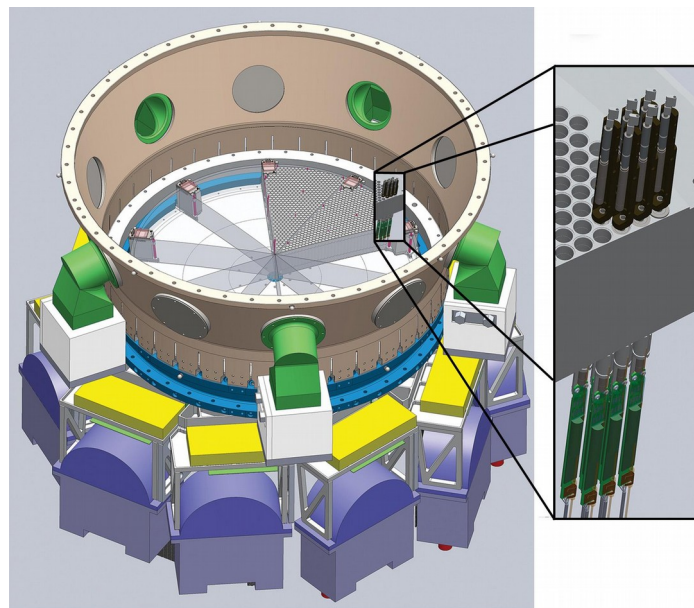
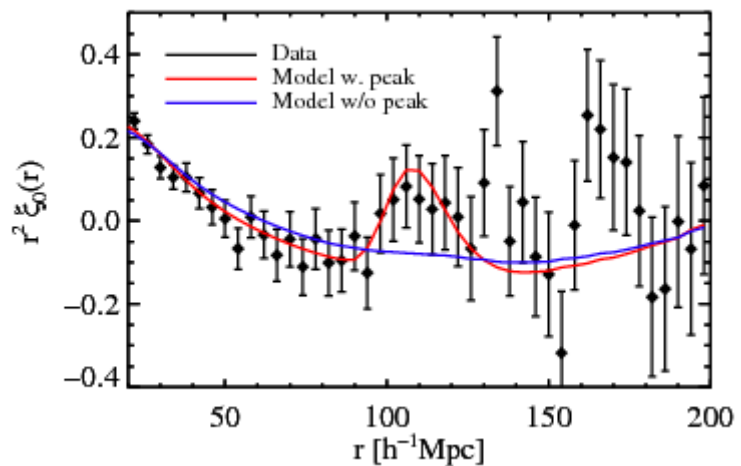


LSST, 8.4-m, 9.6 deg²

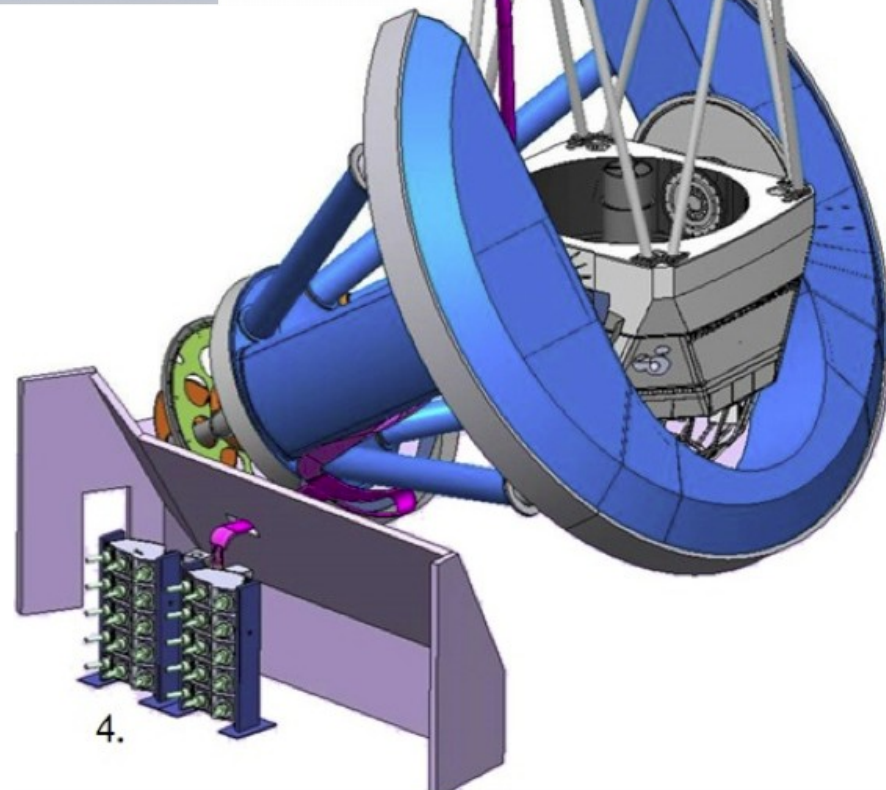
Chile, Cerro Pachon

First light ~ 2022

20000 deg², 1000 visits in 10 y



Scales of certain physical processes, imprinted on the CMB, survive to the late-time universe and can be seen in the clustering signal of galaxy positions



DESI (Mayall), 4-m, 3.2 deg²
Arizona, US, Kitt Peak (KPNO)

Contents

1.1 – A (brief) History of Cosmology

1.2 – Measuring Cosmological Distances

1.3 – Dynamics of the Universe

1.4 – Contents of the Universe

1.5 – A first sketch of the Big Bang model

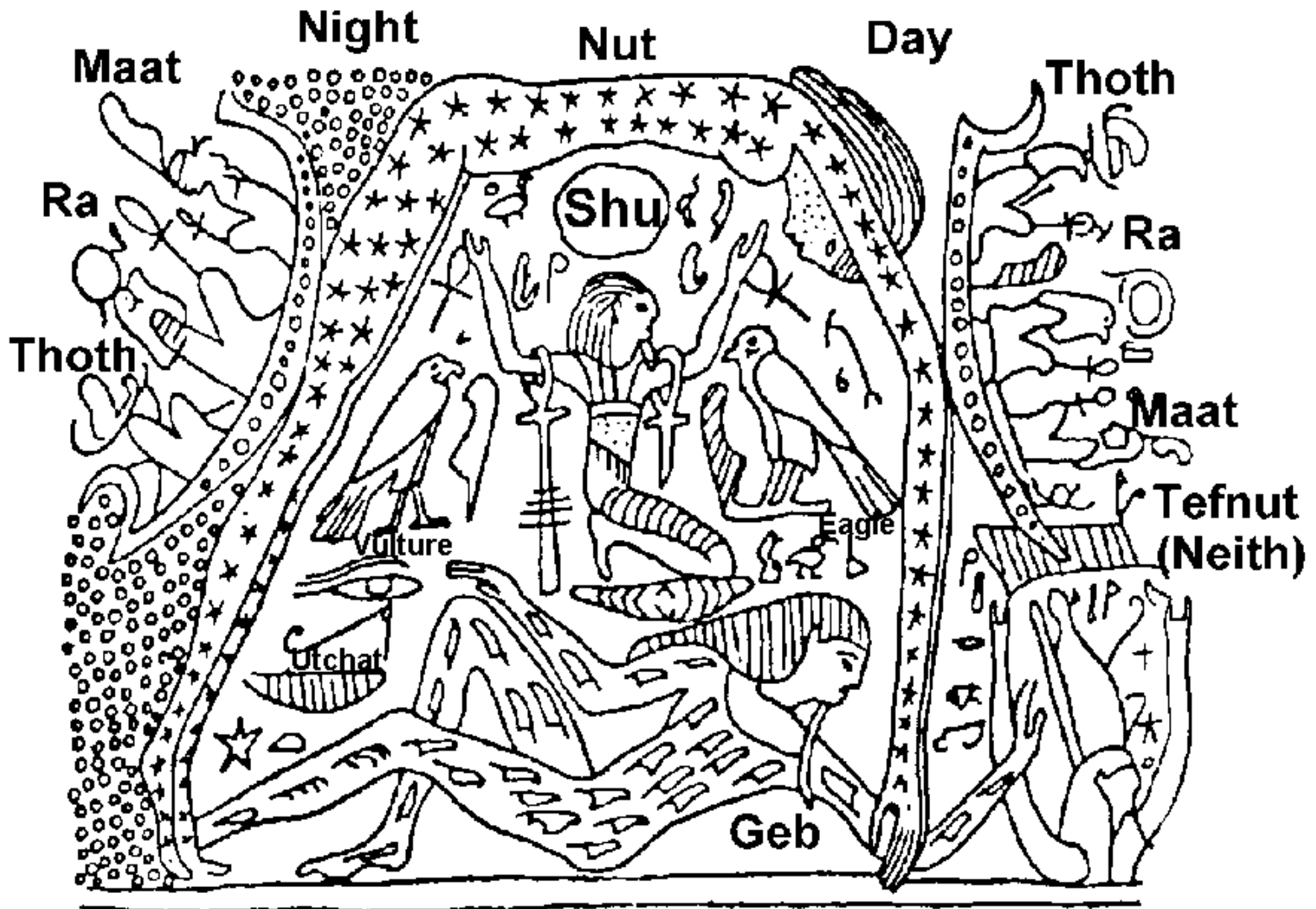
1.1

A (brief) history of Cosmology

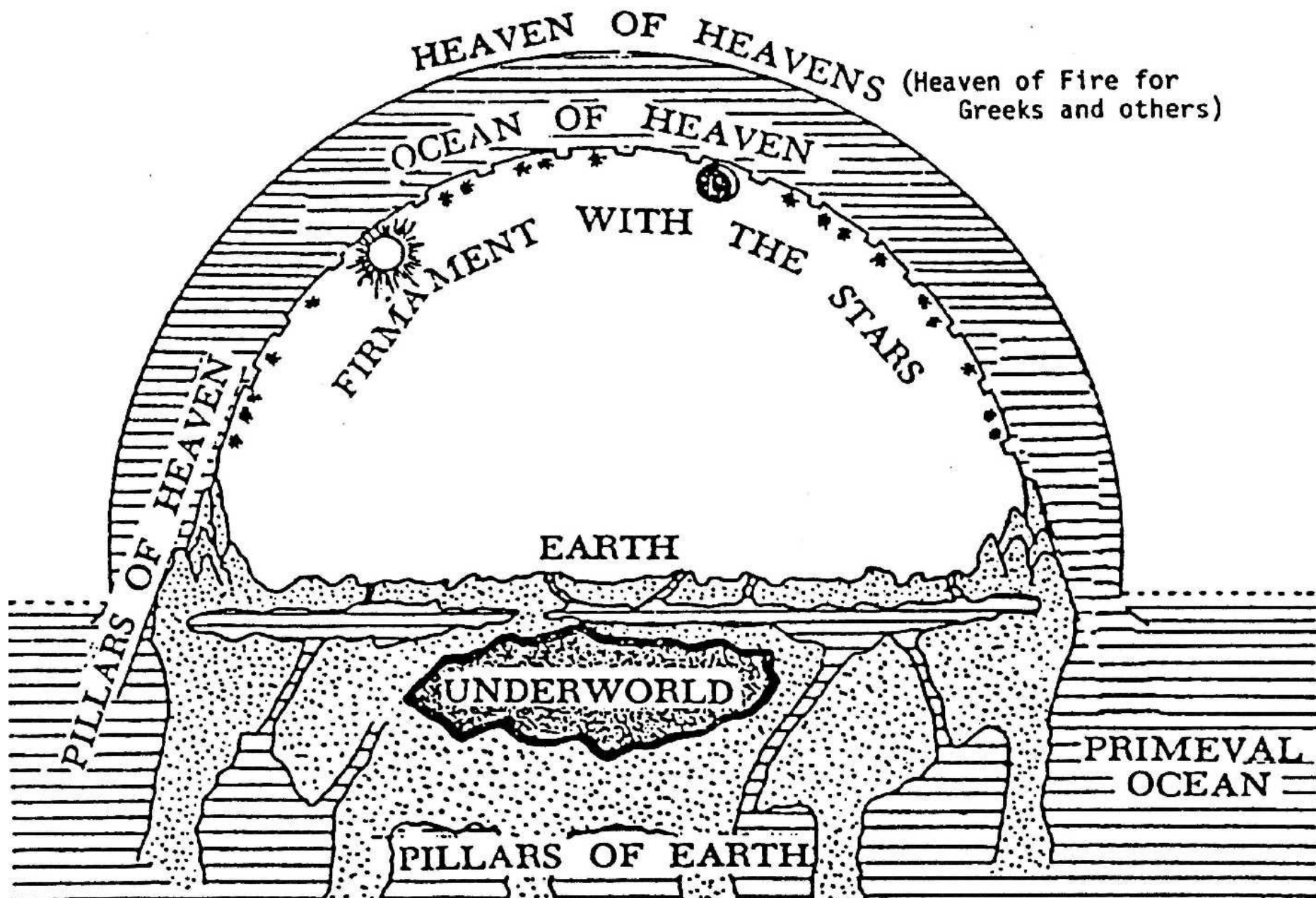
*Worldviews across the ages
Origins of Modern Cosmology*

Egyptian Cosmology





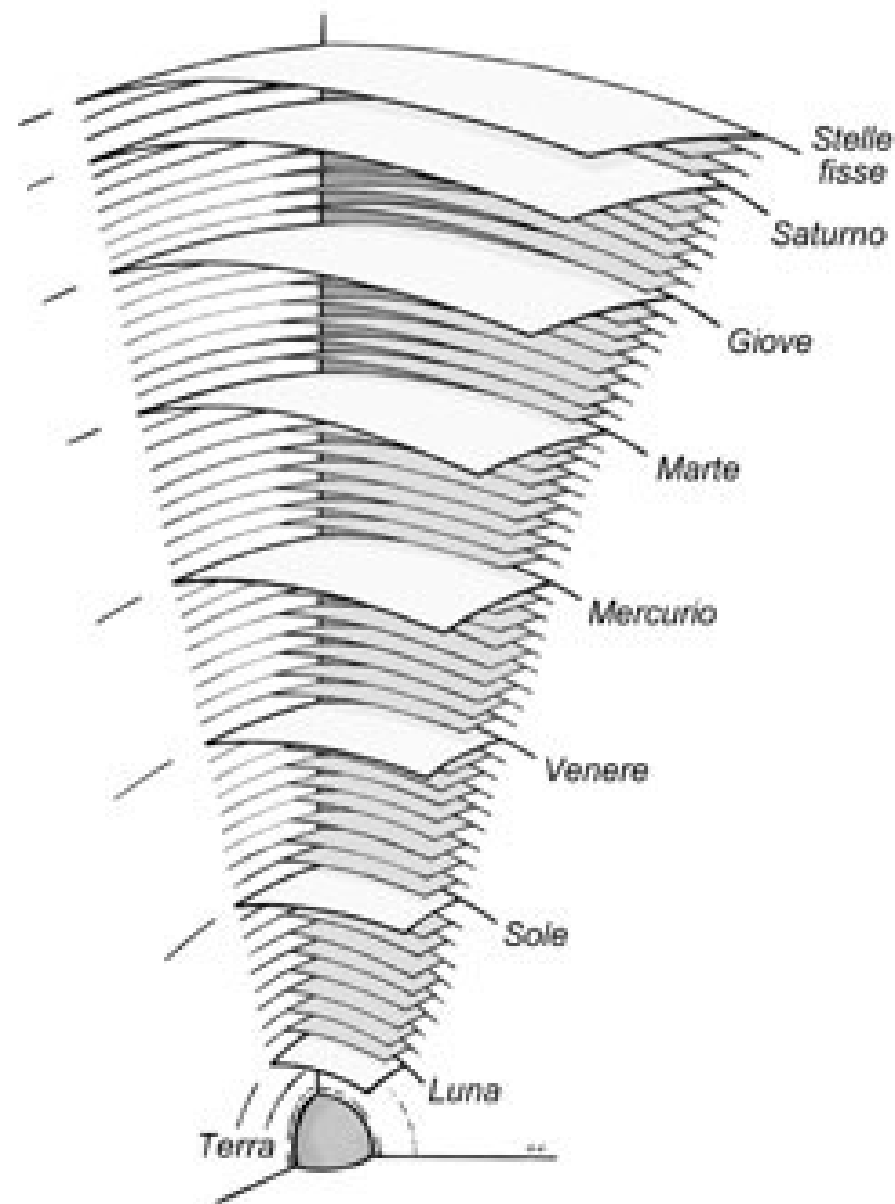
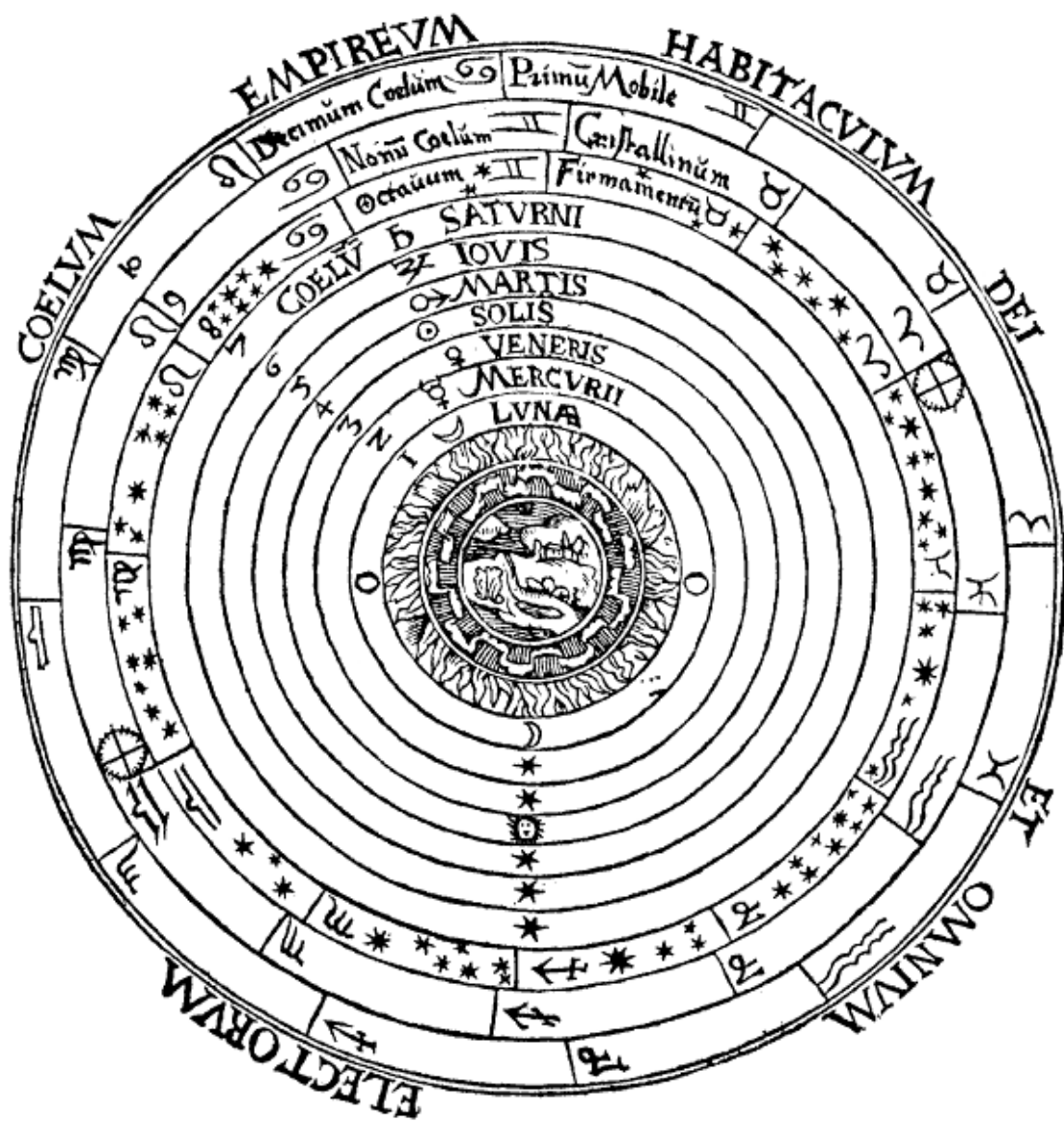
Duat of Osiris



Celestial spheres

Eudoxus of Cnidus (IV century B.C.).

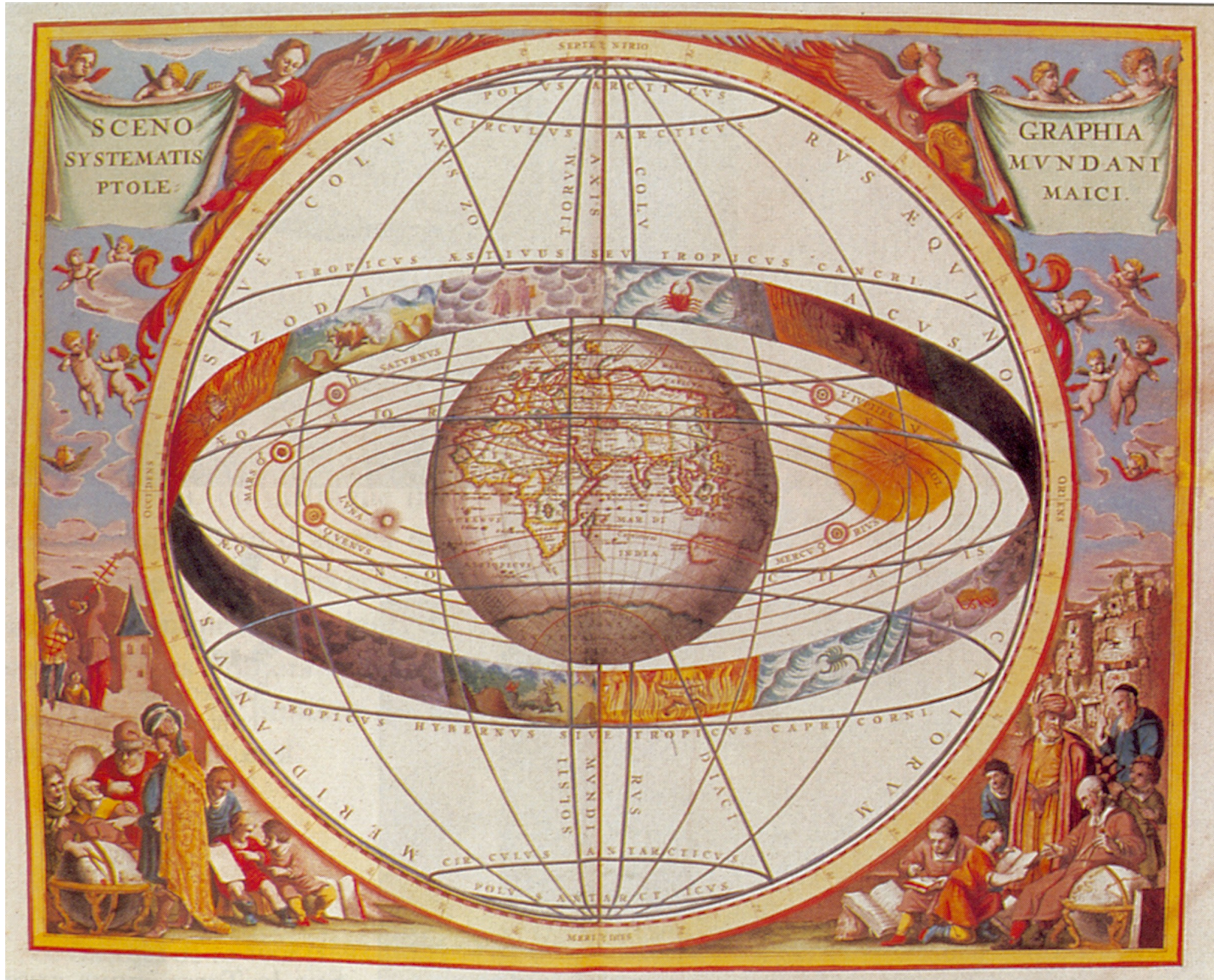
Schema huius præmissæ diuisionis Sphærarum.



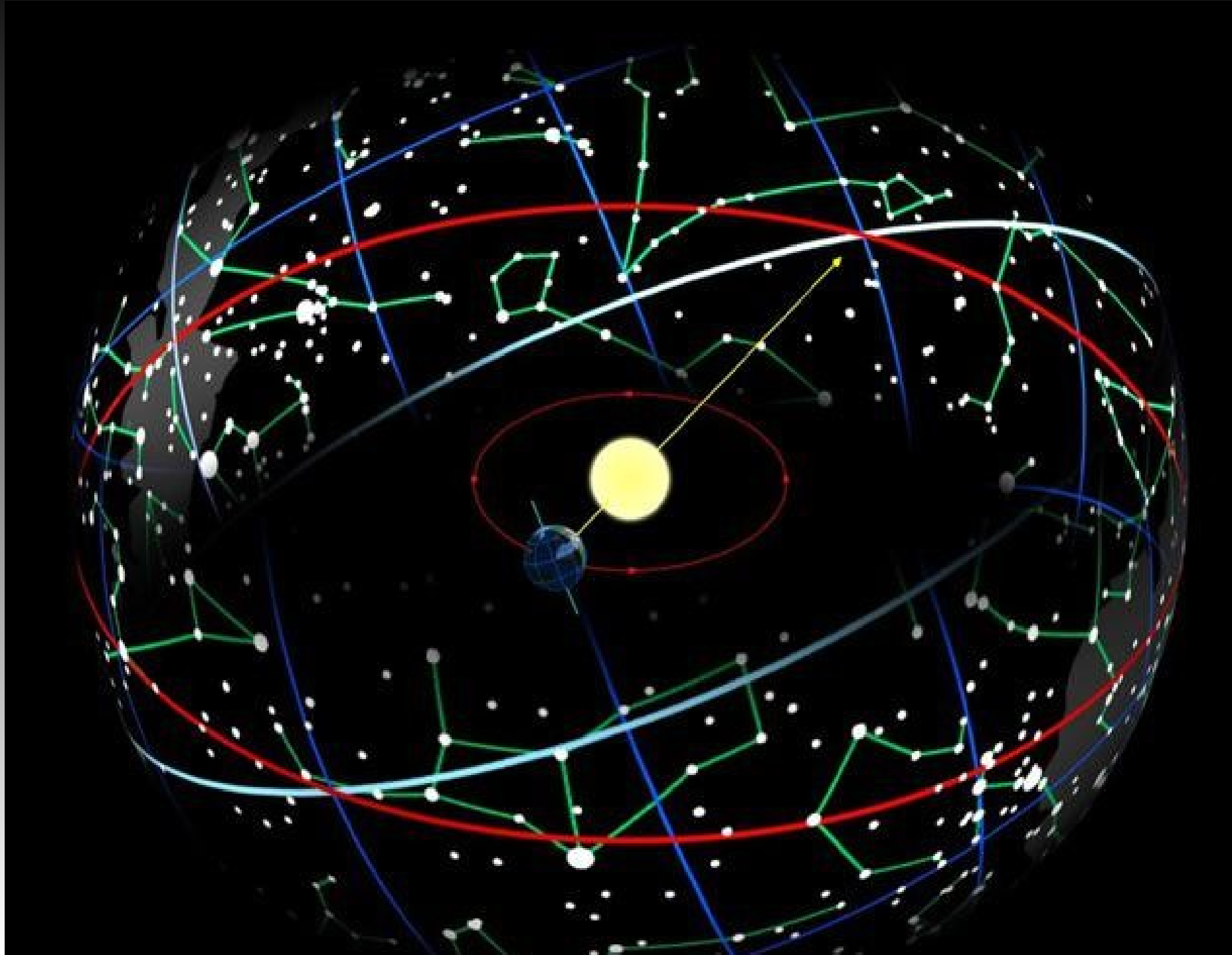
Sphere of fixed stars



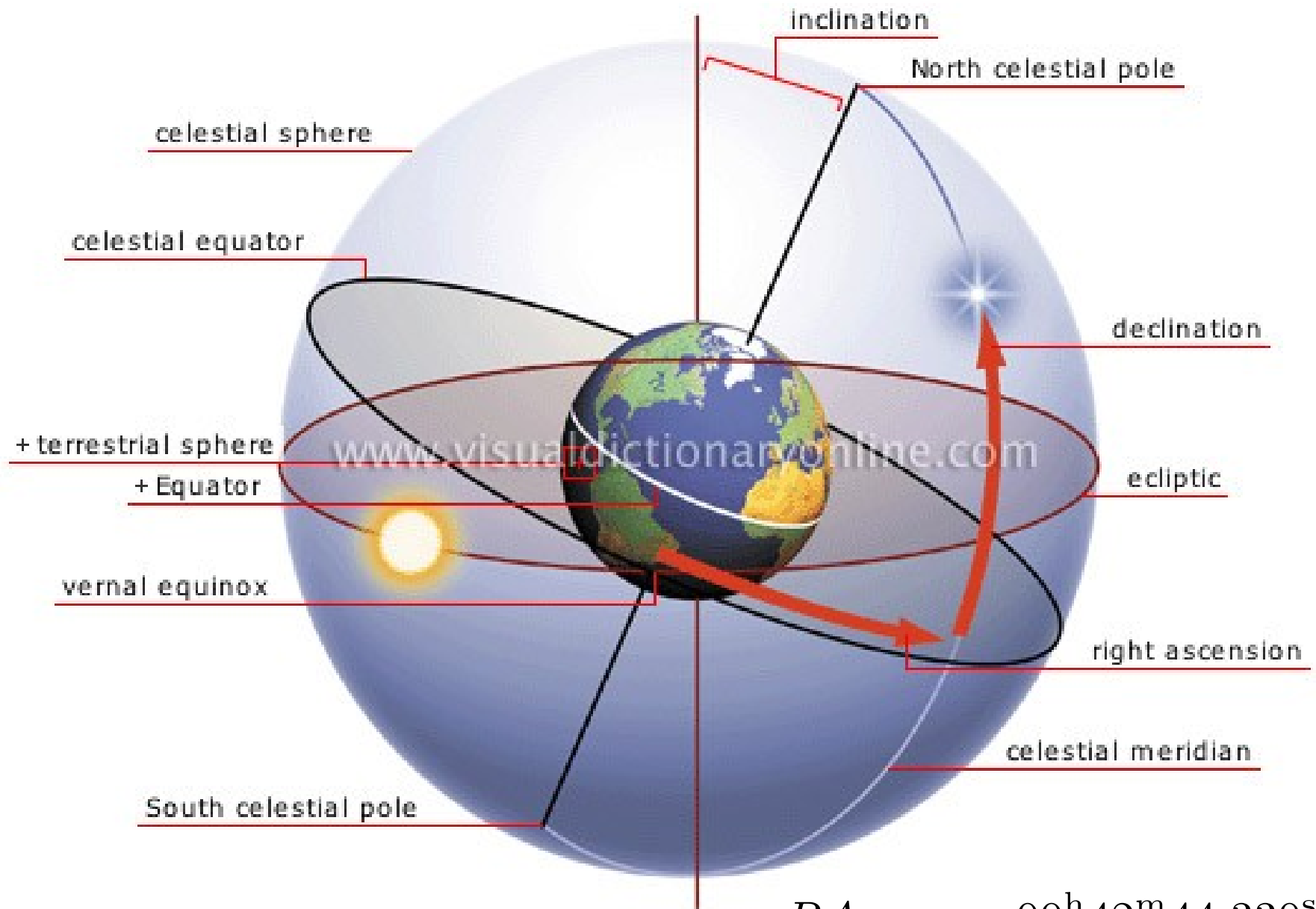
Ancient Greece (Ptolemy)



The Sphere of the « fixed » stars...



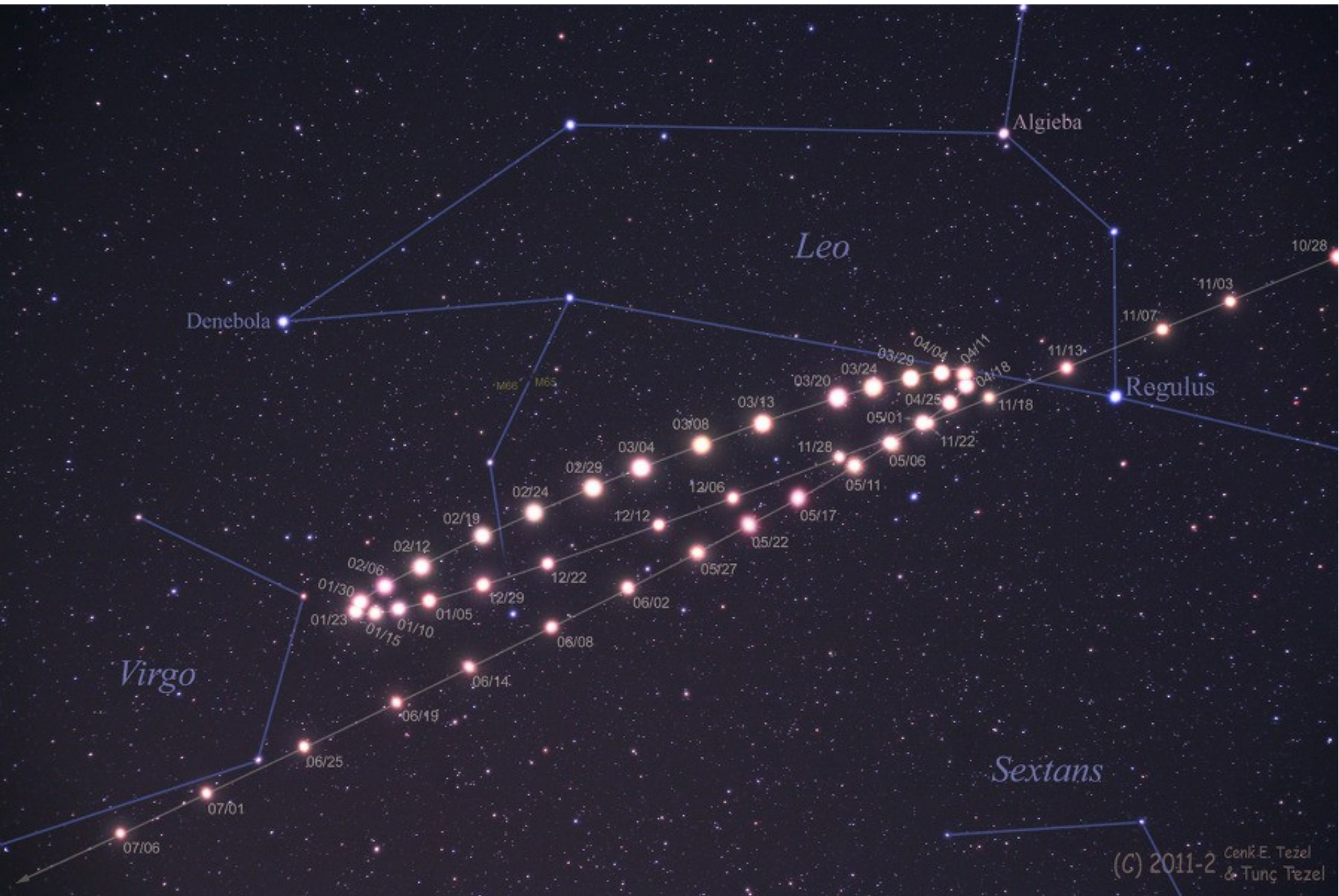
Celestial coordinates

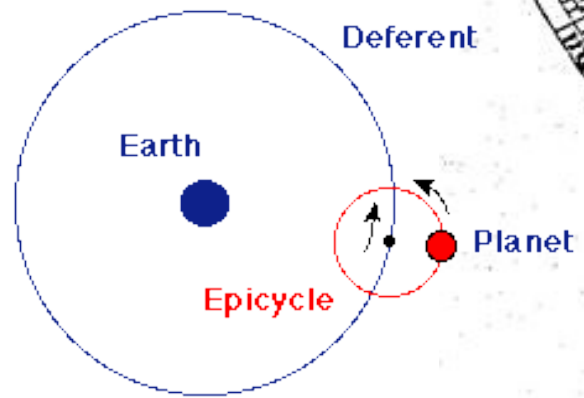
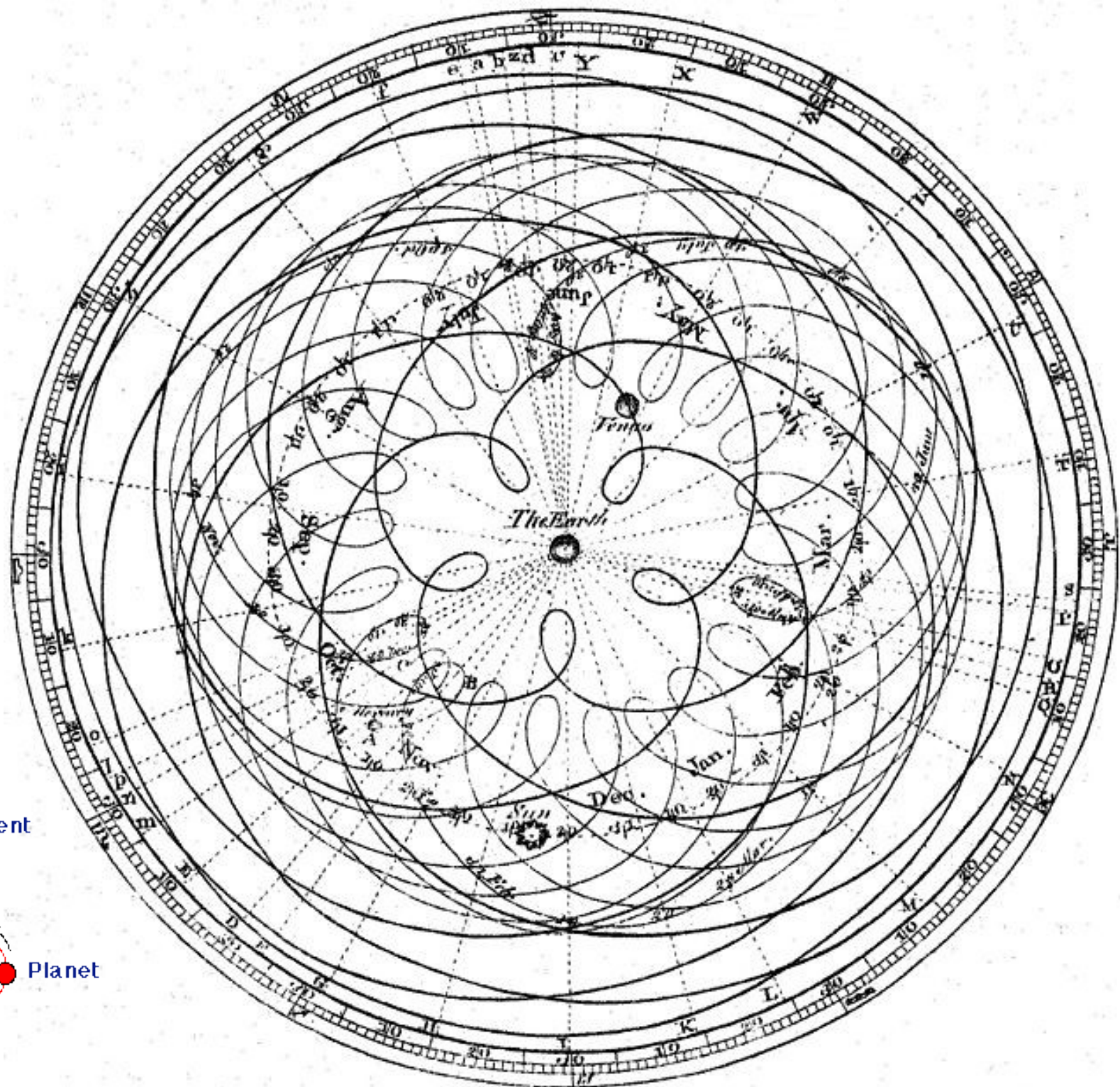


(M31) Andromeda :

$$RA = \alpha = 00^{\text{h}}42^{\text{m}}44.330^{\text{s}}$$
$$DEC = \delta = +41^{\circ}16'07.50''$$

Retrograde movement of Mars





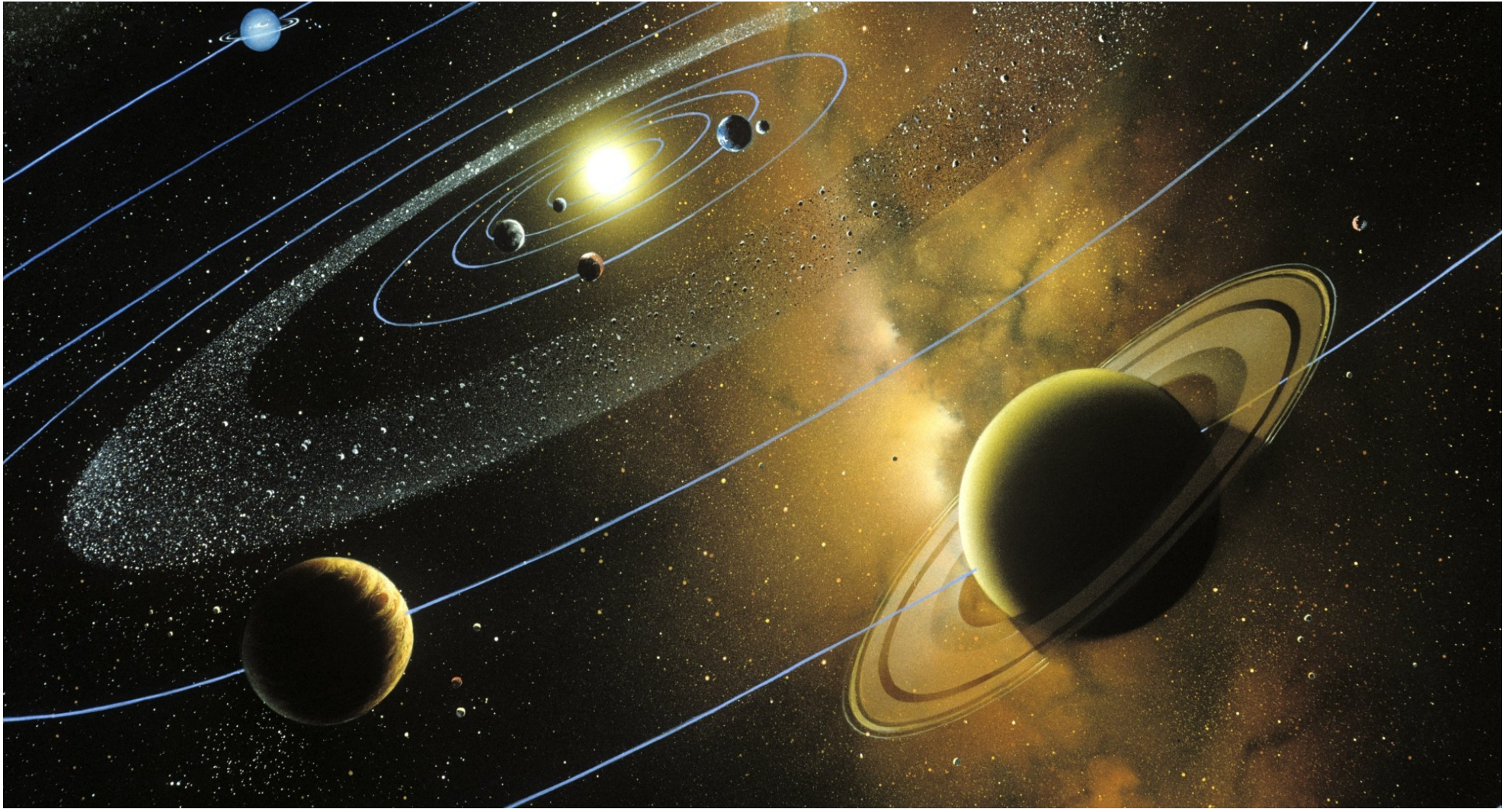
Copernician revolution

Nicolaus Copernicus (1473 – 1543), J. Kepler, ...



The Solar System

A modern view



7 planets + Uranus (Herschel, 1781) + Neptune (Galle/Le Verrier 1846) + Pluton (1930)



Castiopeja.

PERSEUS.

Camelopardalus.

Andromeda.

Auriga.

Triangulum Majus.

Minus.

Musca.

Aries.

Taurus.

Fig. W.

Latitude.

Longitude.

Longitude.

Latitude.

Latitude.

Longitude.

Latitude.

Longitude.

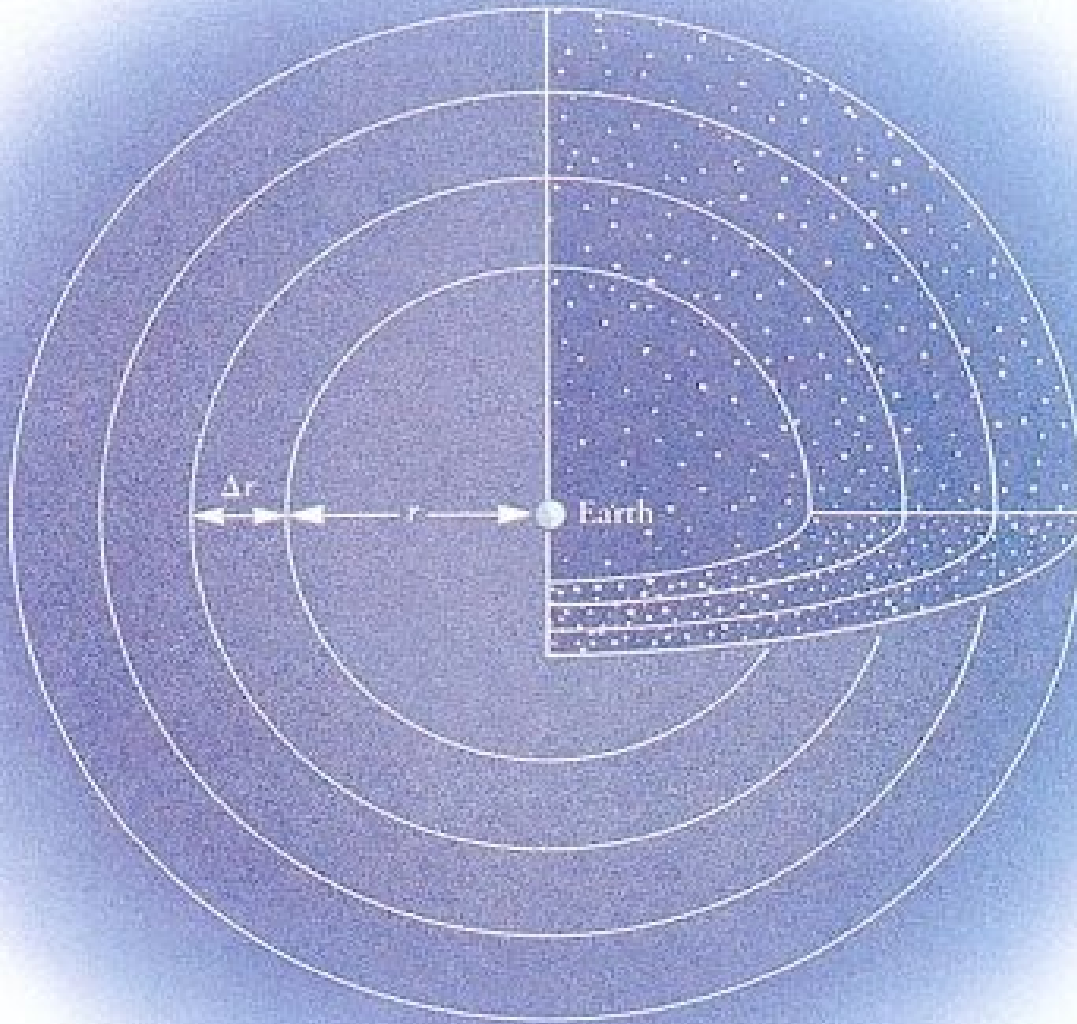
Latitude.

Olbers paradox

J. Kepler (1610), H. Olbers (1823), ...

Why is the night sky dark ?

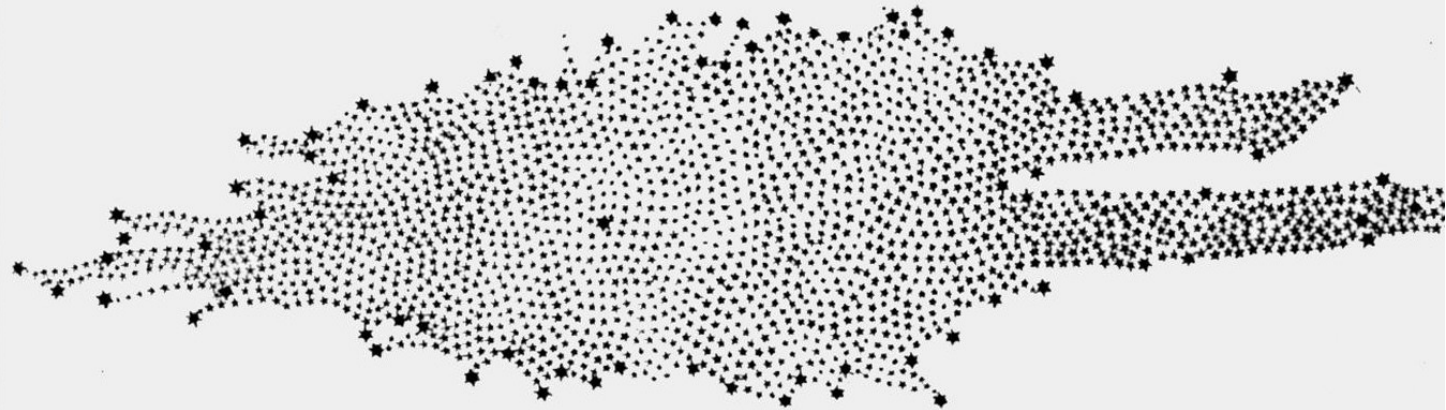
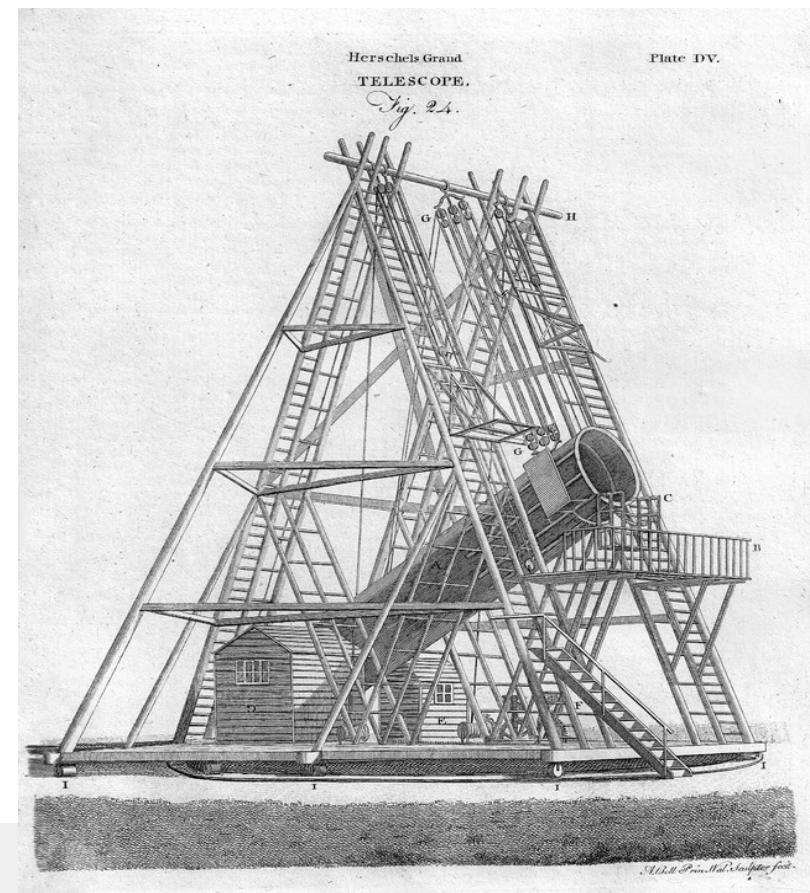
In an infinite and static universe, the sky should be as bright as the sun surface



Our « galaxy »

W. Herschel suggests (1781) from star counts that the shape of the « universe » may be lenticular.

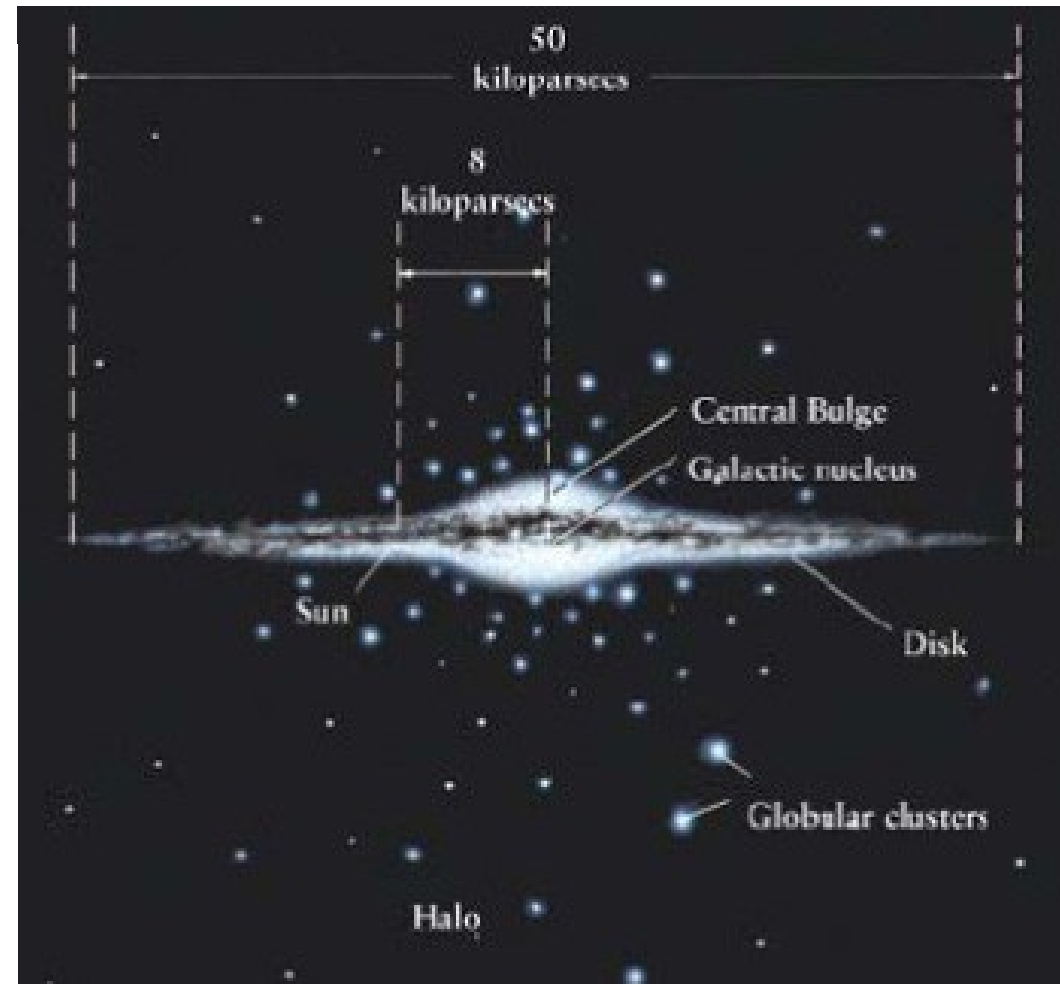
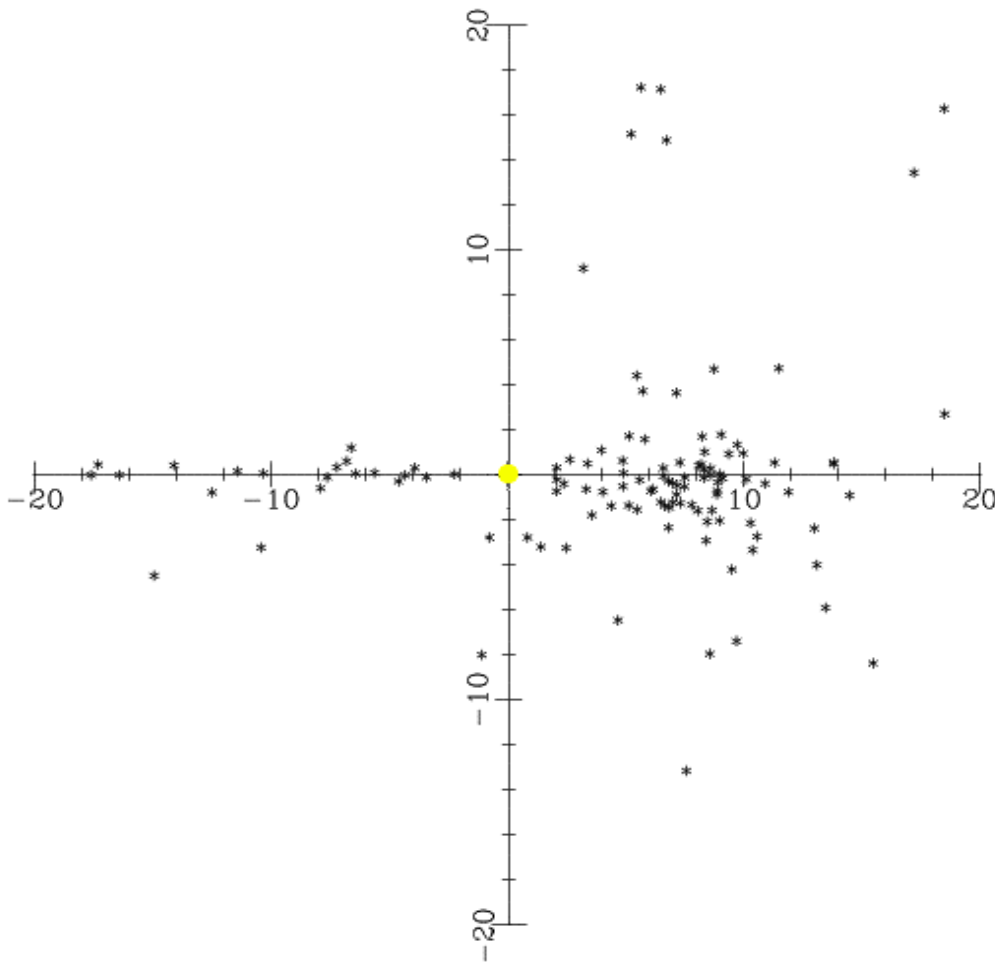
First hints about our own galaxy, and extra galactic objects (« nebulae »).



Our position in the Galaxy

H. Shapley (1885 – 1972)

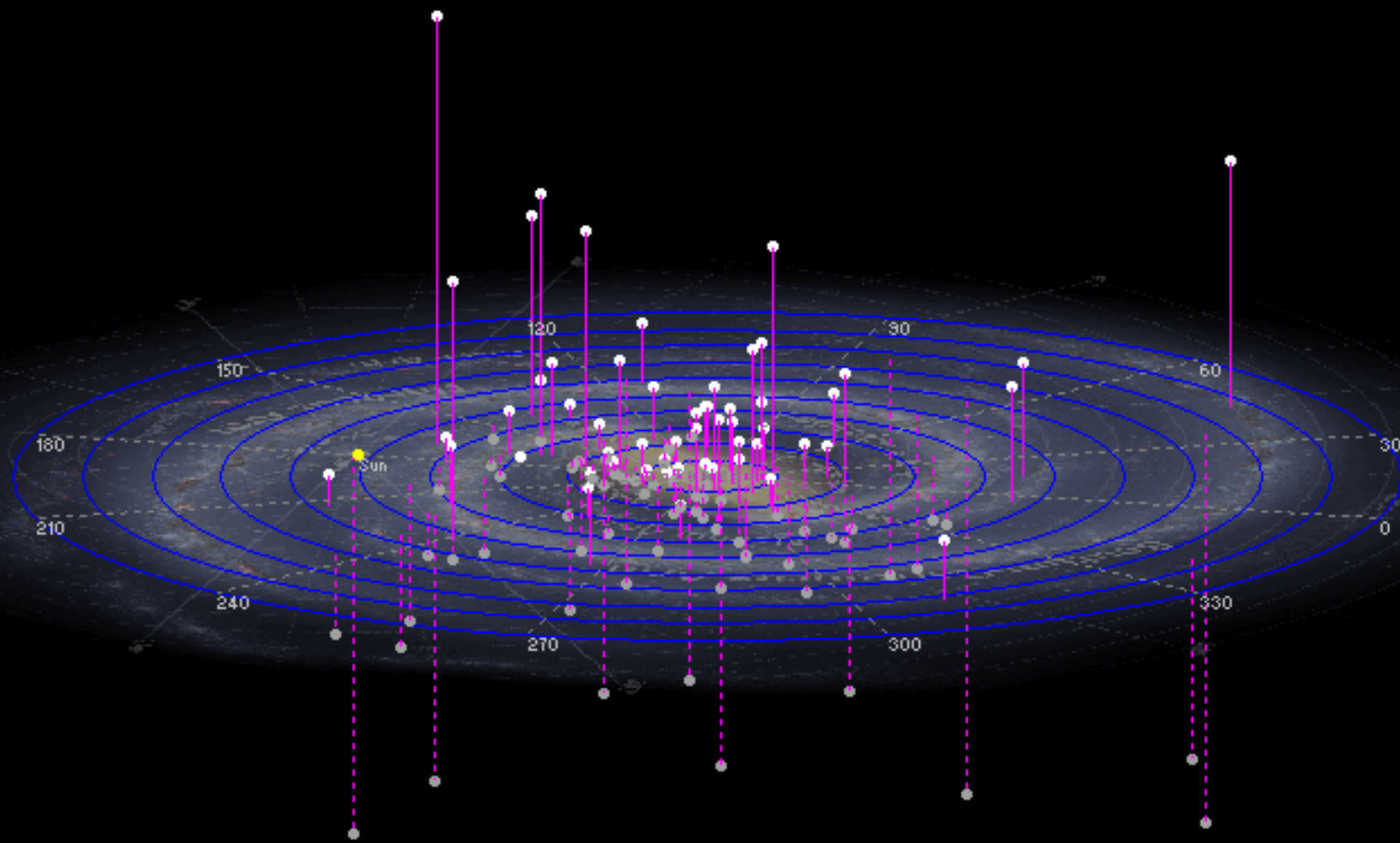
From the distribution of the globular clusters
(distances from RR Lyrae / cepheids)



The 119 globular clusters within 50,000 LY of the galactic centre

Galactic centric (galactic longitude and latitude)

5,000 LY



The Great debate (1920)

H. Shapley (1885 – 1972) – H. Curtis (1872 – 1942)



Are the fuzzy faint « nebulae » extra galactic objects ?

Are they « island-universes » similar to our own ? (Shapley)

Or are they objects close to us ?

Does the « galaxy » constitute the entire Universe ? (Curtis)

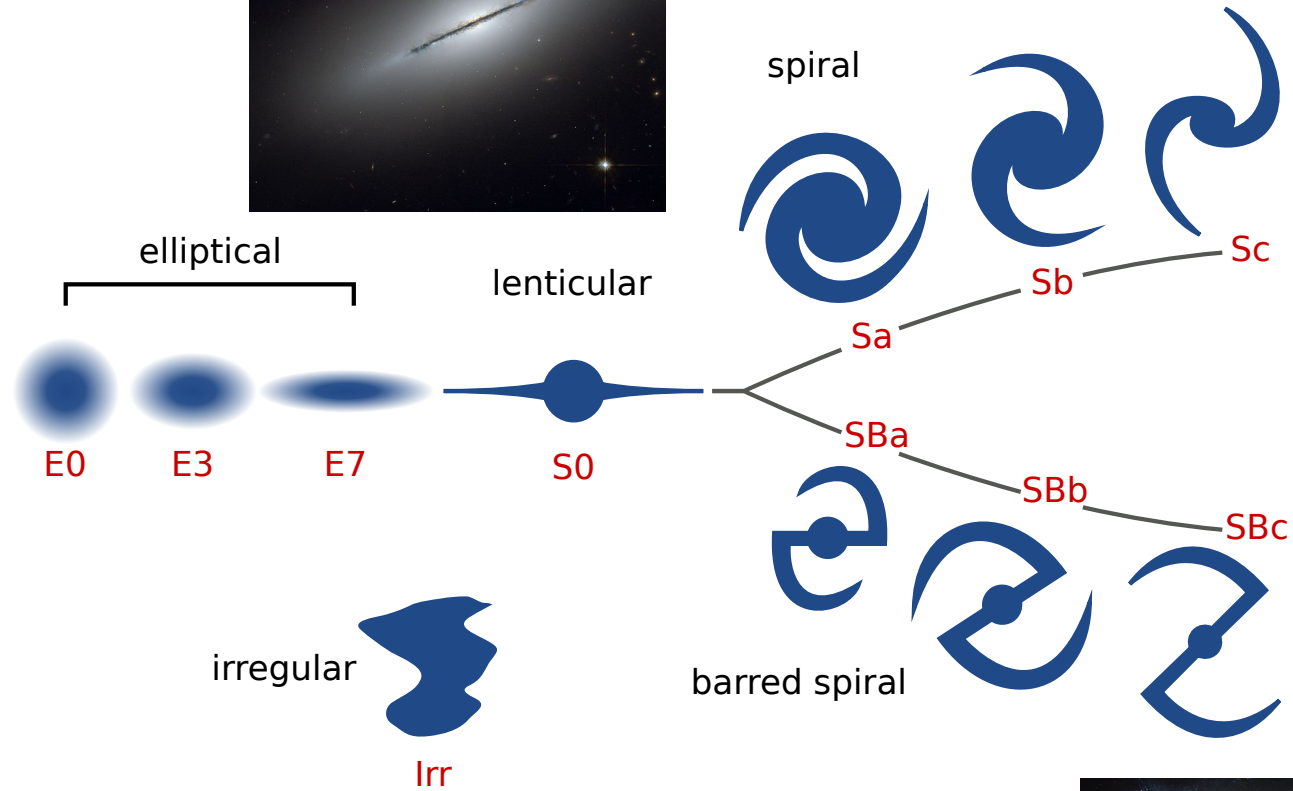
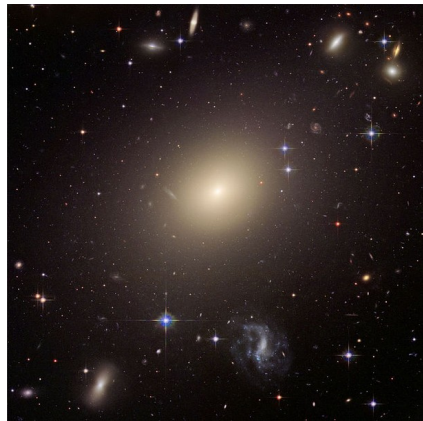
Mid-1920s → Hubble measured the distance to M31
using Cepheids



Conclusion : the Universe is much bigger !

Classification of the galaxies

E. Hubble, M. Humason



Hubble's law

Redshift z :

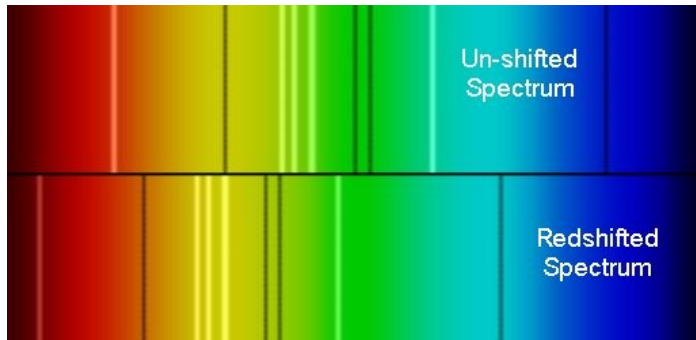
$$\lambda_{\text{obs}} = \lambda_{\text{emit}}(1 + z)$$

$$z = \frac{\lambda_{\text{obs}}}{\lambda_{\text{emit}}} - 1$$

« Doppler » effect :

$$\frac{\lambda_{\text{obs}}}{\lambda_{\text{emit}}} = \sqrt{\frac{1 + \beta}{1 - \beta}} \simeq 1 + \beta$$

$$z \simeq \beta \quad v \simeq cz$$



Galaxy in the constellation

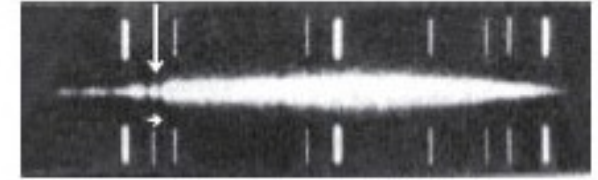
approximate distance/ 10^6 ly

redshift



Virgo

55

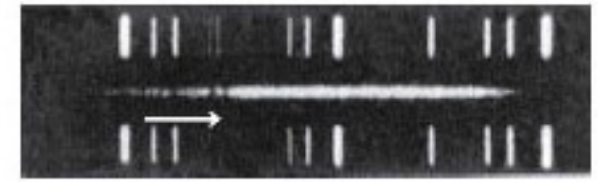


0.004



Ursa Major

700

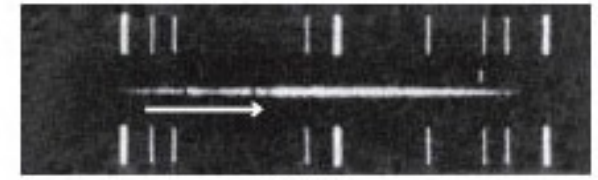


0.050



Corona Borealis

1000



0.073

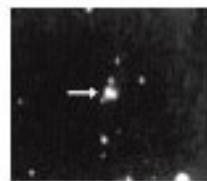


Bootes

2000

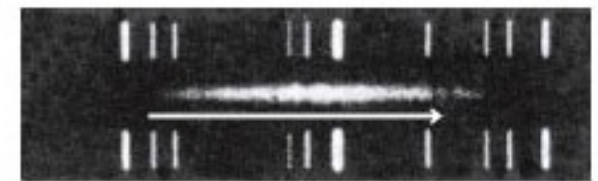


0.130



Hydra

3000



0.203

Hubble's law $v = cz = H_0 d$

(1929)

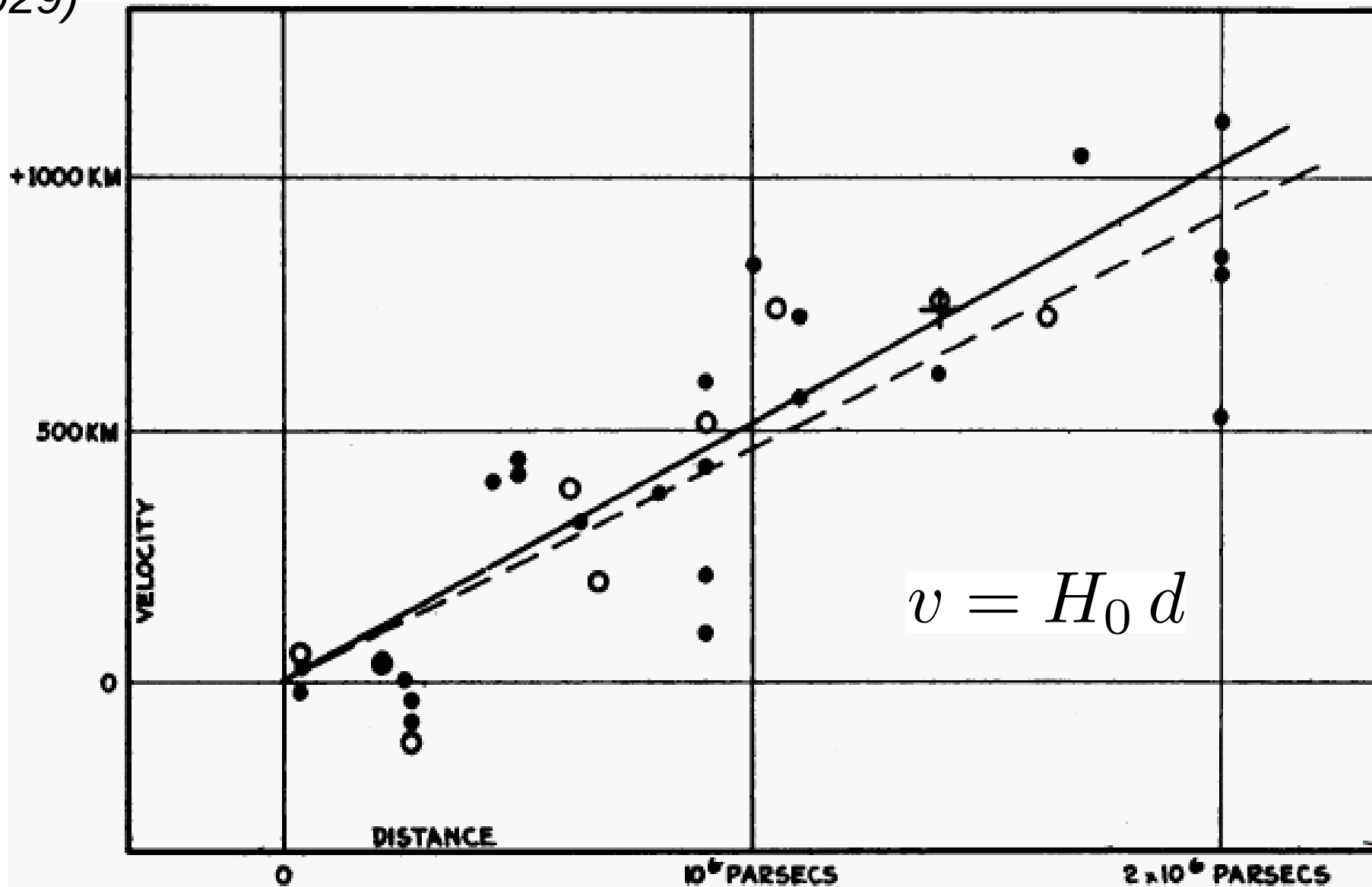


FIGURE 1

Velocity-Distance Relation among Extra-Galactic Nebulae.

The Universe is expanding !

E. Hubble, G. Lemaître, A. Friedmann, ...

The point of view is the same
for any galaxy : no center

Isotropy and homogeneity at
large scale

Typical age of the Universe :
Hubble time :

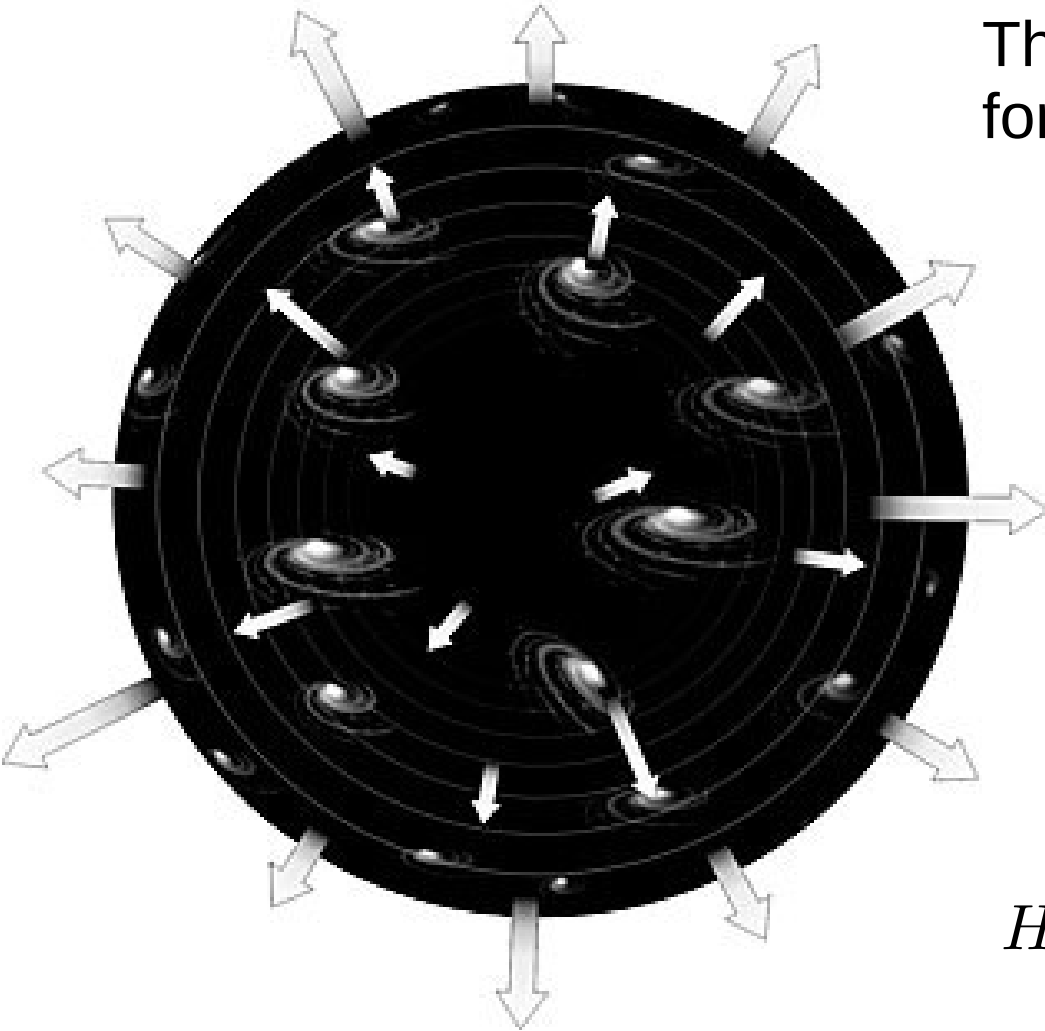
$$t_0 = 1/H_0$$

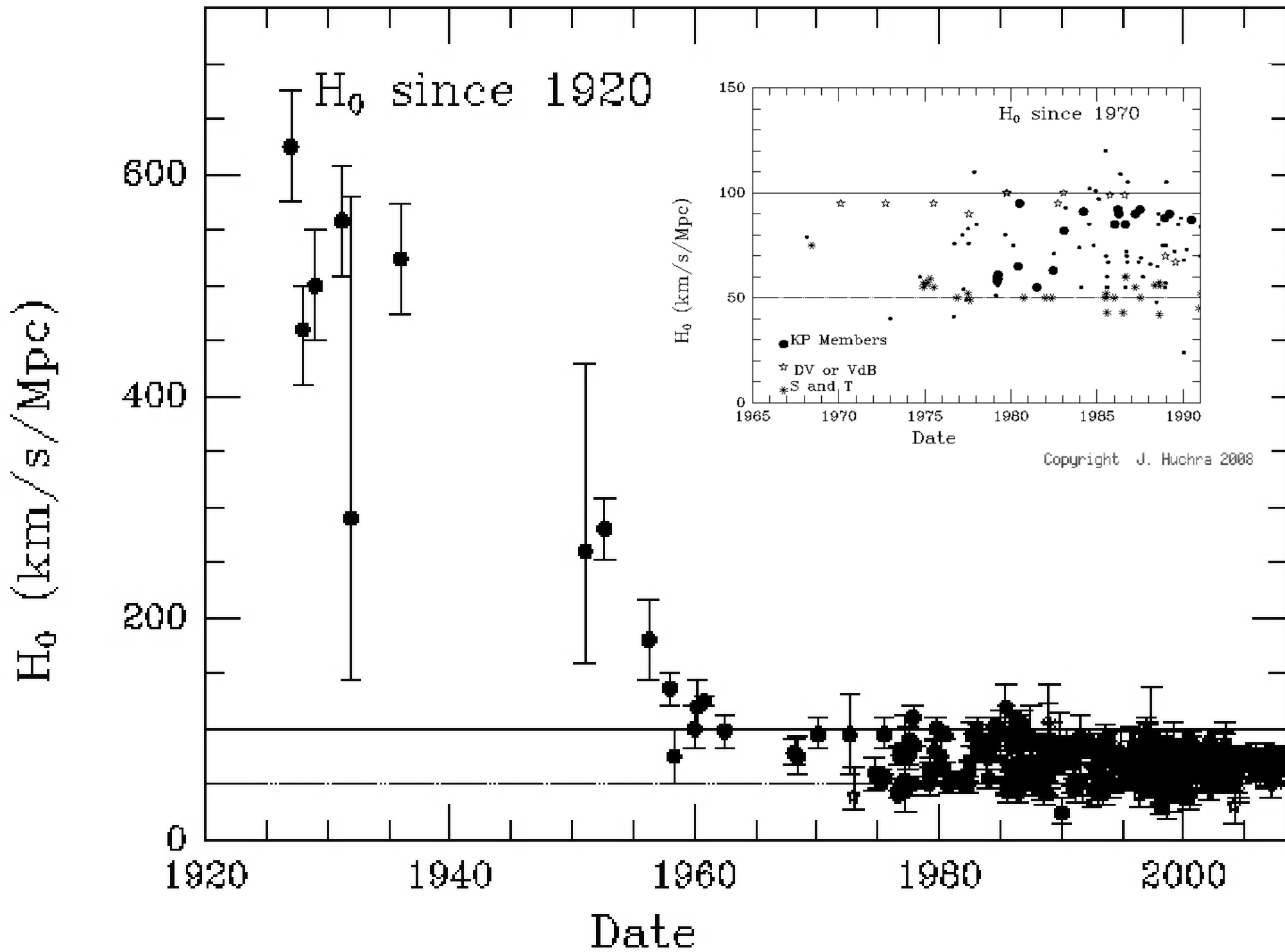
Hubble initial value was much
too high :

$$H_0 \simeq 500 \text{ km/s/Mpc} \quad t_0 \simeq 2 \text{ Gy}$$

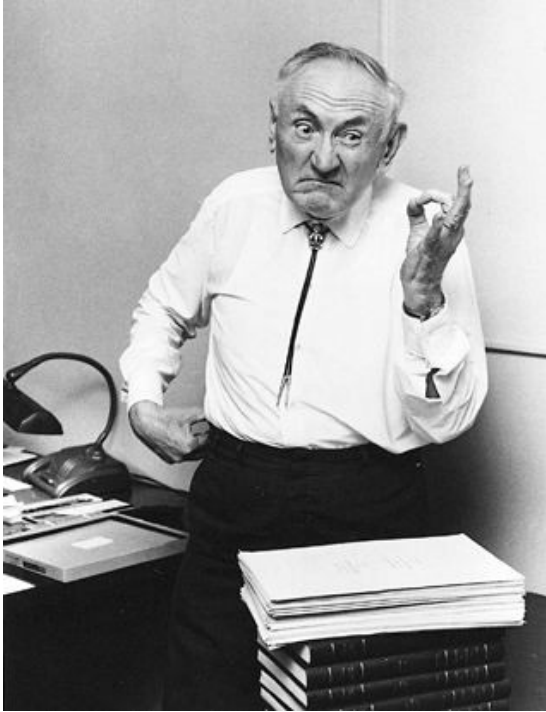
Expanding Universe : The Universe should have been more compact and warmer in the past.

One answer to Olbers' paradox : the age of the Universe is finite





Alternative theories (« Every good story needs a villain »)



F. Zwicky (1898 – 1974)

Theory of the « *tired light* » to explain redshift from far away objects

Prediction of neutrons stars, discoverer of many supernovae.

Original mind, hard to work with...

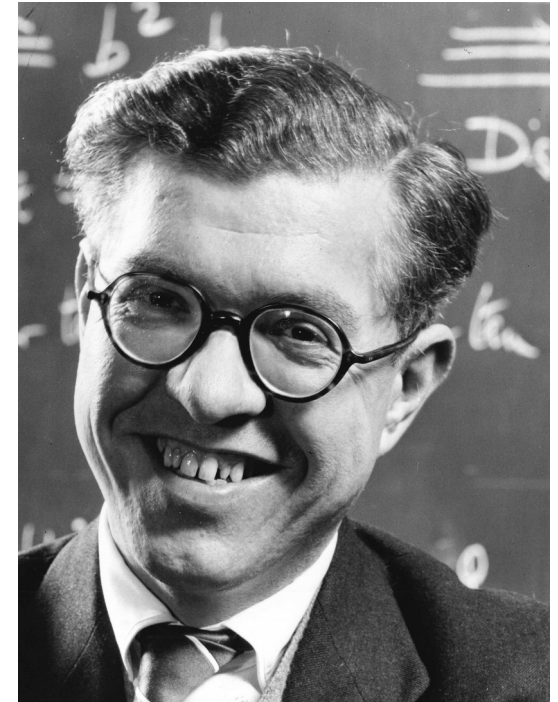
F. Hoyle (1915 – 2001)

First papers on nucleosynthesis in stars.

Predicted the triple alpha reactions.

Strong detractor of the « Big Bang » theory.

Alternative theory of matter creation to compensate expansion (« *steady state* »).



Cosmic Microwave Background (CMB)

A. A. Penzias & R. W. Wilson (1965)

Black Body spectrum (microwave)

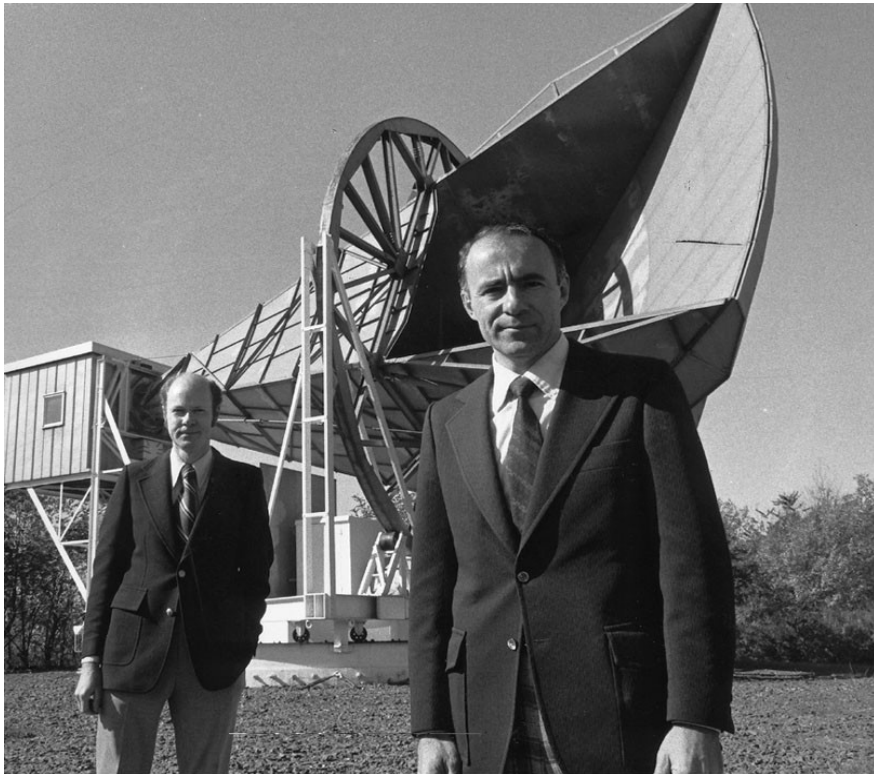
$T \sim 2.73 \text{ K}$ (BB spectrum up to 10^{-5})

Predicted by the « hot Big Bang » model

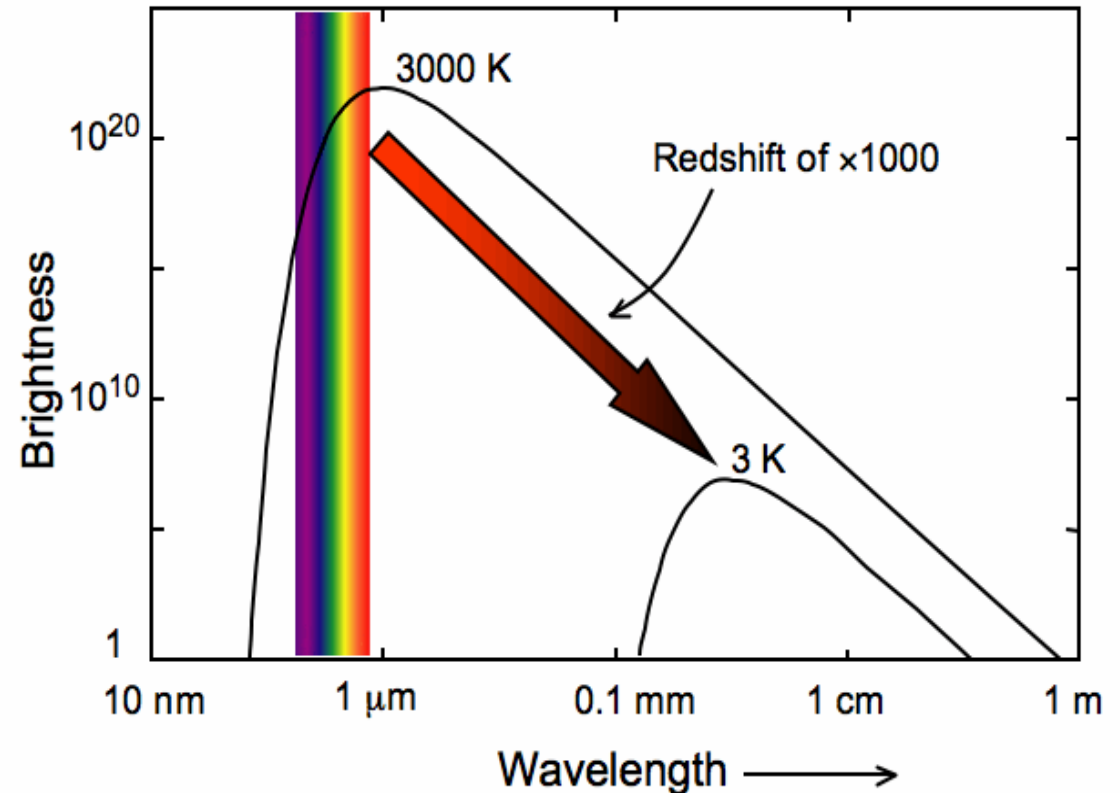
Hot Universe in the Past

Black body spectrum emitted at the time of last ionizations (decoupling of photons and matter)

$T \sim 2.73 \text{ K}$ implies a redshift of ~ 1100



© 2004 Thomson - Brooks/Cole

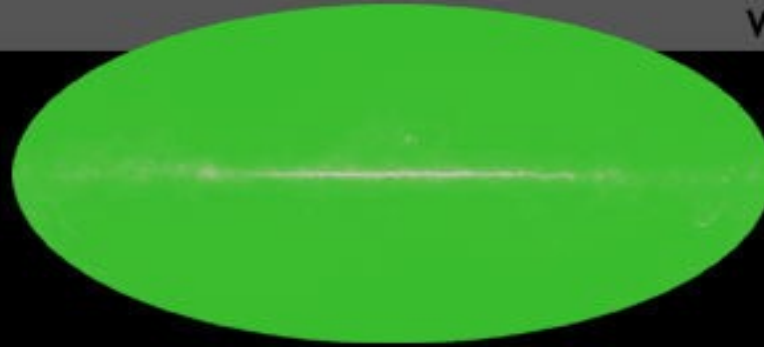


Cosmic Microwave Background (CMB)

1965



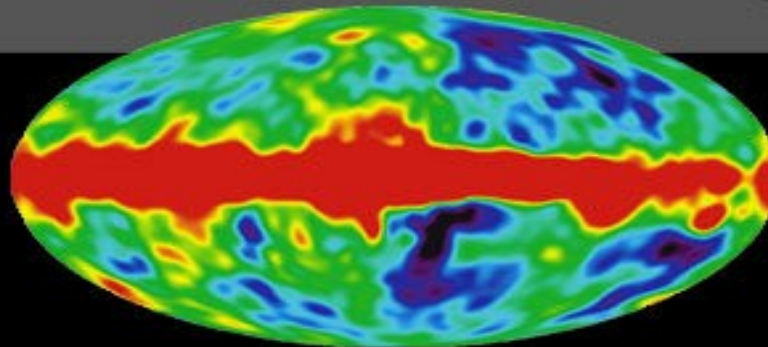
Penzias and
Wilson



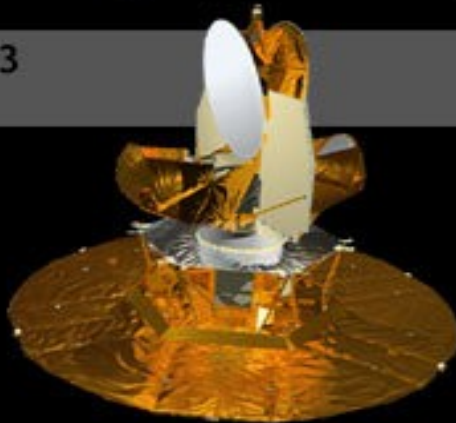
1992



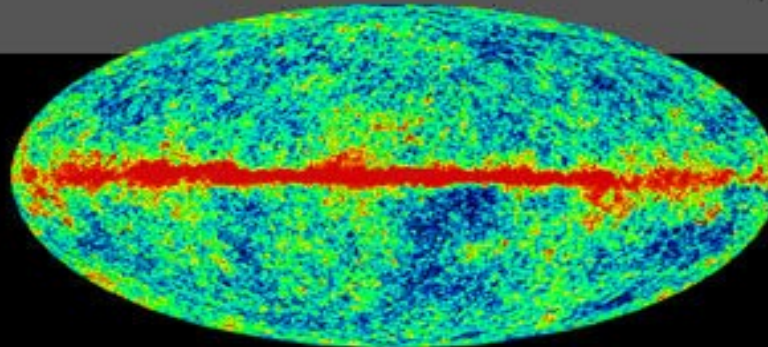
COBE



2003



WMAP

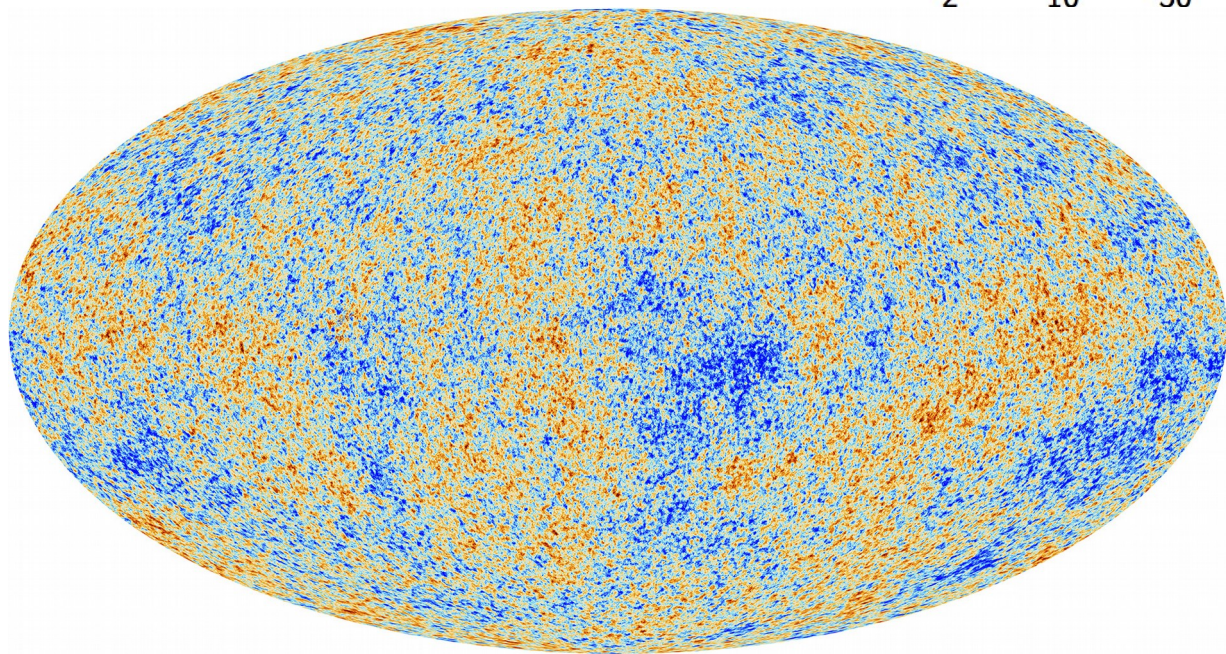
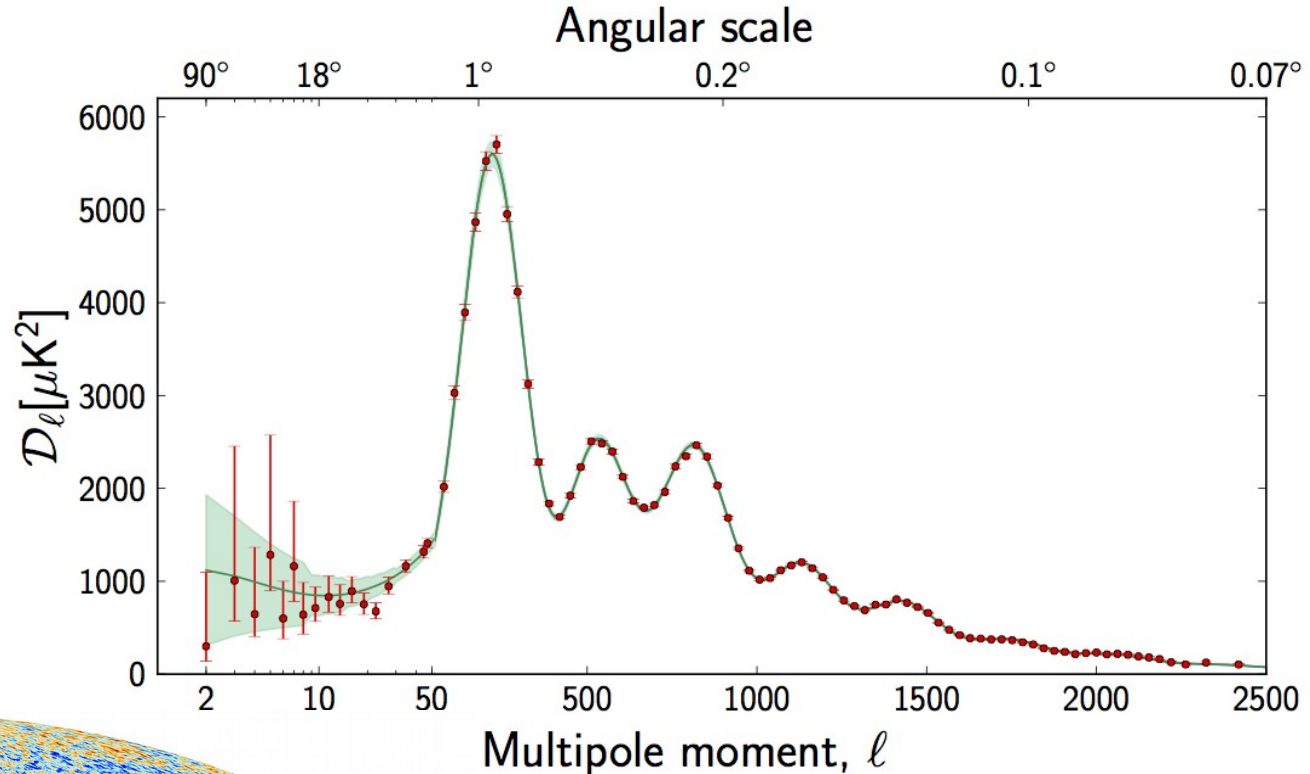


Cosmic Microwave Background (CMB)

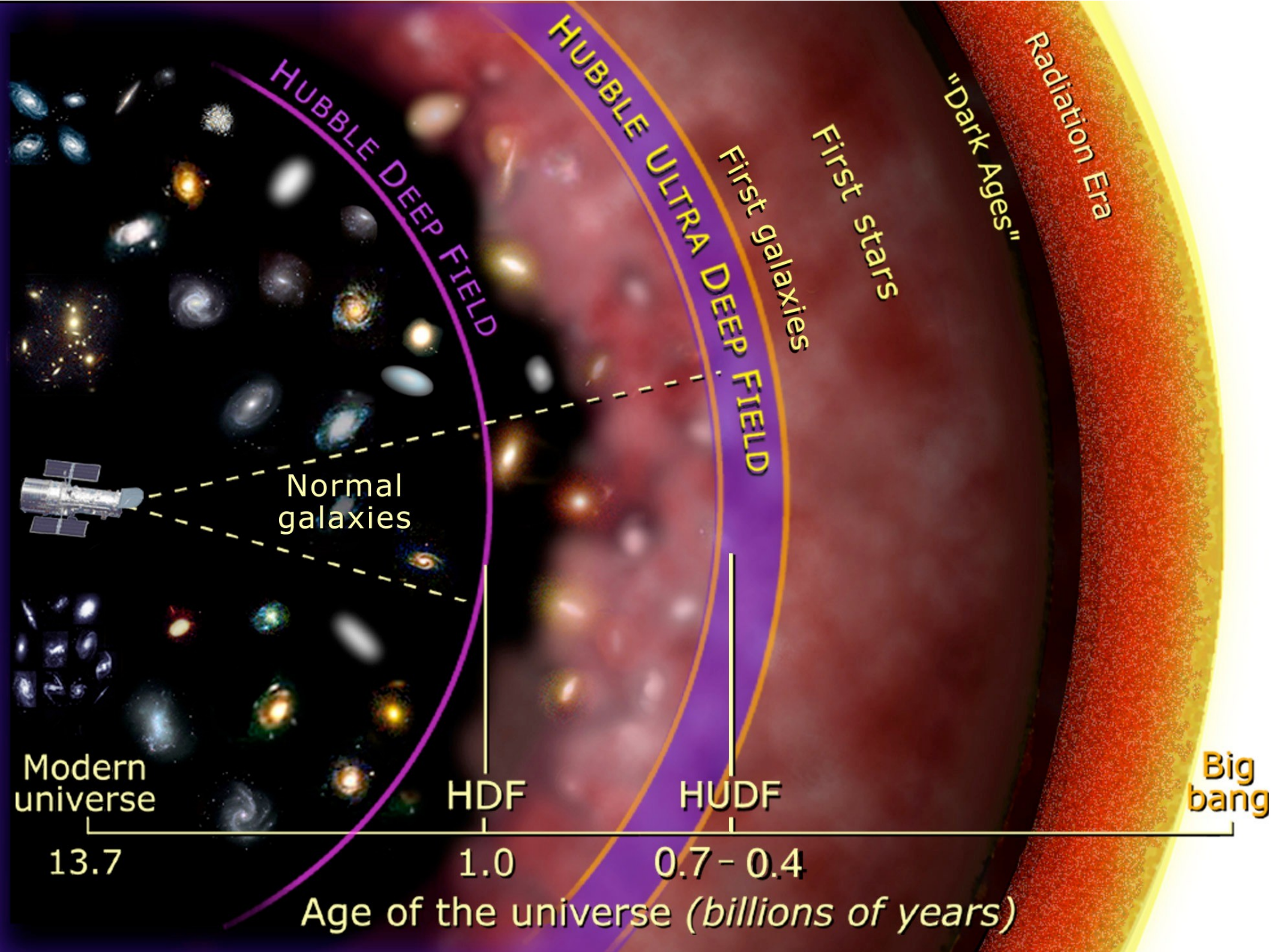
Anisotropies reveal
densities fluctuations
at the time of
emission of the CMB

$z \sim 1100$

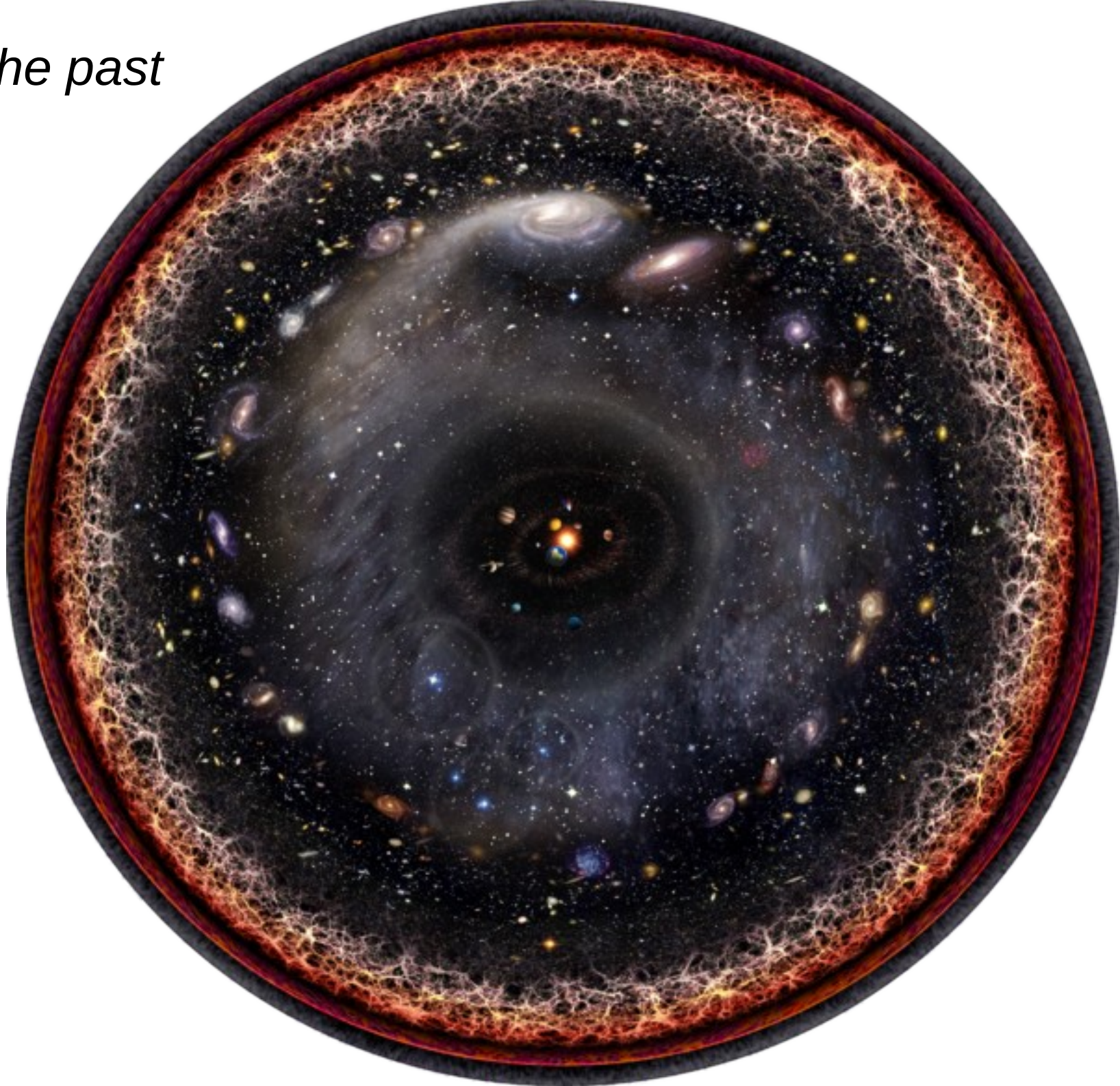
Age $\sim 370\,000$ y



Planck
2009-2012



Looking into the past



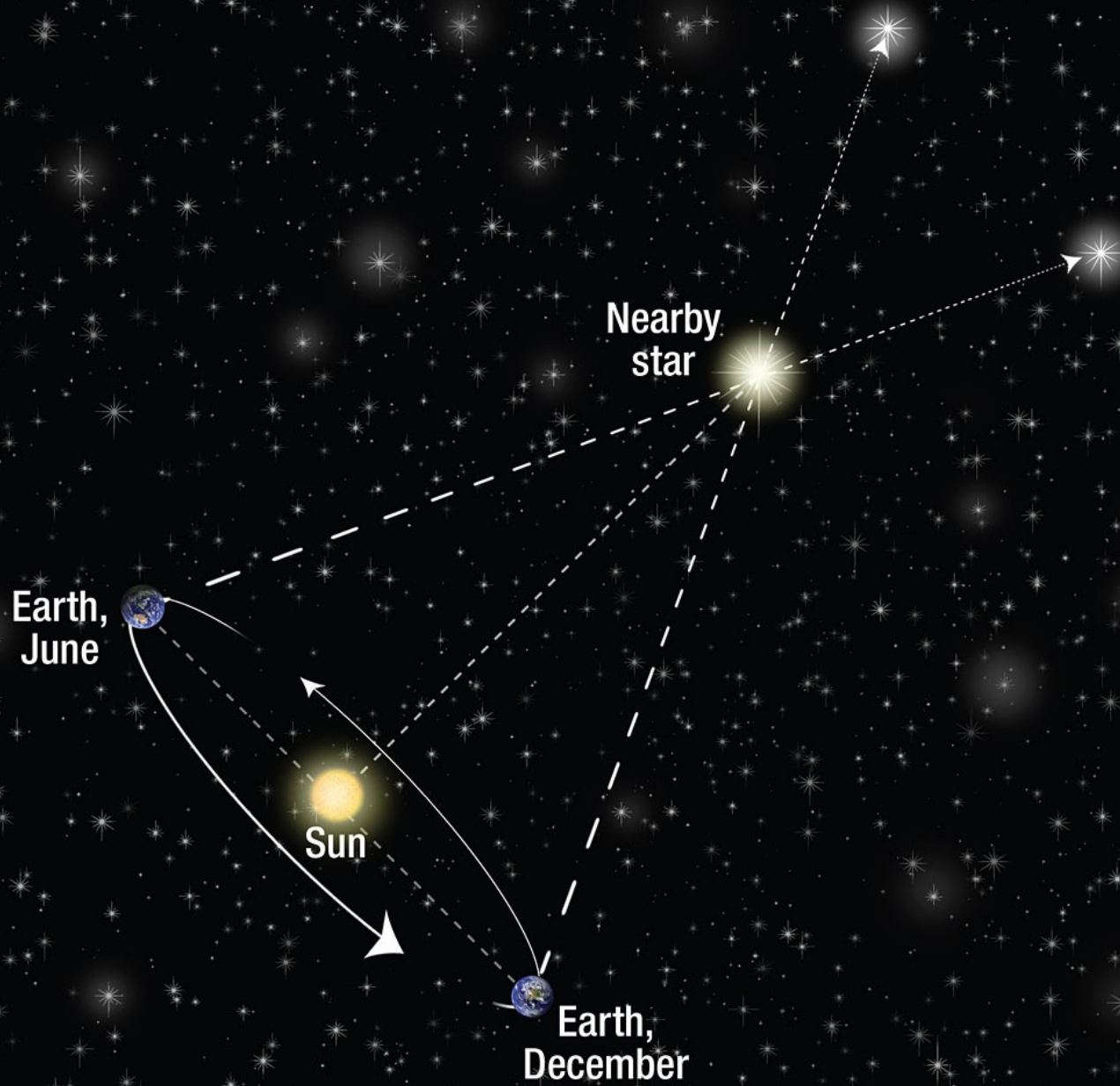
1.2

How do we measure distances in such a big Universe ?

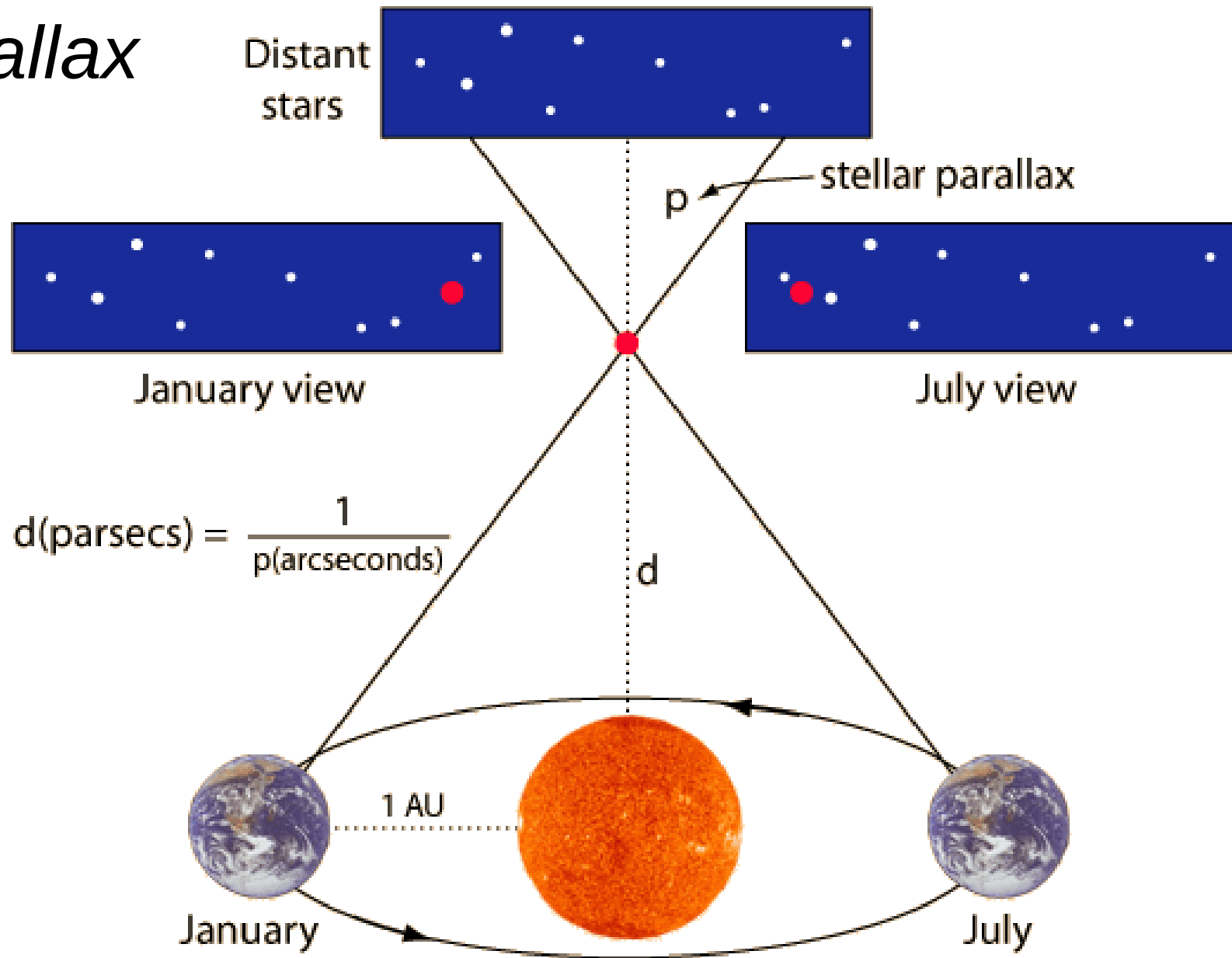
*To do Cosmology we often
need to measure distances to
objects sitting very far away...*

The « Cosmic ladder »

Stellar parallax



Stellar parallax



$$d(\text{parsecs}) = \frac{1}{p(\text{arcseconds})}$$

$$1 \text{ pc} = 3.26 \text{ ly} = 3.09 \times 10^{16} \text{ m}$$

Stellar parallax

- Only nearby stars (19-20th century)

~ 1000 stars, up to ~100 pc

- **Hipparcos** (1989-1993)

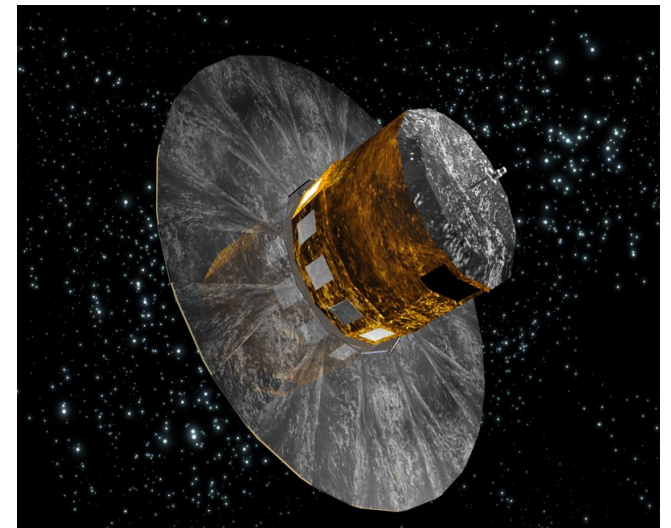
~ 100 000 stars, 0.001 arcsec,
distances up to ~500 pc



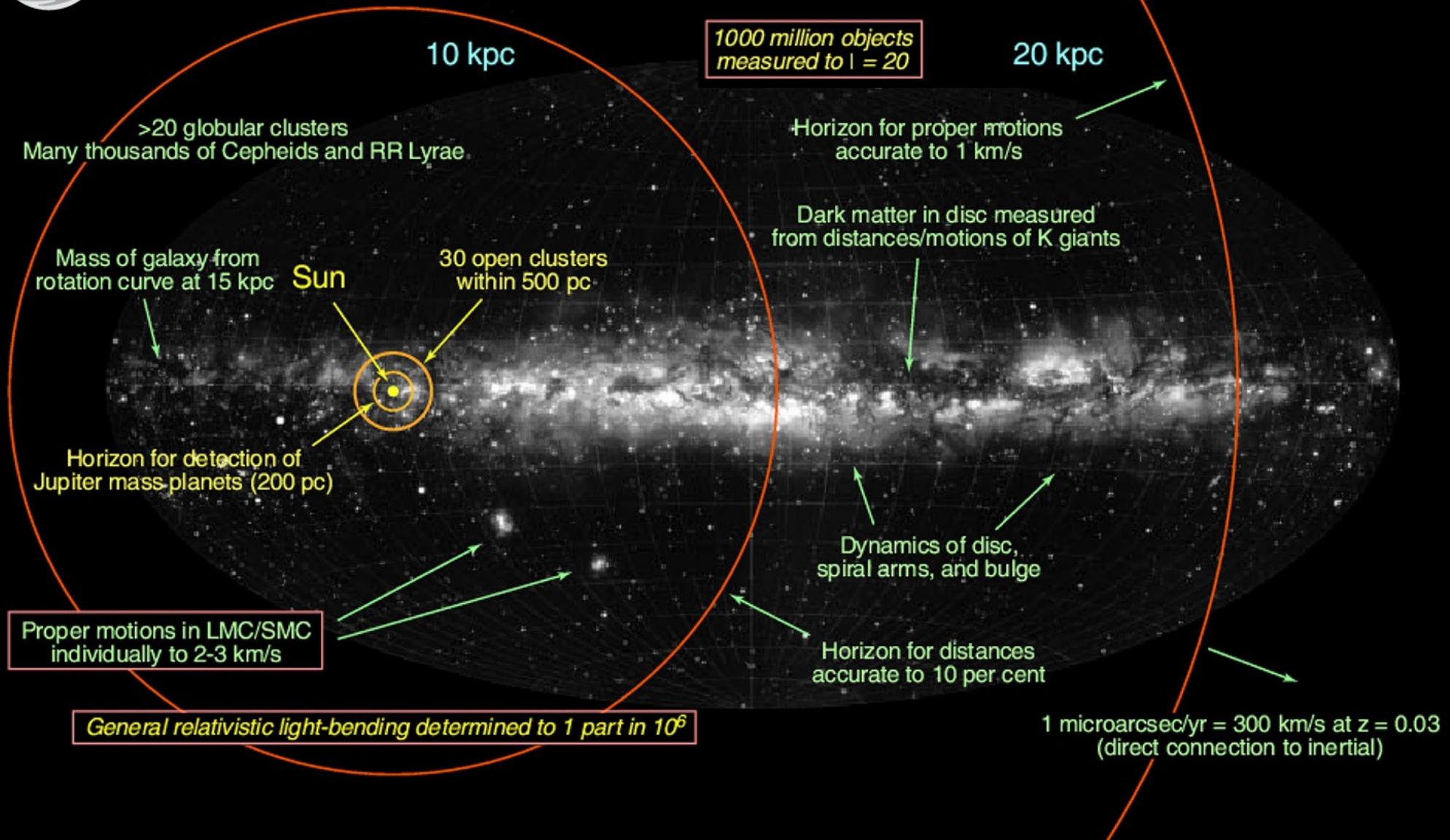
- **GAIA** (2013-)

3D catalog of ~ 10^9 stars

reliable distances up to ~10 kpc



Stellar parallax



Stellar properties

Stellar spectrum : a black body $\rightarrow T^\circ$
+ emission/absorption lines

Spectral classification : OBAFGKM

Temperature, Color and Luminosity

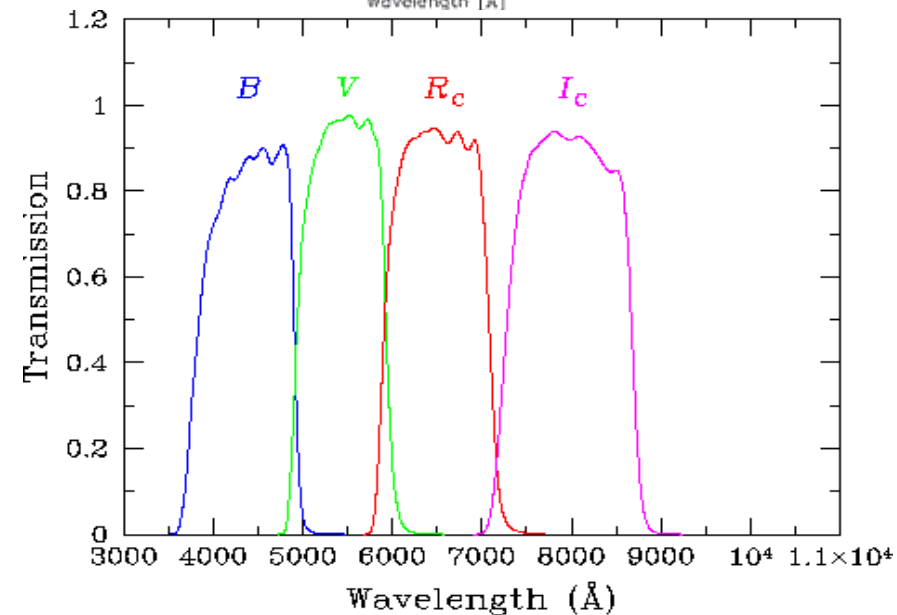
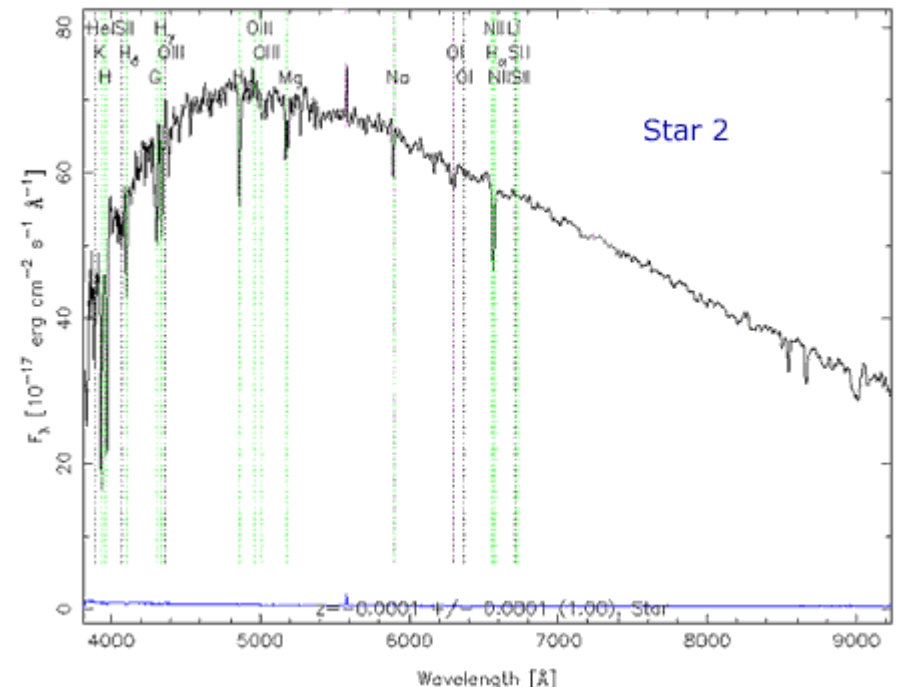
Stellar apparent magnitude (in filters)
(ref = ref star, Vega) :

$$m_x - m_{x,\text{ref}} = -2.5 \log_{10} \left(\frac{I_x}{I_{x,\text{ref}}} \right)$$

Absolute magnitude M : magnitude
of the same object seen at 10 pc

Distance modulus μ :

$$m - M = \mu = 5 \log_{10} \left(\frac{d}{10 \text{ pc}} \right) = 5 \log_{10} d - 5$$



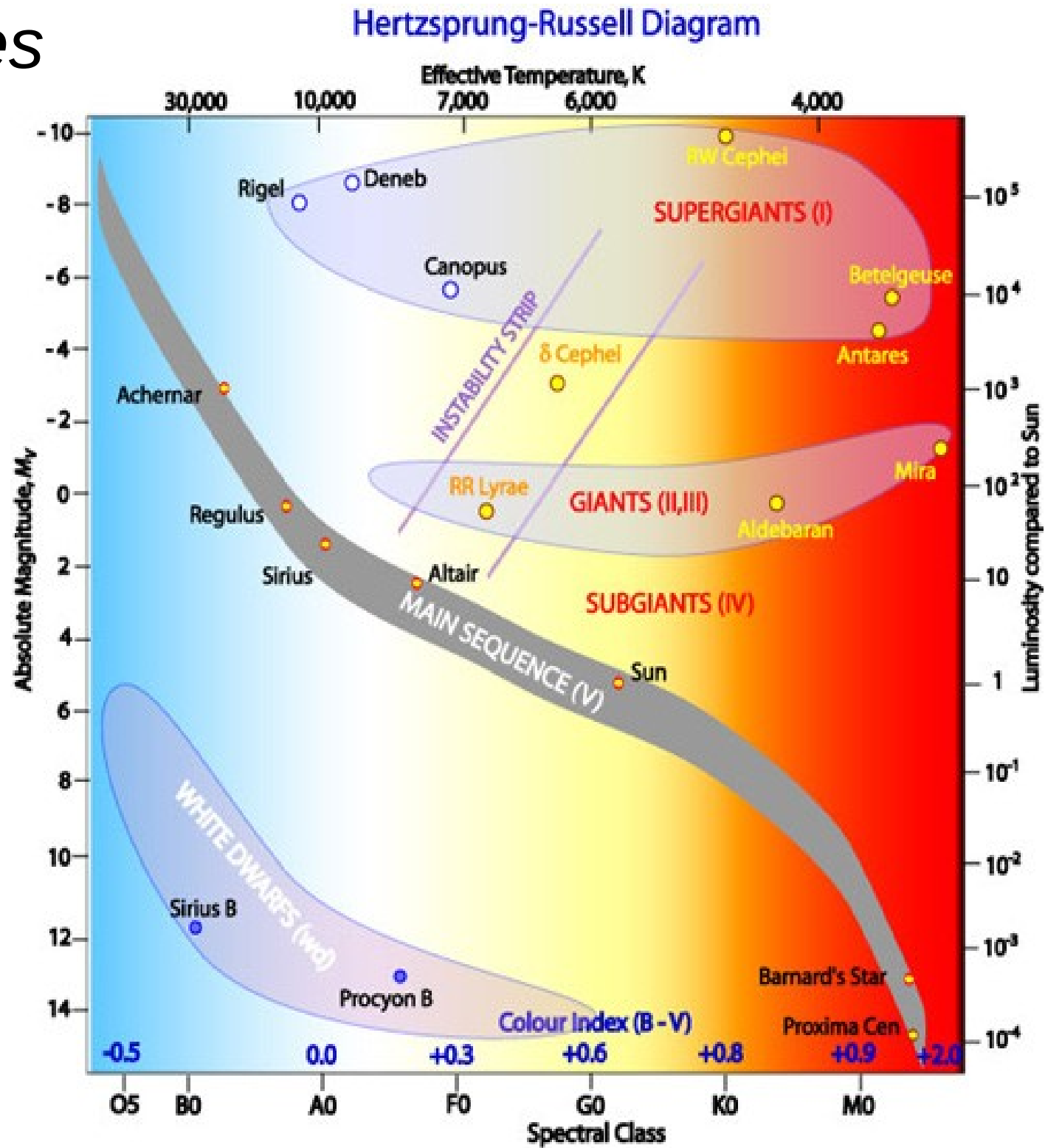
Stellar properties

Absolute Mag.
Temp° & color

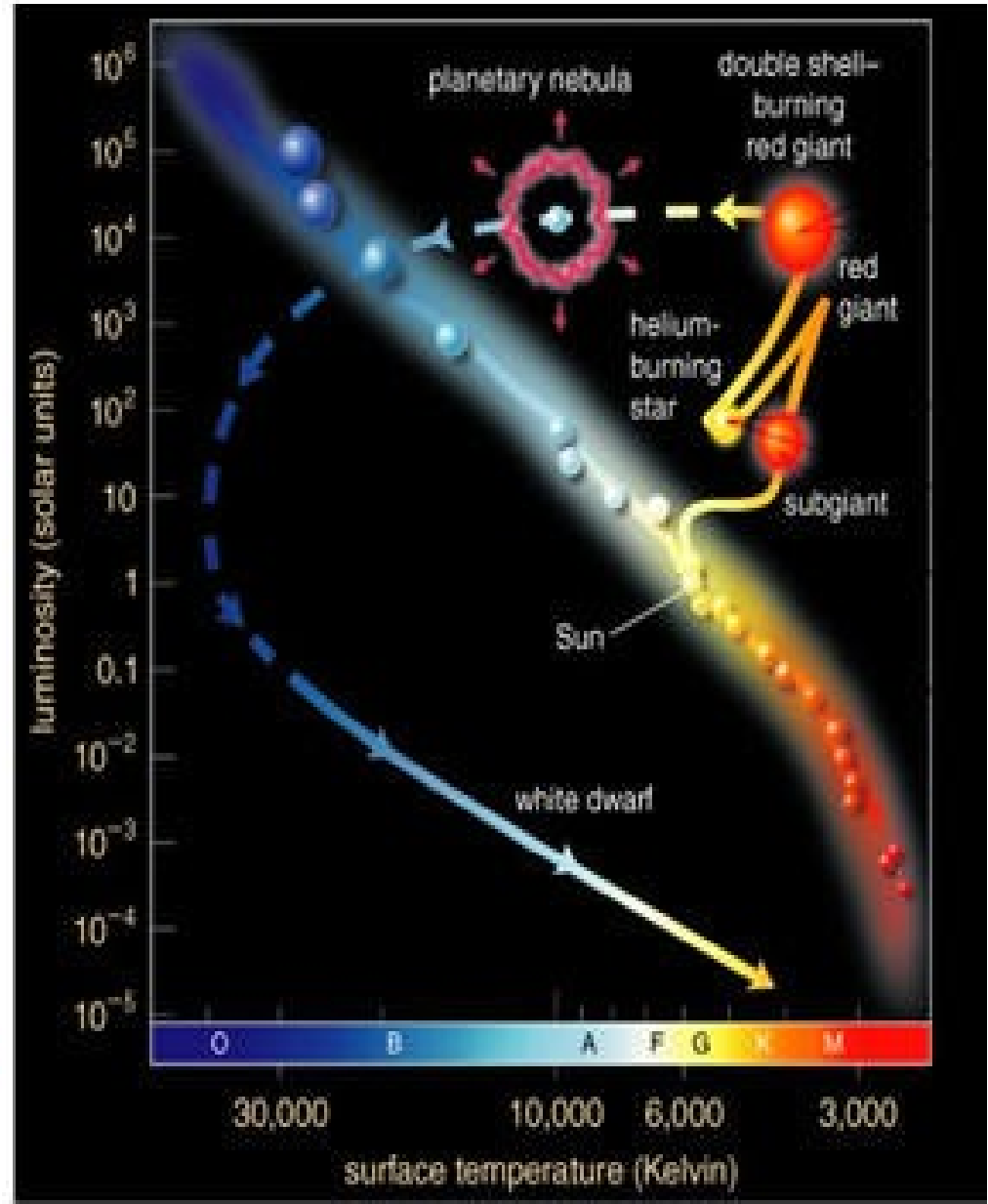
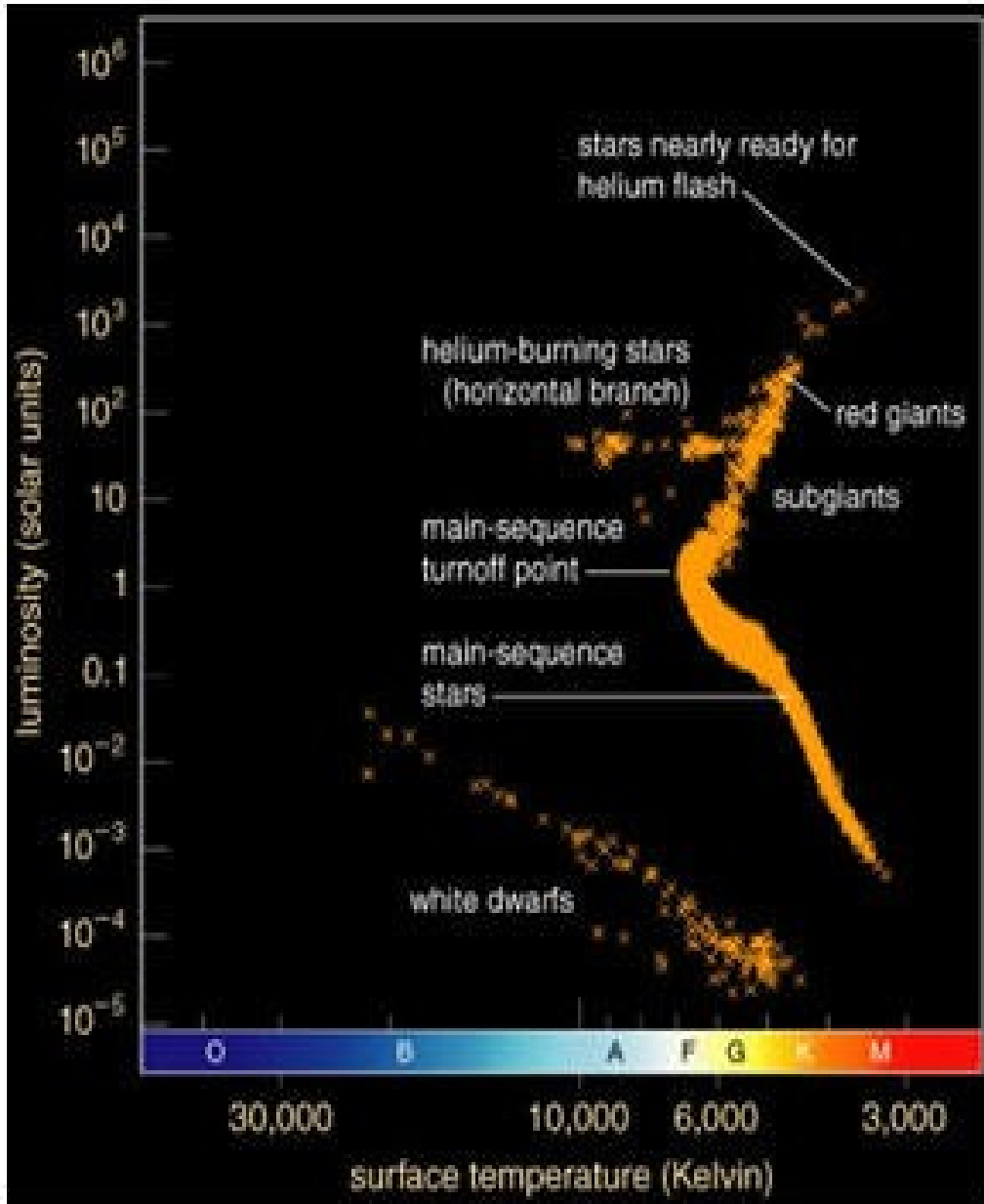
Stellar classes

Stellar Evolution

H-R diagram can
be build for stars
at the same
distance (clusters)

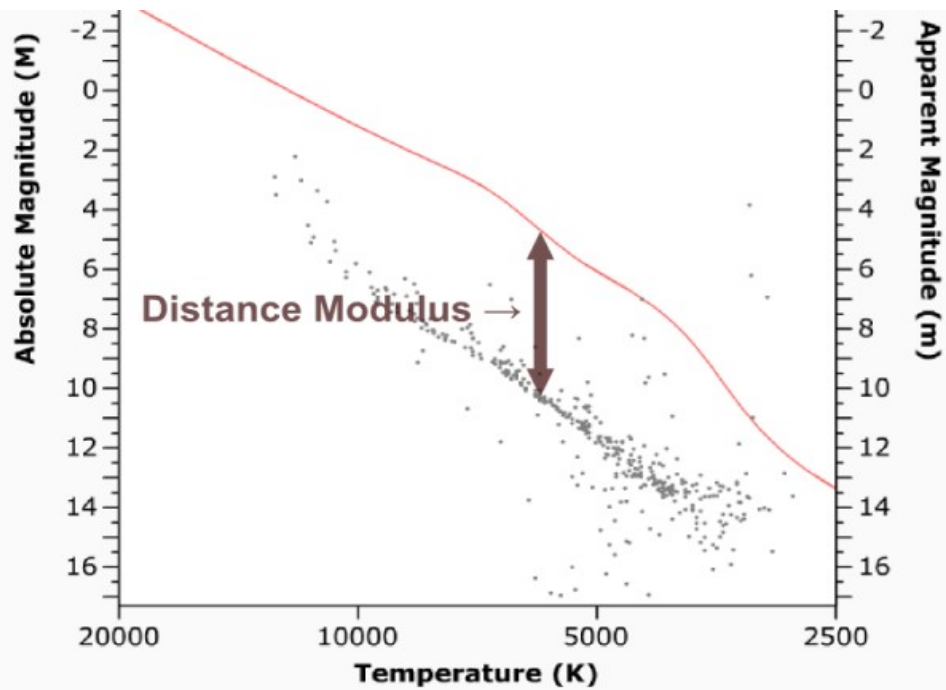
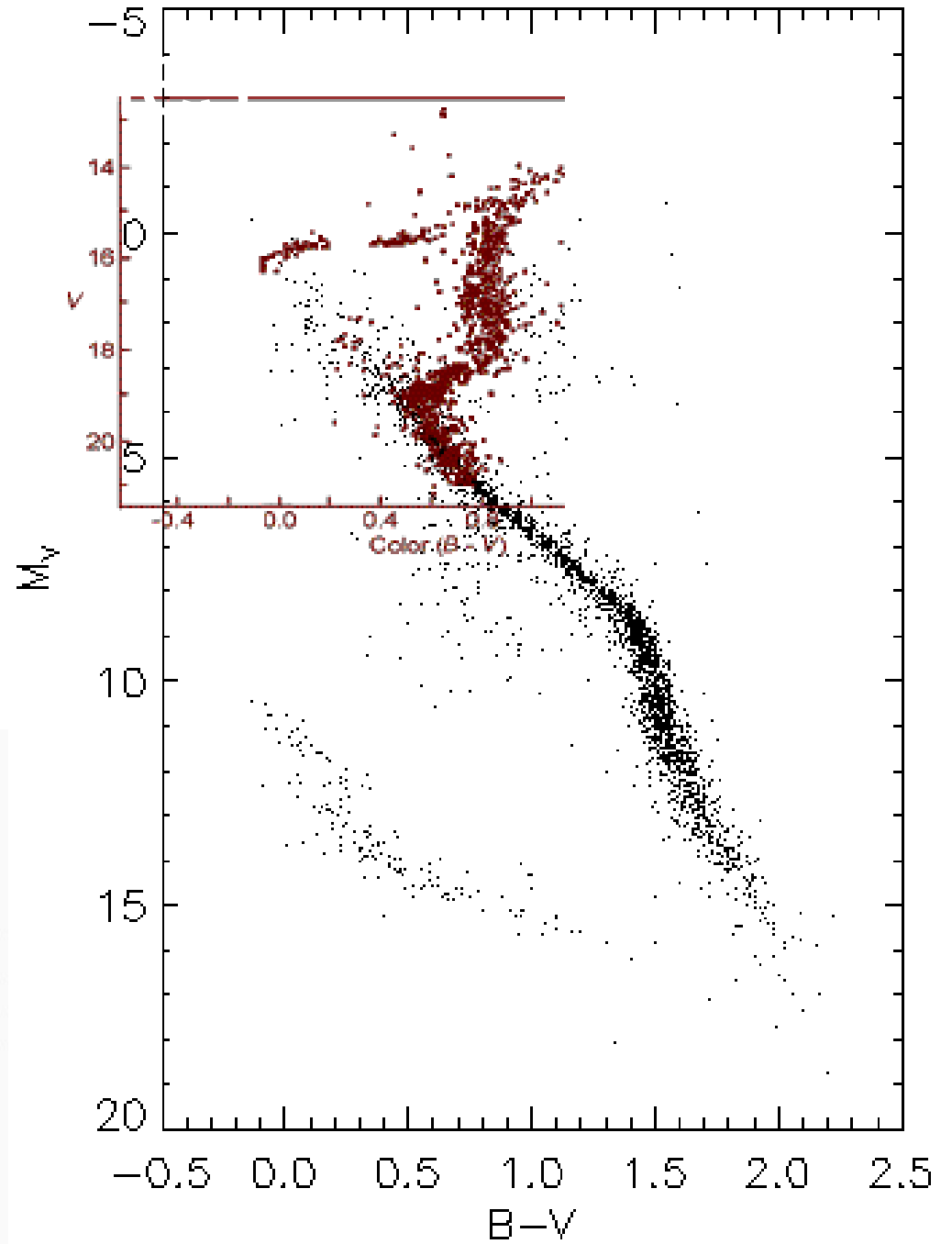
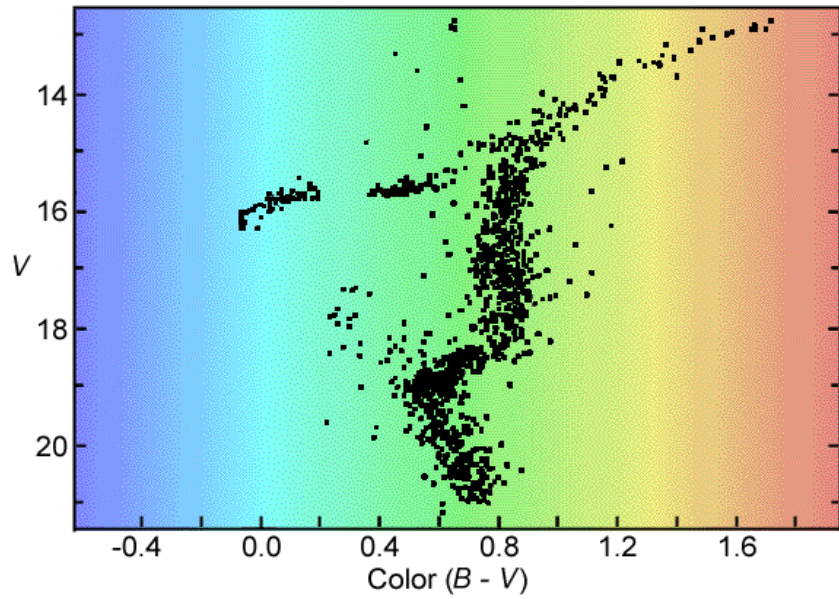


Stellar evolution



Main sequence fitting

Stars from M3 (cluster)

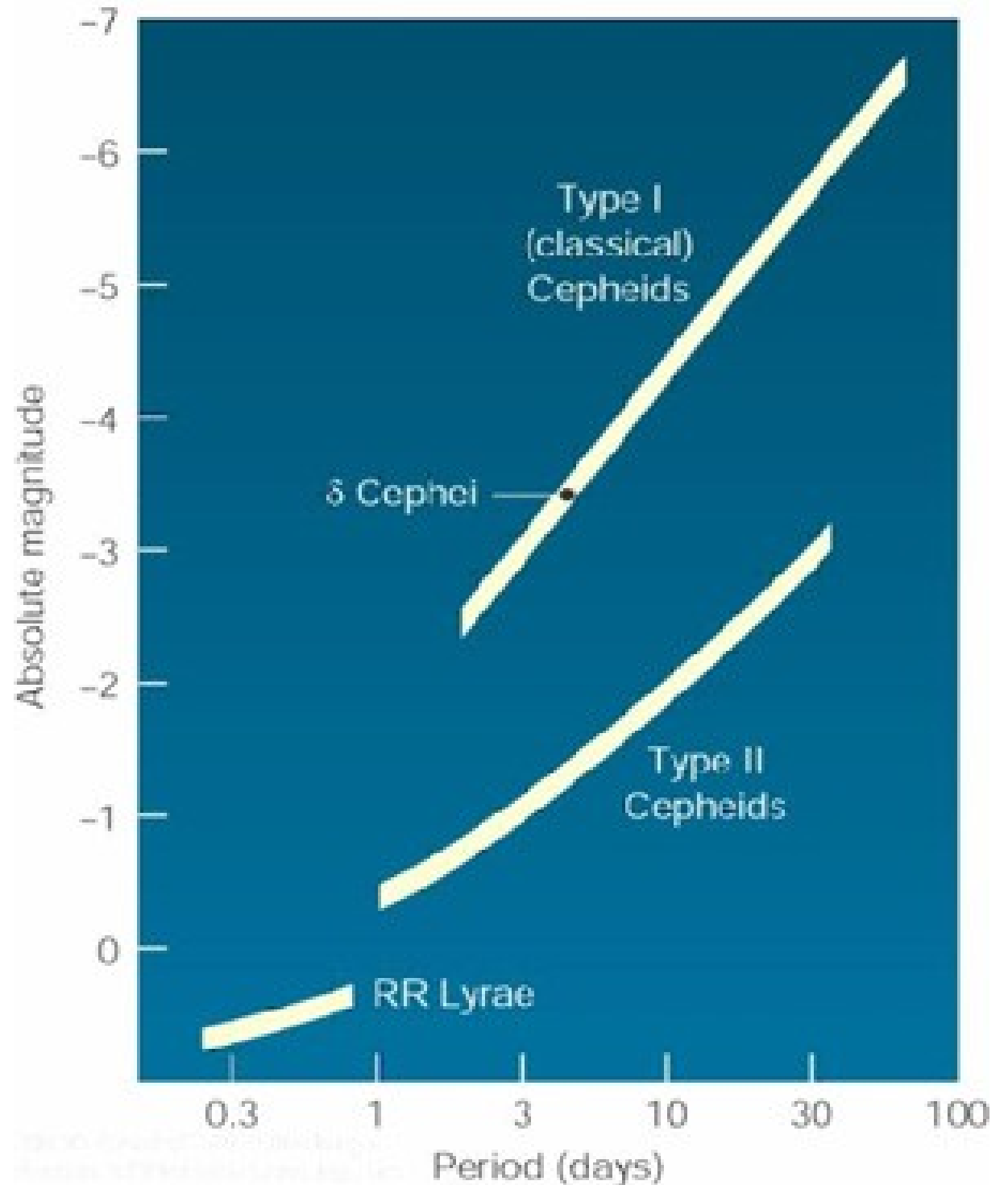


Variable stars

« Standard candles »

Objects with a known absolute luminosity

Distance modulus and distance may be deduced



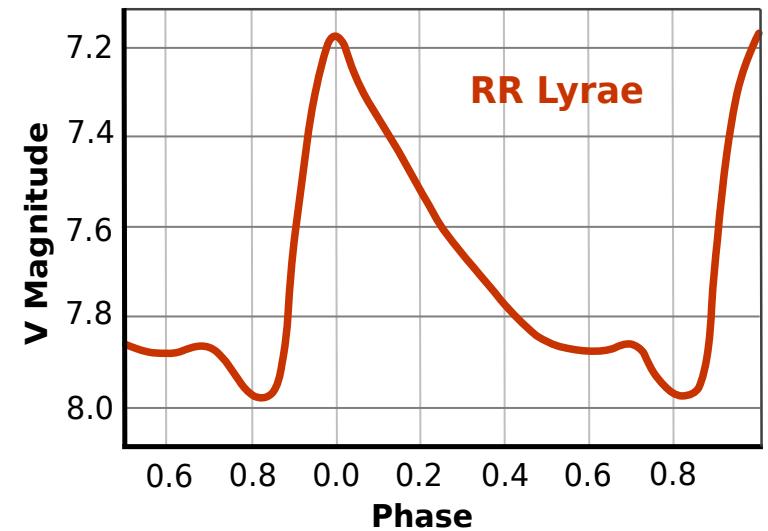
RR Lyrae

Pulsating variables

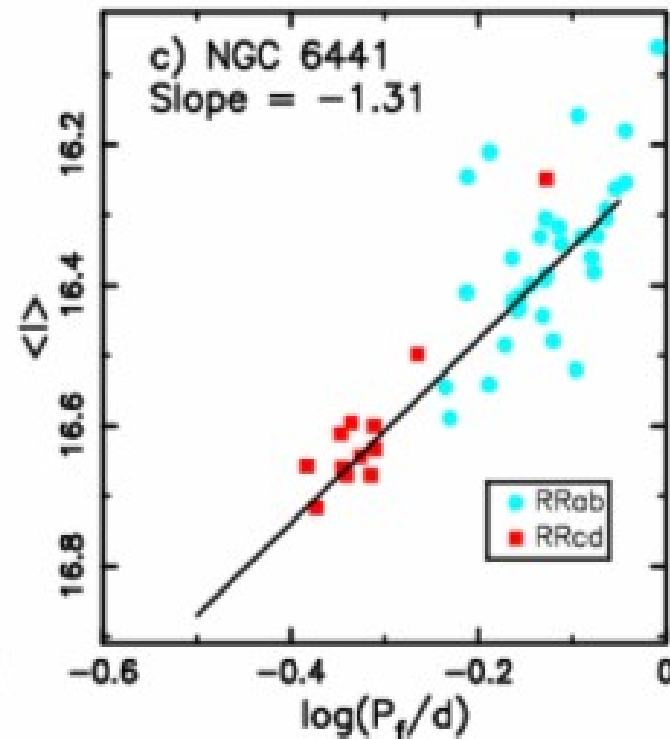
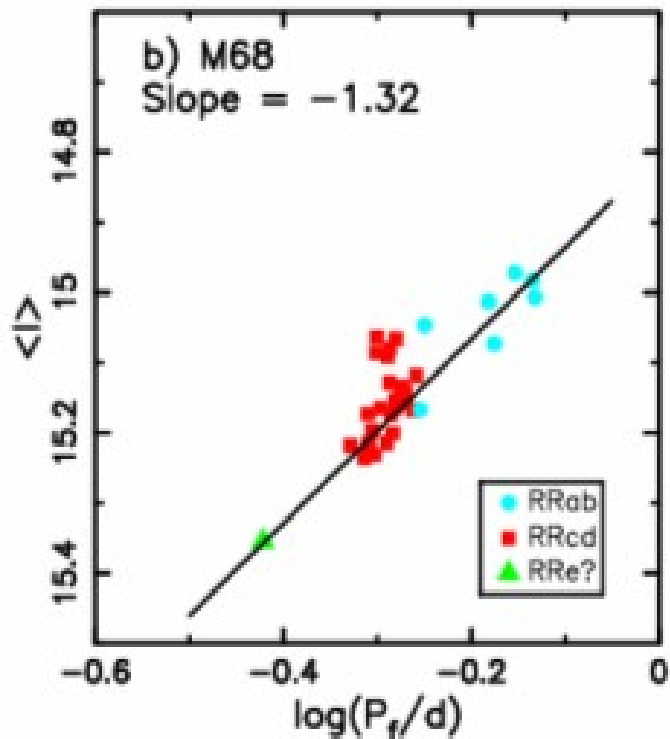
$P \sim 0.1 - 1$ day

Period-luminosity relation

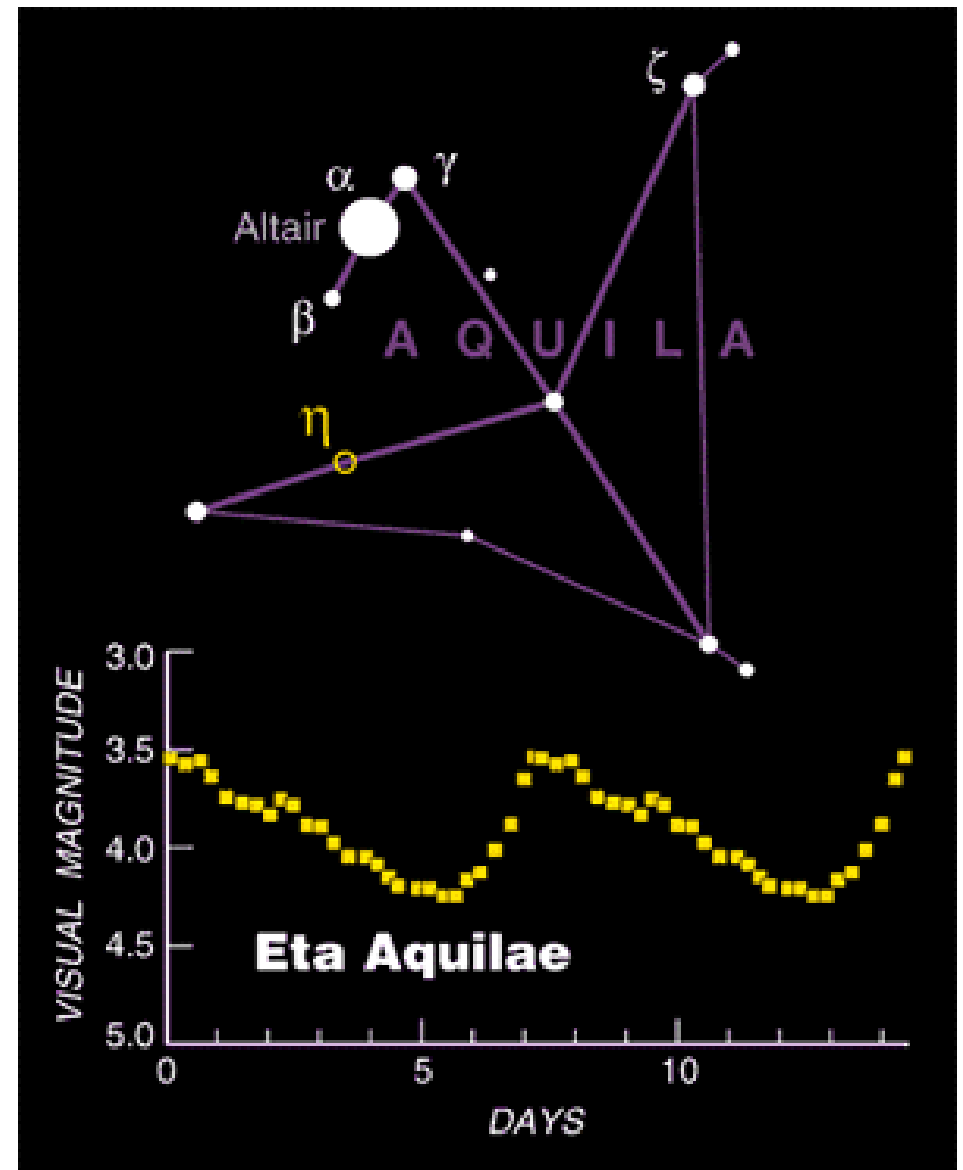
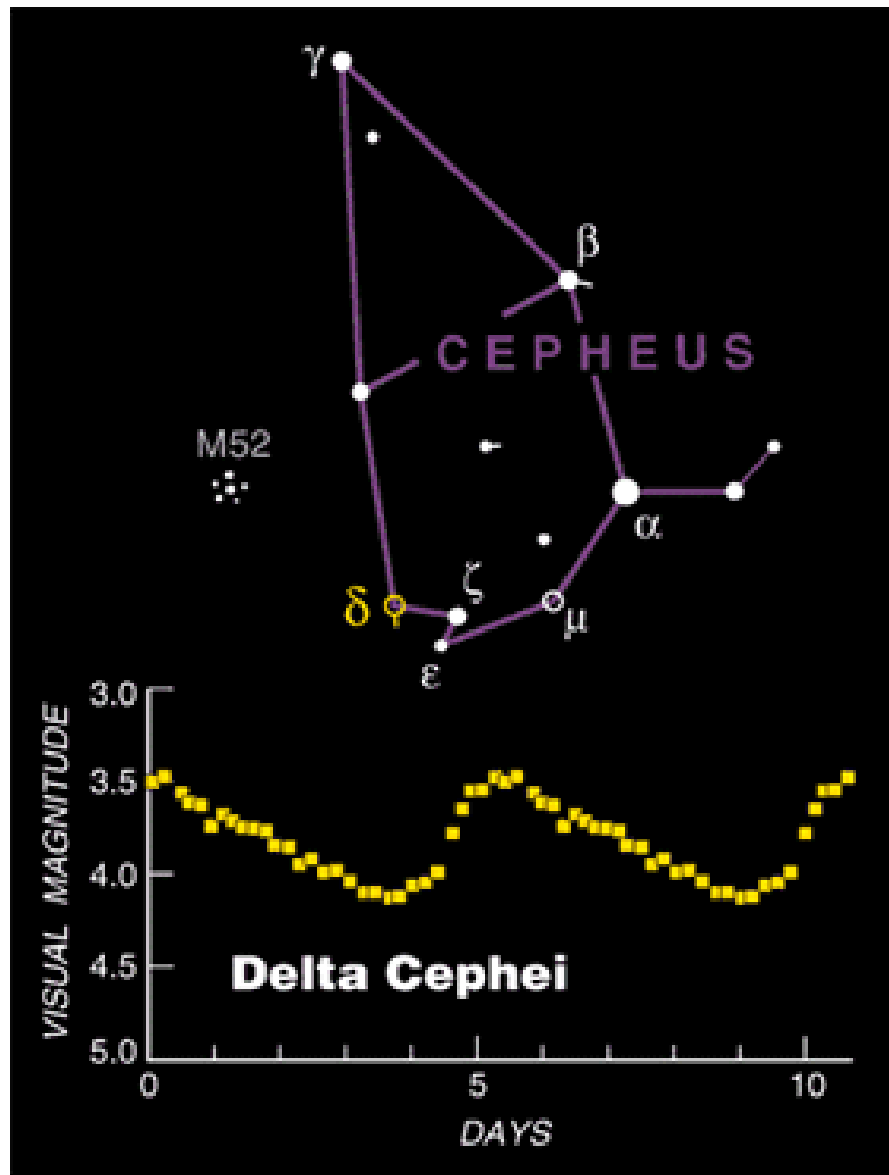
Abs Mag $\sim +0.75$



Distance indicator inside the galaxy / local group



Cepheids : pulsating stars



Cepheids : luminosity-distance relation

Pulsating variables

$P \sim 1 - 50$ days

Very bright stars

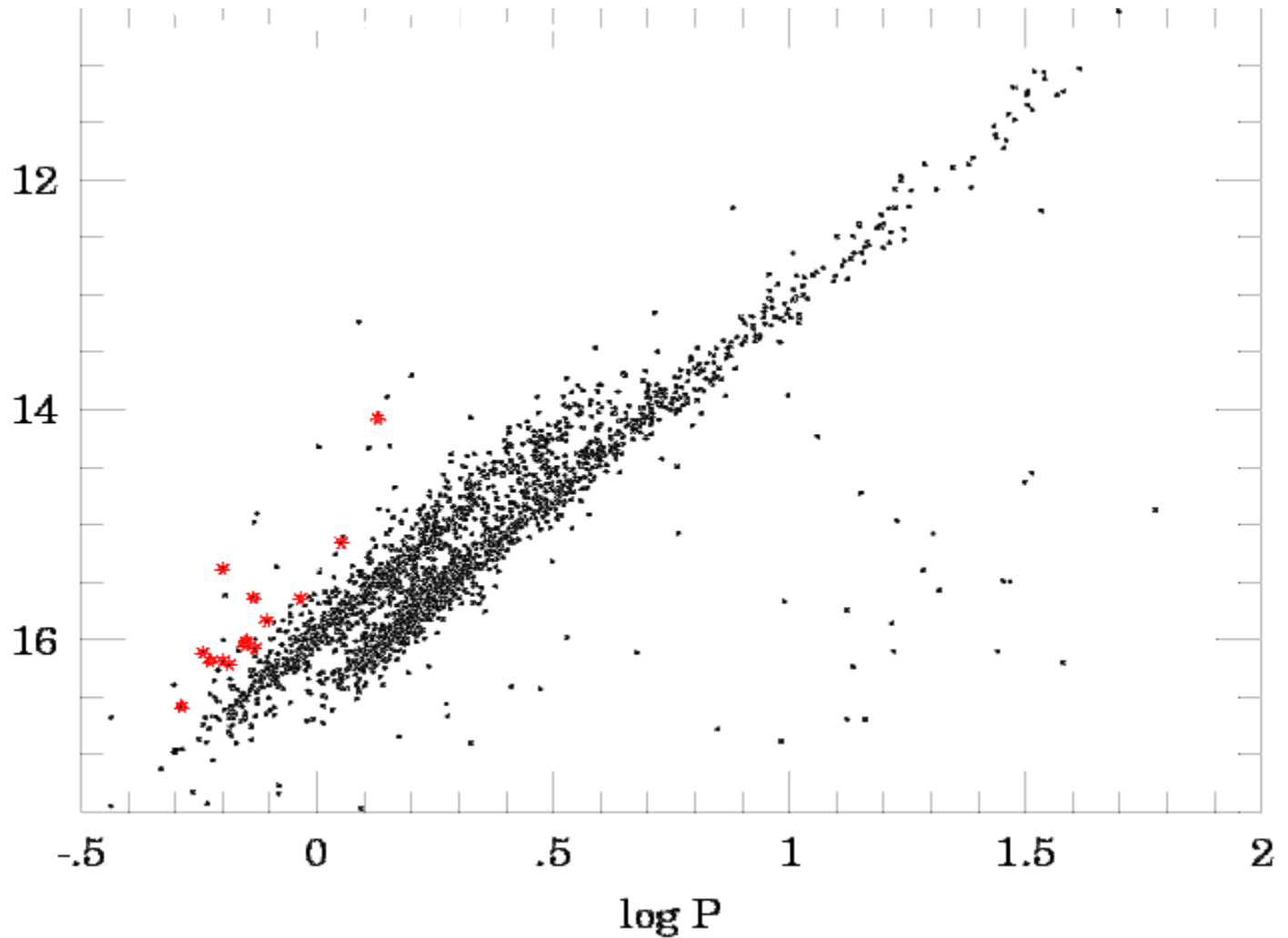
Rare objects

2 main populations

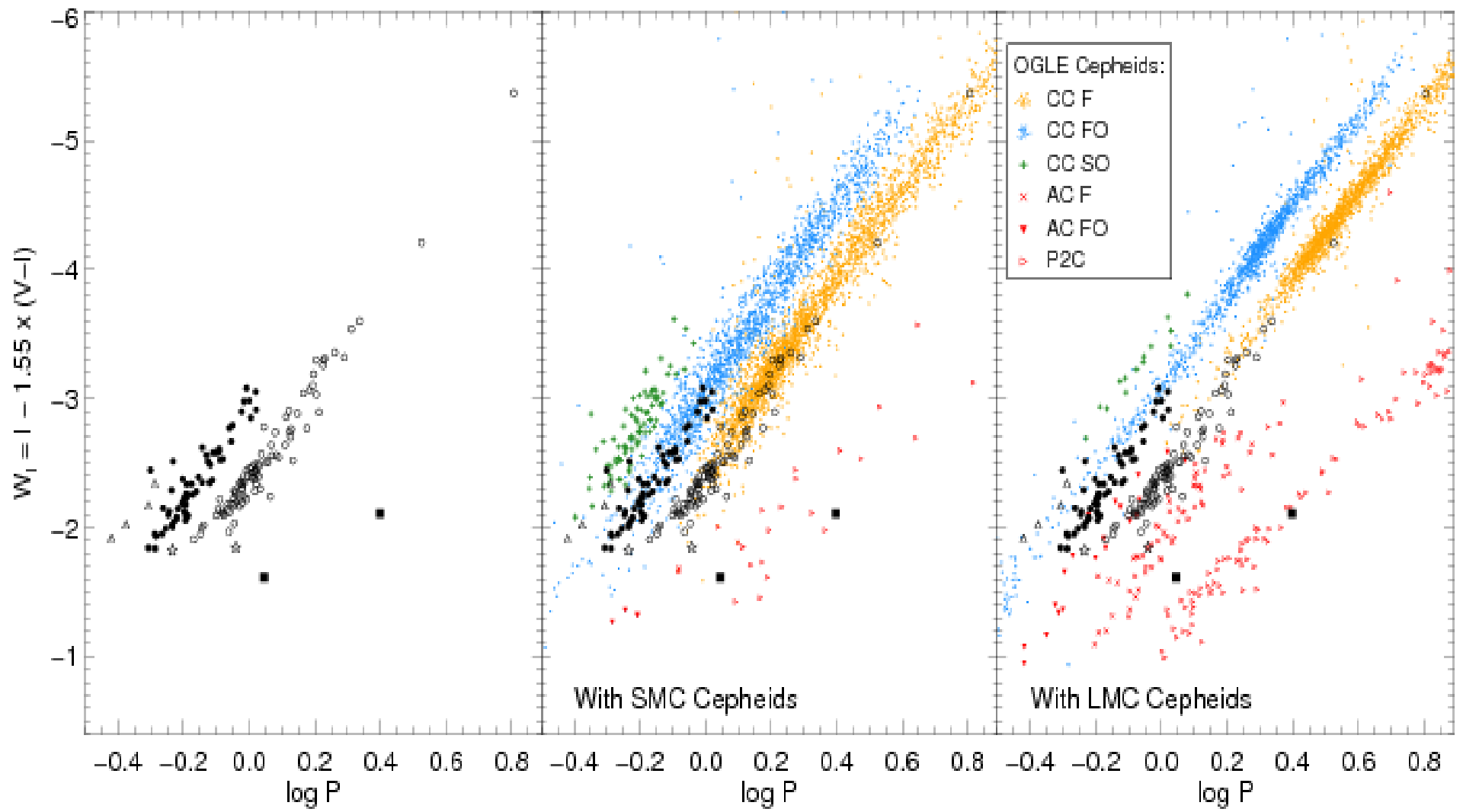
- Classical Cepheids
- Type II (overtone)

Distance indicator

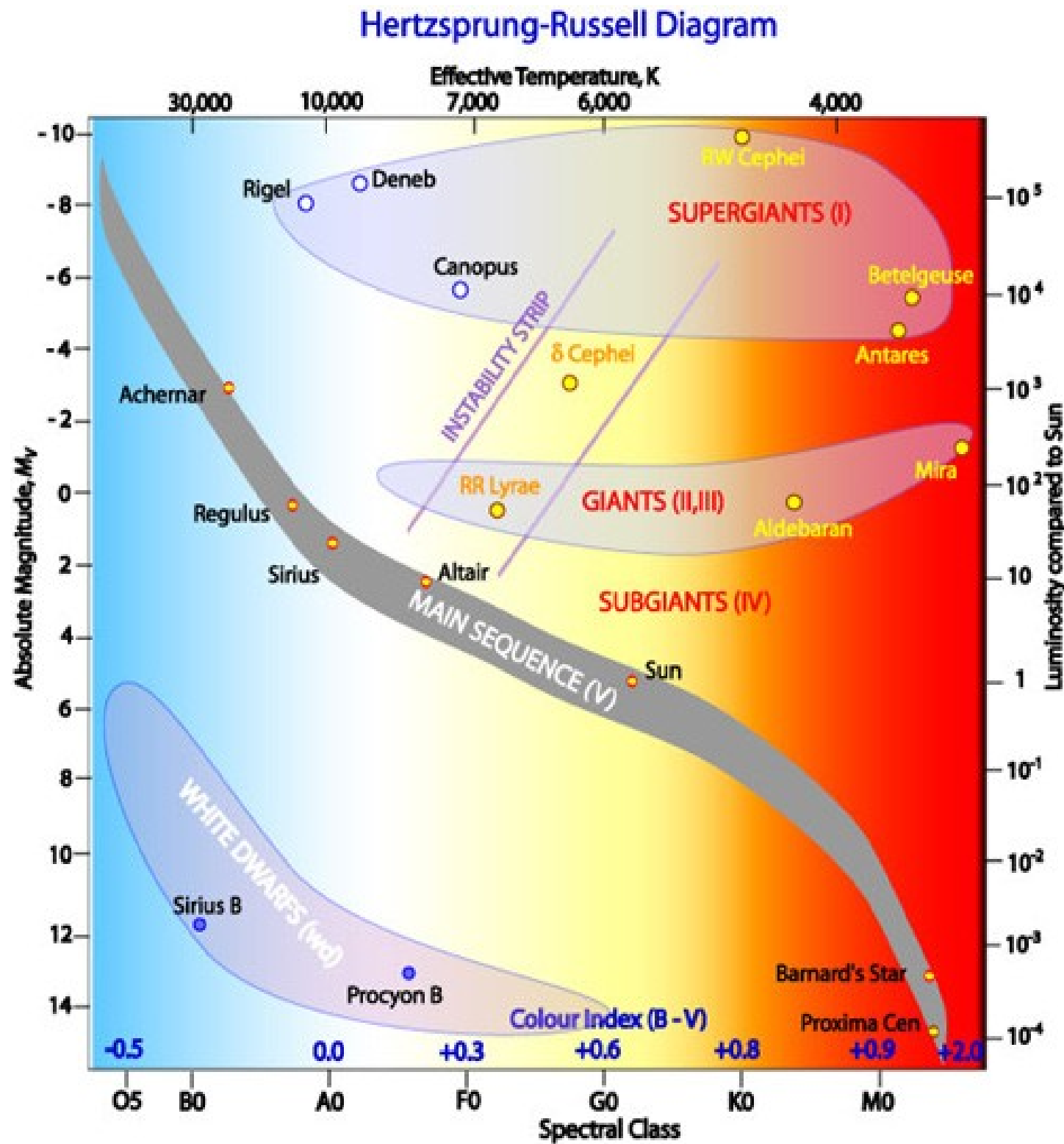
Up to ~ 30 Mpc



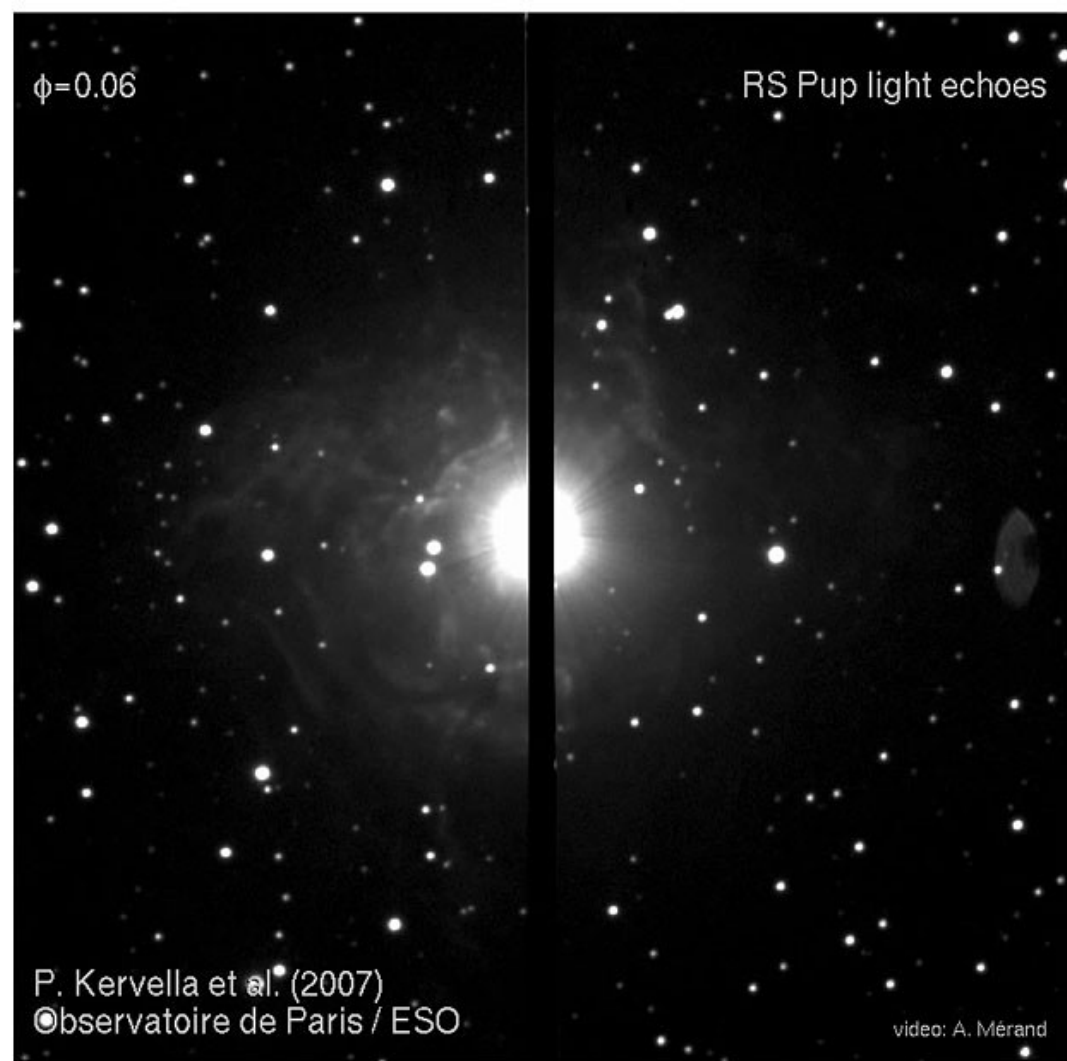
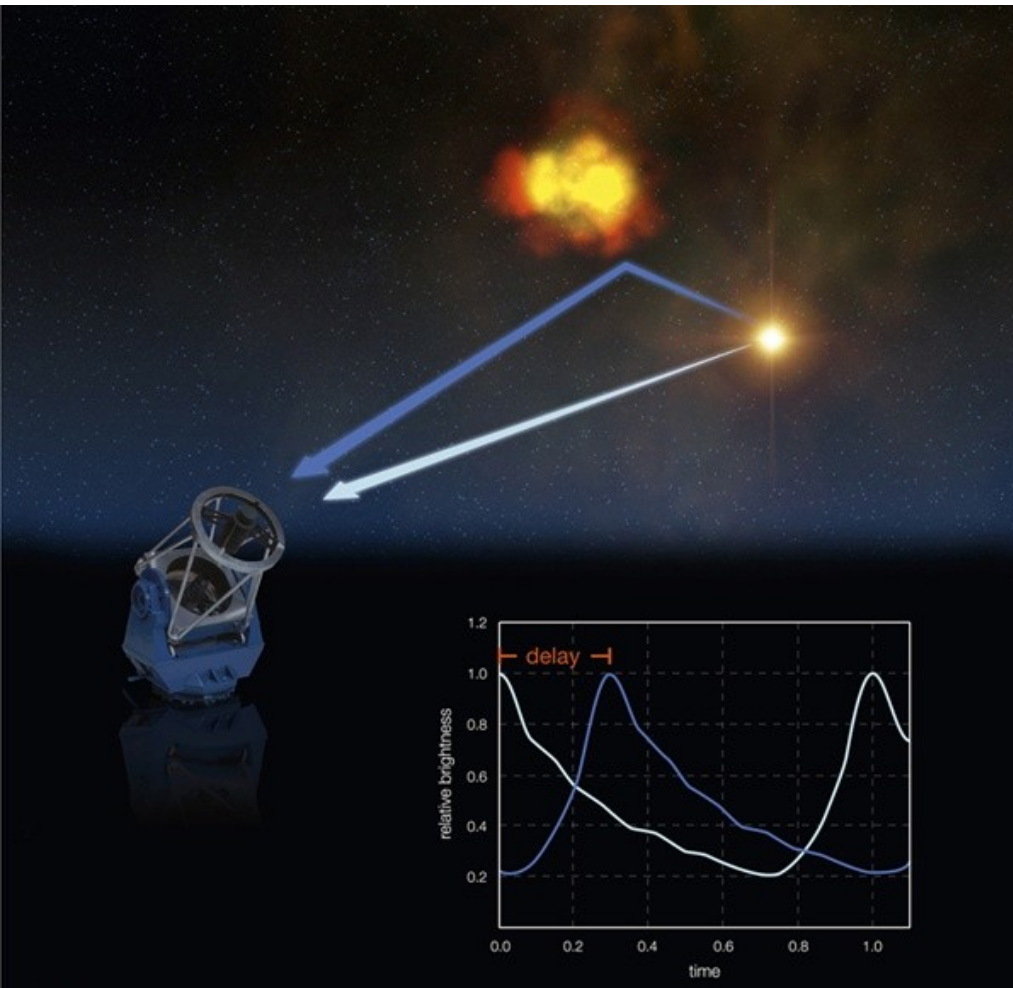
$$M_v = (-2.43 \pm 0.12)(\log_{10}(P) - 1) - (4.05 \pm 0.02)$$



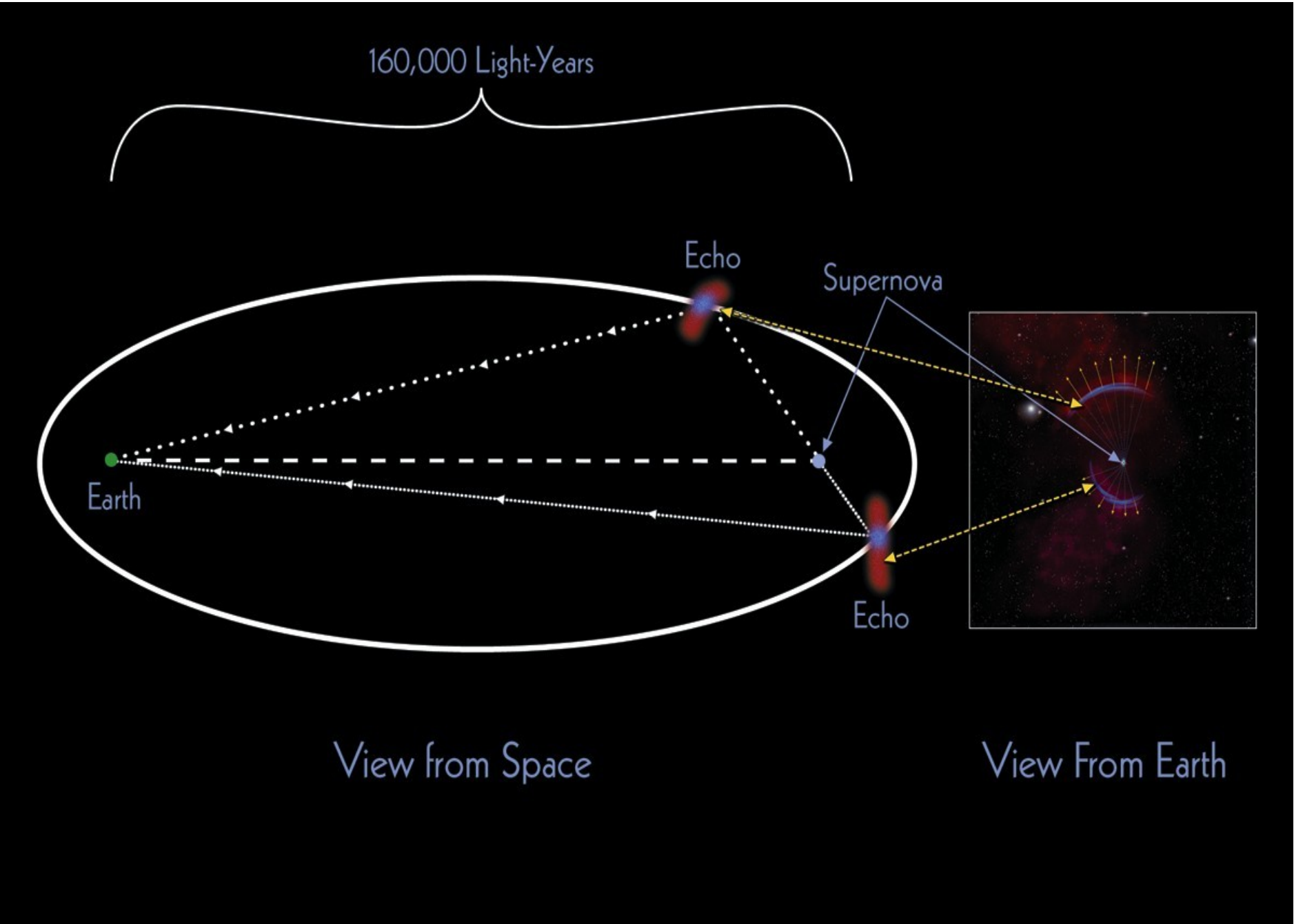
Cepheids



Light echoes



Light echoes



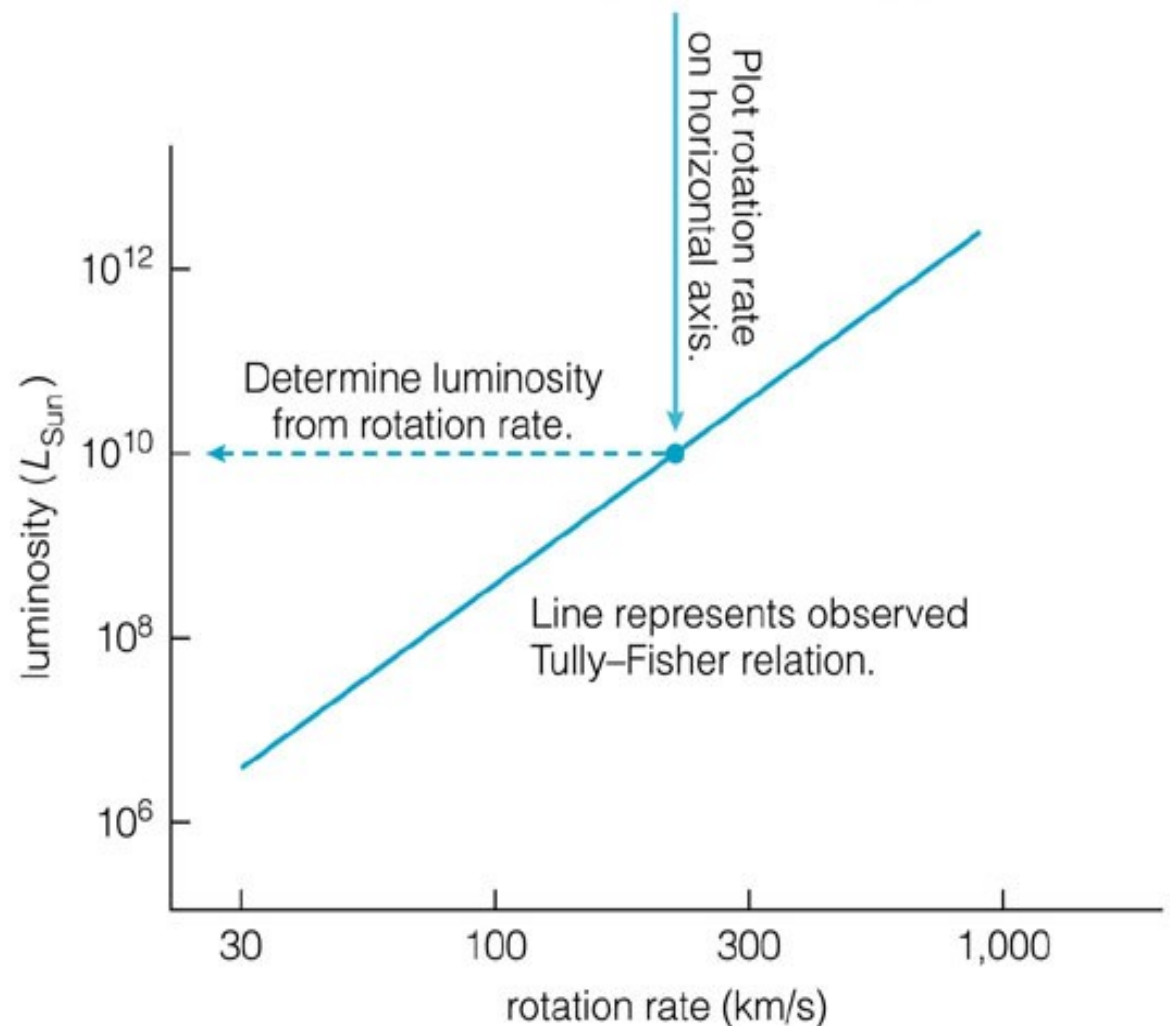
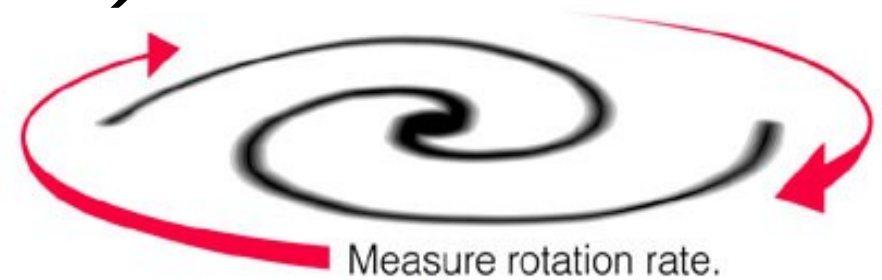
Tully-Fisher (spiral galaxies)

R. Tully & J. Fisher (1977)

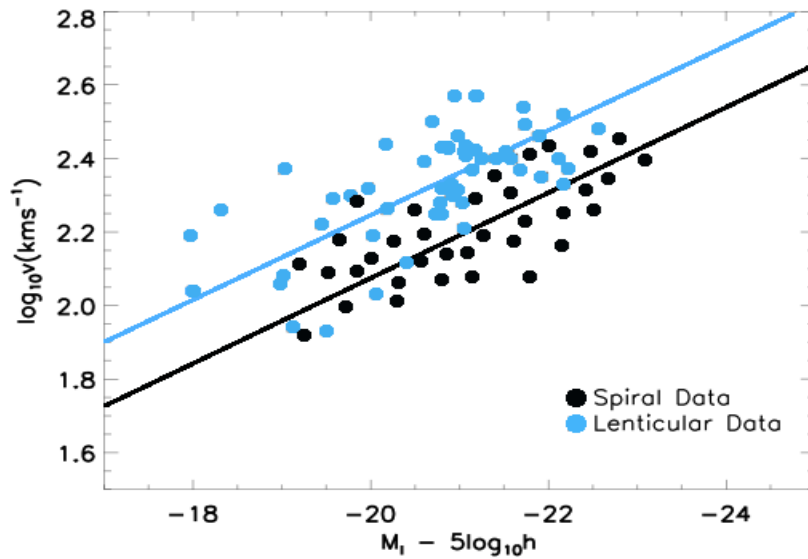
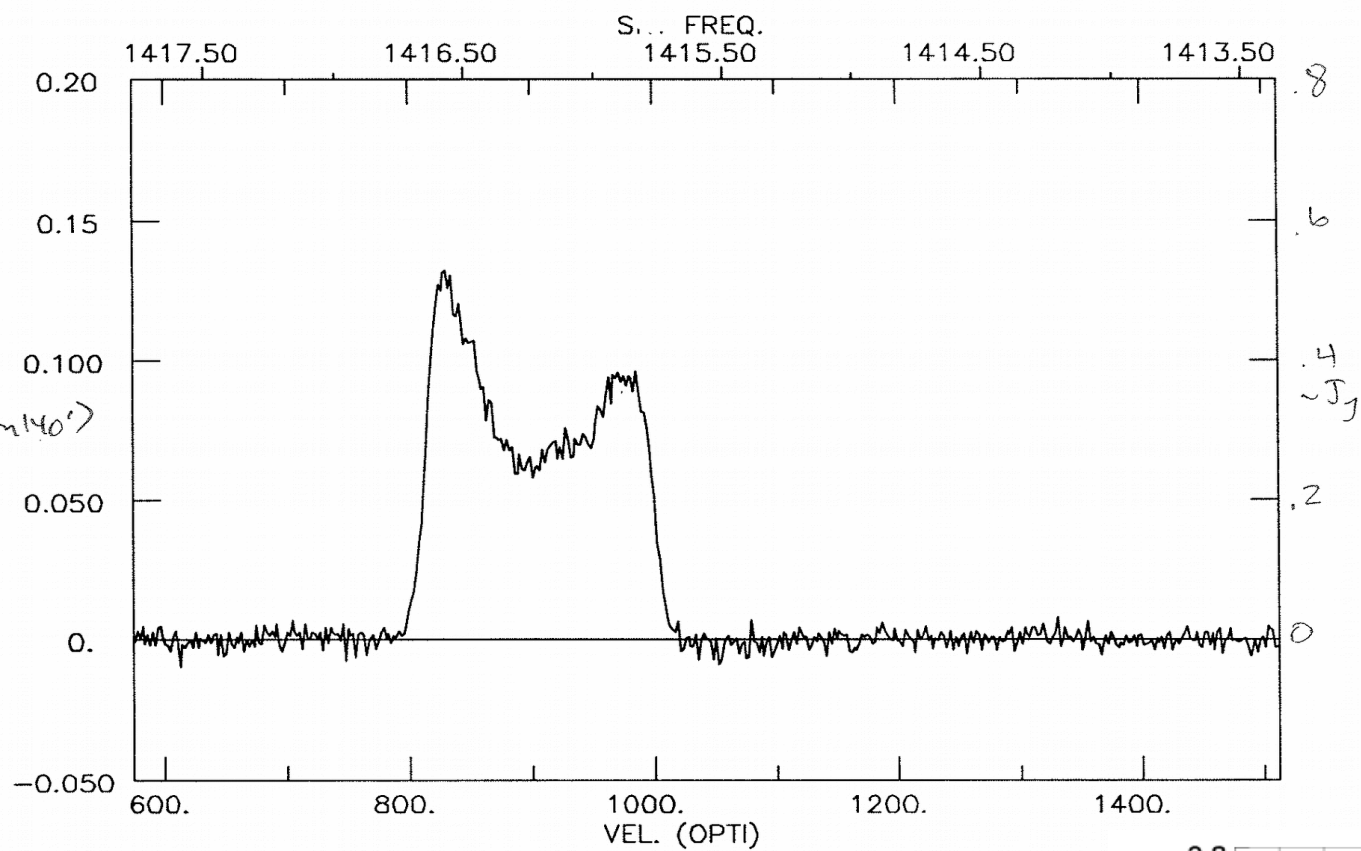
Relation between the rotation velocity of a galaxy (deduced from the width of its emission lines) and its absolute magnitude

By comparing with the apparent magnitude, gives the distance modulus

Similar relations for elliptical galaxies (Faber-Jackson, Fundamental plane)



Tully-Fisher



Type Ia supernovae : standard candle

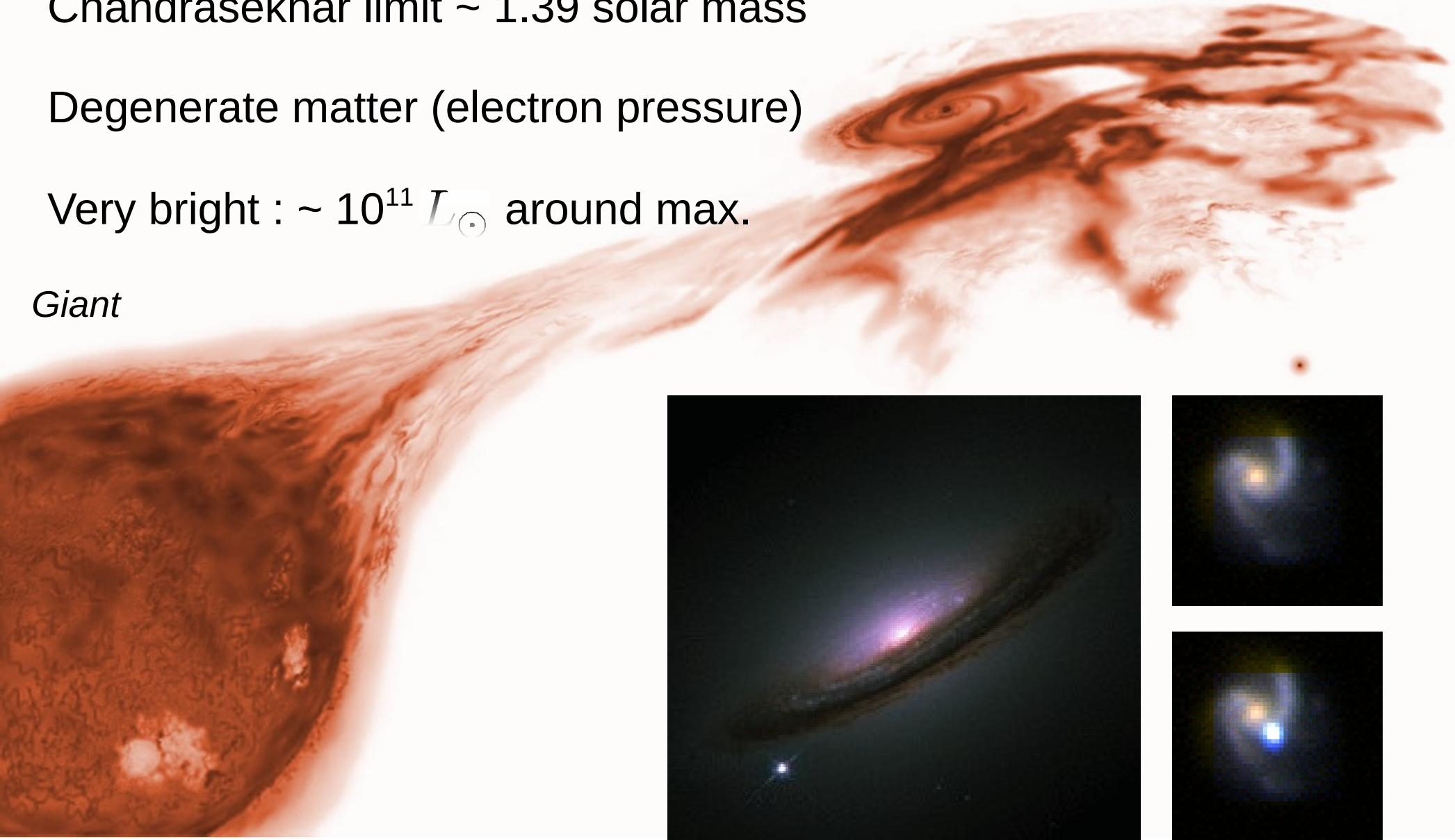
White dwarf accreting mass from its partner
Chandrasekhar limit ~ 1.39 solar mass

White dwarf

Degenerate matter (electron pressure)

Very bright : $\sim 10^{11} L_{\odot}$ around max.

Giant



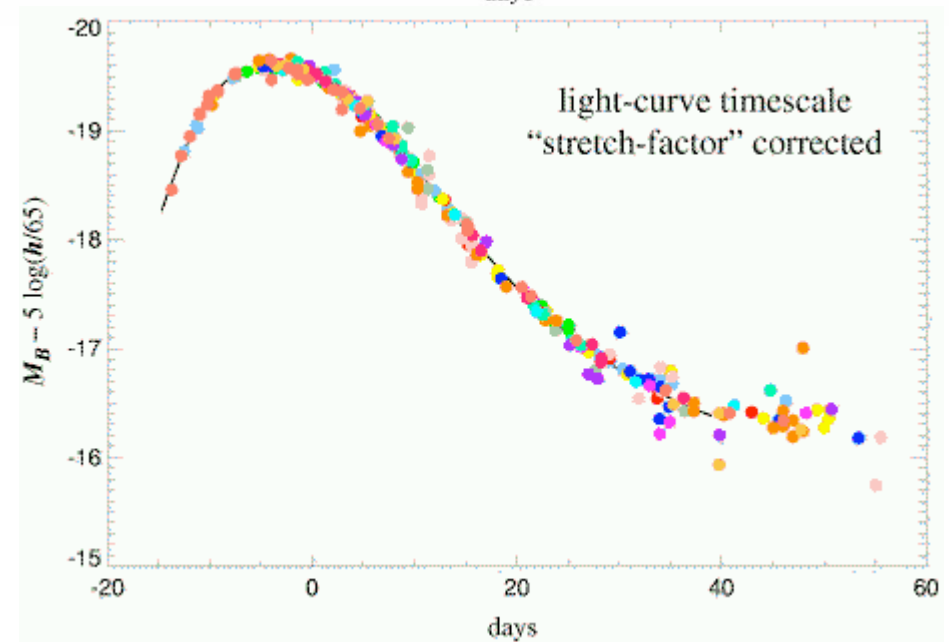
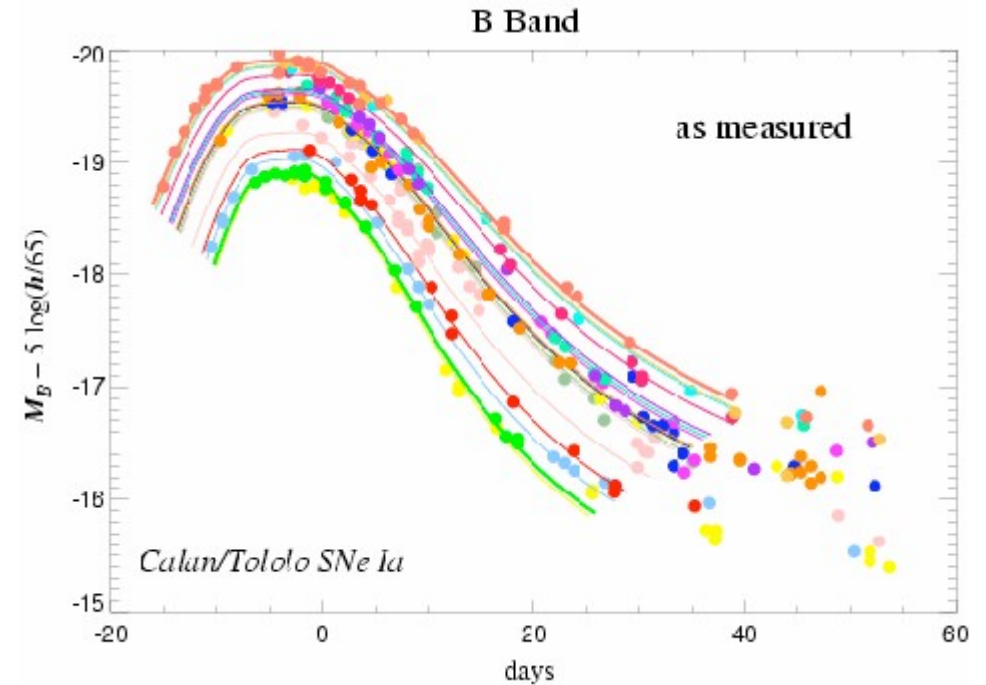
Type Ia supernovae : standard candle

SNIa = standard candles

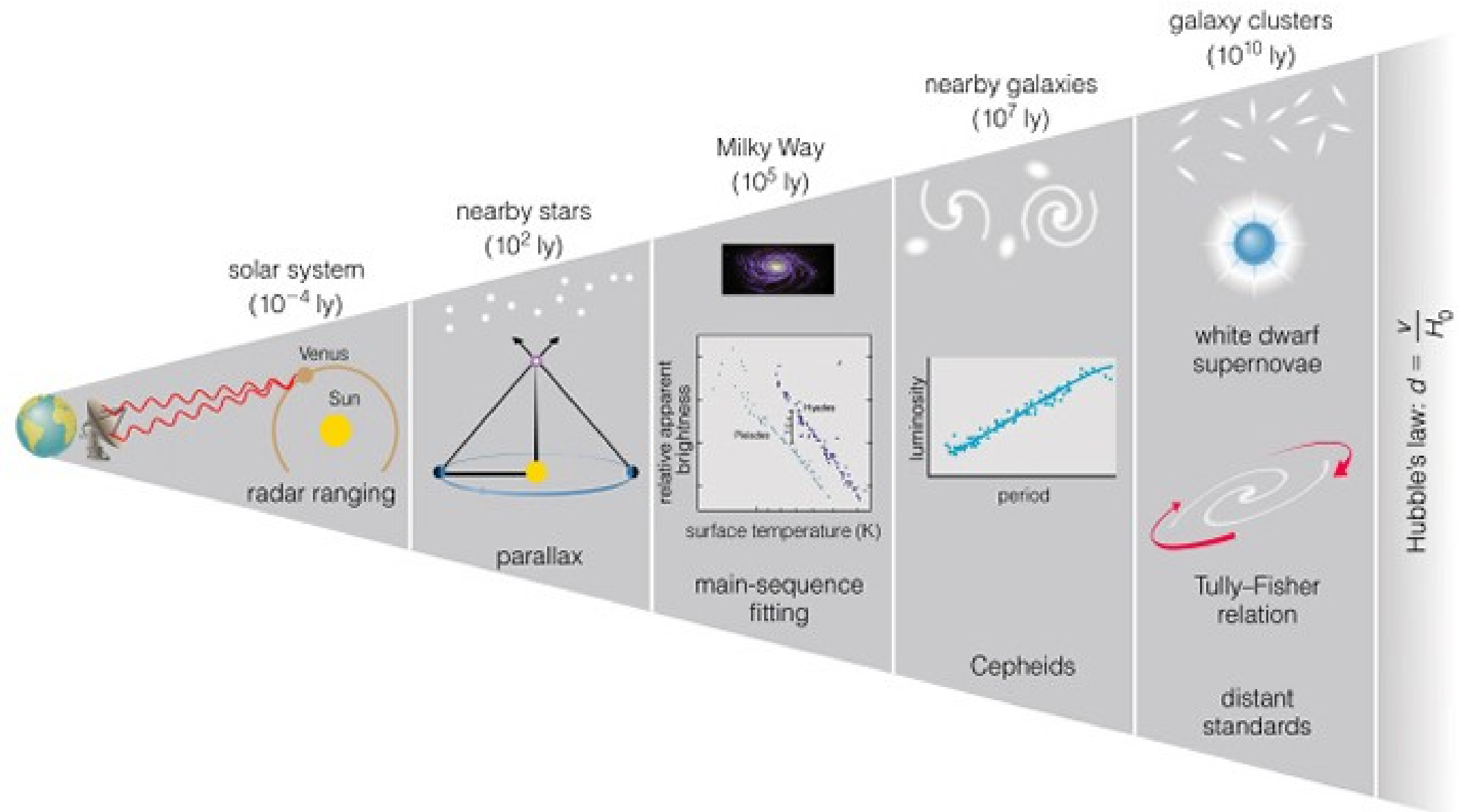
Raises in a few weeks,
decreases in a few months

« Standardization » :
Correlation of the peak
luminosity with the decay
(« stretch ») and the color.

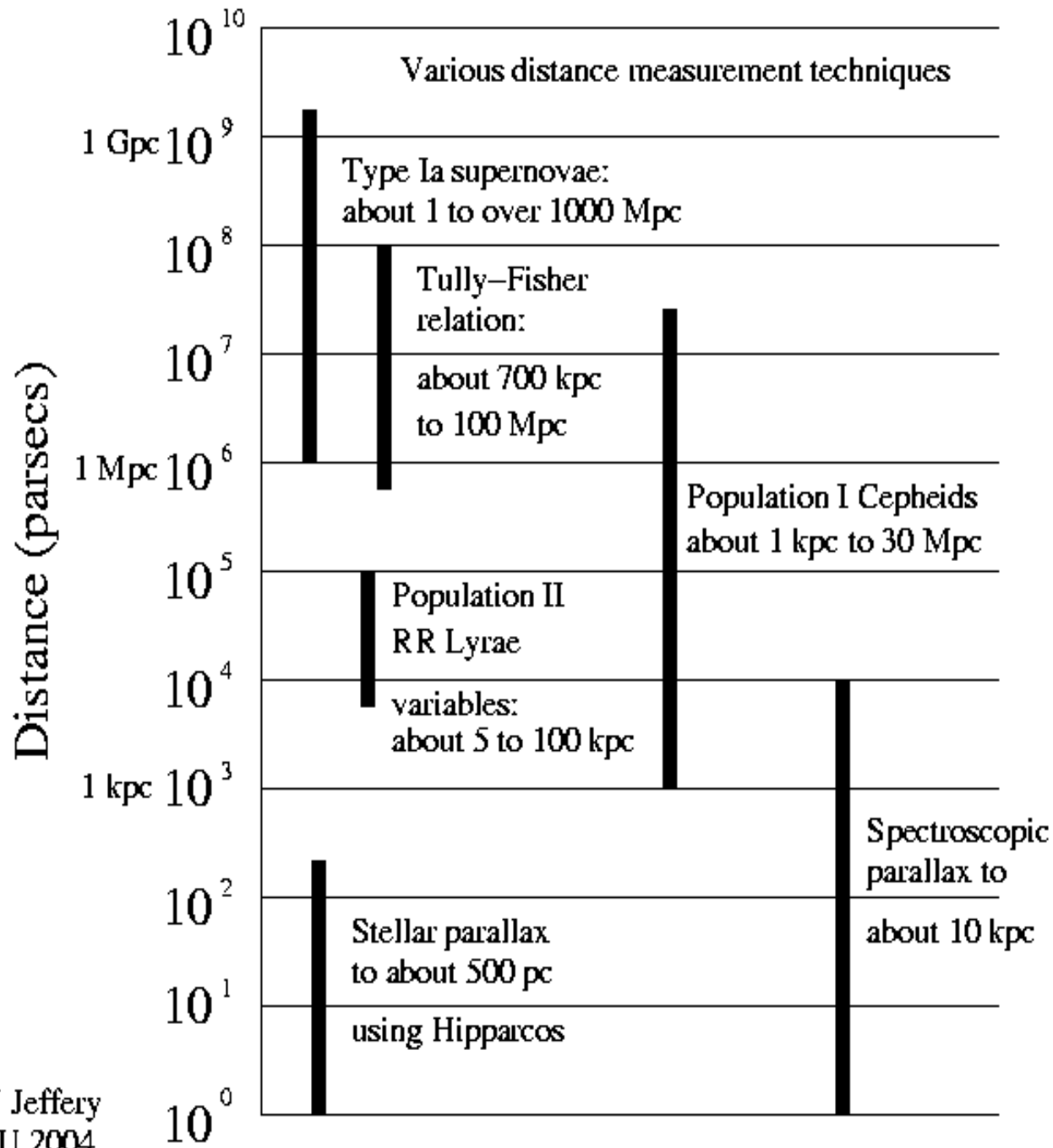
Distance estimator up to $z \sim 1$
and beyond.

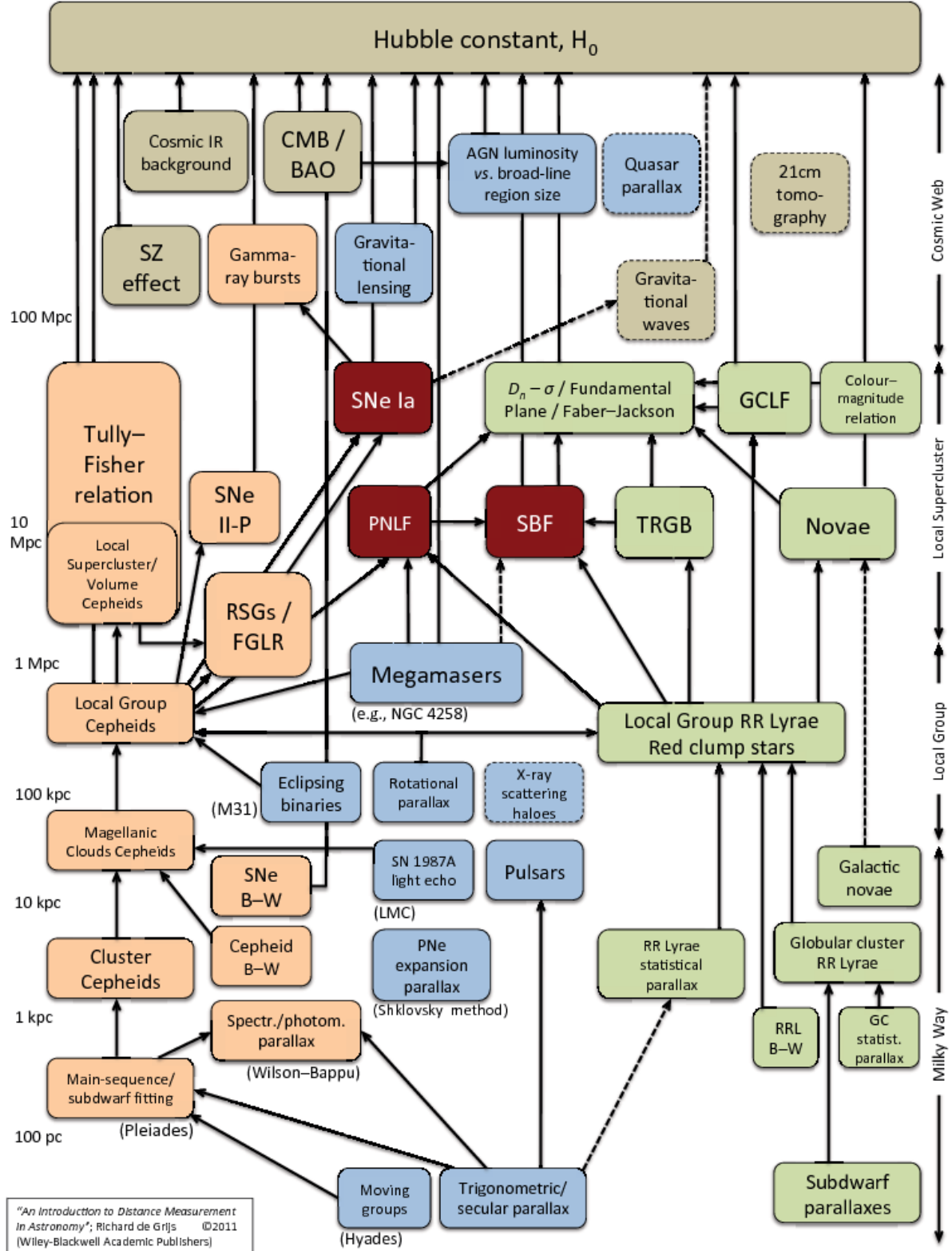


Cosmic distance ladder



A Cartoon of the Distance Ladder





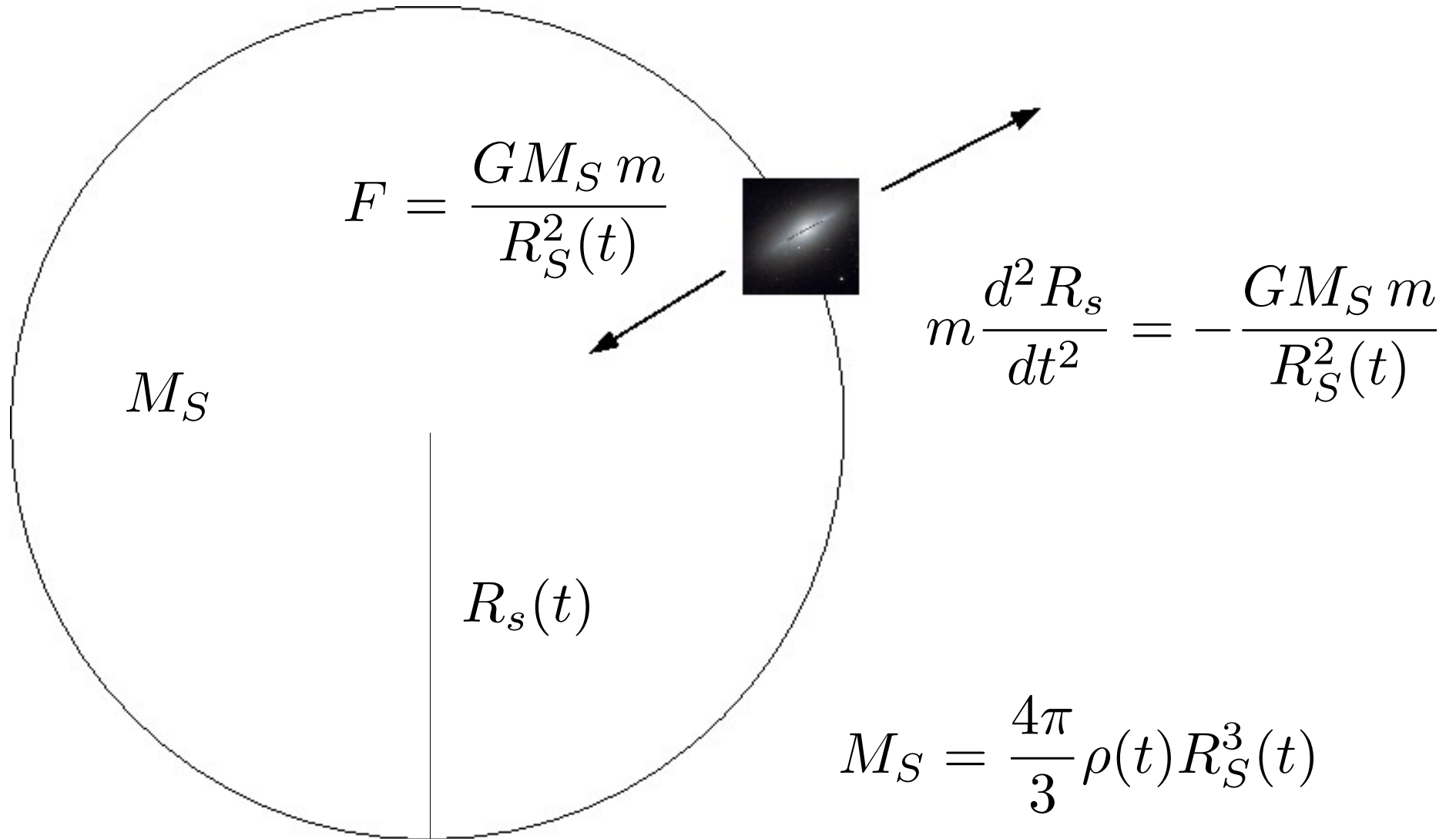
"An Introduction to Distance Measurement in Astronomy"; Richard de Grijs ©2011 (Wiley-Blackwell Academic Publishers)

1.3

Dynamics of the Universe

*How a simple (and wrong) Newtonian description
may (partly) capture the Universe dynamics*

A very simple (and wrong) model



Newtonian dynamics

$$R_s = r_s a(t)$$

$$\ddot{a} = -\frac{4\pi G\rho(t)}{3} a(t)$$

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G\rho(t)}{3} + \frac{2C}{r_s^2 a^2(t)}$$

If $C > 0$: expansion will never stop.

If $C < 0$: expansion stops at $a_{\max}^2 = -GM_S/Cr_s$
Followed by contraction

$C = 0$: limit case (expand forever but $\dot{a} \rightarrow 0$)

Analogy with the dynamic of a ball bouncing on Earth

From Newton to General Relativity

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G\rho(t)}{3} + \frac{2C}{r_s^2 a^2(t)} \quad \longrightarrow \quad \left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G\rho(t)}{3} - \frac{kc^2}{R_0^2} \frac{1}{a(t)^2}$$
$$\frac{2C}{r_s^2} \quad \longrightarrow \quad -\frac{kc^2}{R_0^2}$$

$$k = -1, +1, 0$$

$$R(t) = a(t)R_0$$

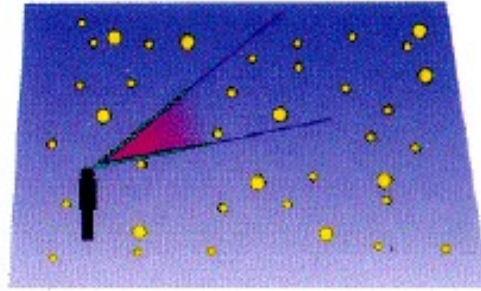
If $k = -1$ ($C > 0$) : negative curvature
expansion will never stop.

If $k = +1$ ($C < 0$) : positive curvature
expansion stops, contraction follows

If $k = 0$ ($C = 0$) : flat universe (limit case)

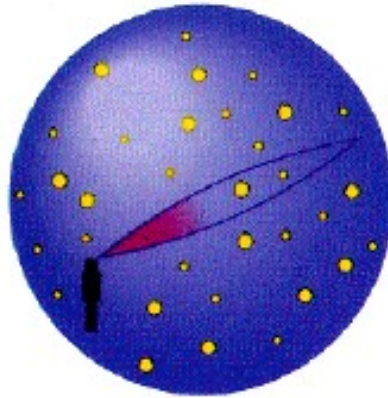
[GR, dynamics...]

$$k = 0$$



Flat universe

$$k = +1$$

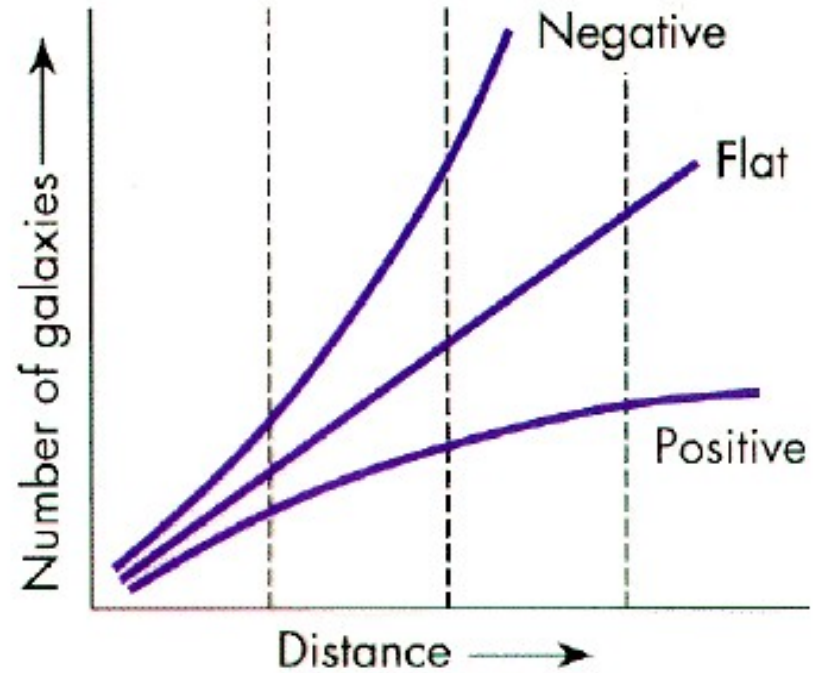


Positively curved universe

$$k = -1$$



Negatively curved universe



[Curvature, General Relativity...]

Hubble rate and critical density

$$H(t) = \frac{\dot{R}(t)}{R(t)} = \frac{\dot{a}(t)}{a(t)}$$

$$z = \frac{\lambda_{\text{obs}}}{\lambda_{\text{emit}}} - 1$$

$$1 + z = \frac{\lambda_{\text{obs}}}{\lambda_{\text{emit}}} = \frac{a(t_{\text{obs}})}{a(t_{\text{emit}})} = \frac{1}{a(t)}$$

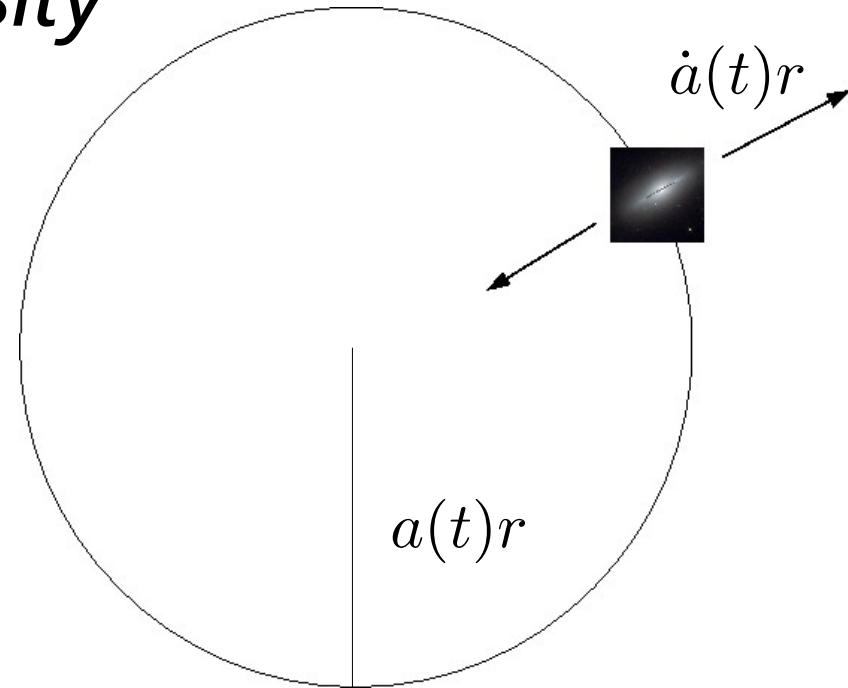
$$H^2(t) = \frac{8\pi G}{3} \rho_c(t) = \frac{8\pi G}{3c^2} \varepsilon_c(t)$$

$$H_0 = H(t_0) \simeq 70 \text{ km/s/Mpc}$$

$$\rho_c(t) = \frac{3}{8\pi G} H^2(t) \quad \varepsilon_c(t) = \frac{3c^2}{8\pi G} H^2(t)$$

$$\rho_{c,0} = \rho(t_0) \simeq 9 \times 10^{-27} \text{ kg m}^{-3} \simeq 1.4 \times 10^{11} \text{ M}_\odot \text{ Mpc}^{-3}$$

$$\varepsilon_{c,0} \simeq 5200 \text{ MeV m}^{-3} \quad 5 \text{ protons per m}^3$$



r comobile distance
 $a(t)$ scale factor

$$a(t_0) = 1$$

Friedmann Equation

$$H^2(t) = \left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G}{3} \rho - \frac{kc^2}{R_0^2 a^2}$$

$$H^2(t) = \left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G}{3c^2} \varepsilon - \frac{kc^2}{R_0^2 a^2}$$

$$-\frac{kc^2}{R_0^2 a^2(t) H^2(t)} = 1 - \frac{\rho(t)}{\rho_c(t)} = 1 - \Omega(t)$$

$$\Omega(t) = \frac{\rho(t)}{\rho_c(t)} = \frac{\varepsilon(t)}{\varepsilon_c(t)}$$

$$\Omega_k(t) = -\frac{kc^2}{R_0^2 a^2(t) H^2(t)} = 1 - \Omega(t)$$

Fluid equation

Adiabatic expansion : $\delta Q = dU + pdV = 0$ i.e. $dS = 0$

Fluid equation :
$$\frac{\ddot{a}}{a} = -\frac{4\pi G}{3c^2}(\varepsilon + 3p)$$

Friedmann equations :

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G}{3c^2}\varepsilon - \frac{kc^2}{R_0^2 a^2}$$

$$\frac{\ddot{a}}{a} = -\frac{4\pi G}{3c^2}(\varepsilon + 3p)$$

$$\dot{\varepsilon} + 3\frac{\dot{a}}{a}(\varepsilon + p) = 0$$

Needed : equation of state : $p = w\varepsilon = w\rho/c$

Equations of state

Non-relativistic matter : $p_m \ll \varepsilon_m = \rho_m/c^2$ $w_m \simeq 0$

Light, relativistic matter (photons, neutrinos, ...) :

$$p_r = \frac{\varepsilon_r}{3} \quad w_r = \frac{1}{3}$$

Density evolution with time :

$$\dot{\varepsilon} + 3\frac{\dot{a}}{a}(\varepsilon + p) = 0 \quad \varepsilon(t) = \varepsilon(t_0)a(t)^{-3(1+w)}$$

$$\varepsilon_m(t) = \varepsilon_m(t_0)/a(t)^3$$

$$\varepsilon_r(t) = \varepsilon_r(t_0)/a(t)^4$$

Eras ! Light era, matter era, ...

[Dynamics, Thermal history...]

Friedmann Equations

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G}{3c^2}\varepsilon - \frac{kc^2}{R_0^2 a^2}$$

$$\dot{\varepsilon} + 3\frac{\dot{a}}{a}(\varepsilon + p) = 0$$

$$\frac{\ddot{a}}{a} = -\frac{4\pi G}{3c^2}(\varepsilon + 3p)$$

Describe an evolving, non static Universe (scale factor $a(t)$)

Einstein → How to stabilise it to get a static Universe ?

Friedmann Equations (with Λ)

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G}{3c^2}\varepsilon - \frac{kc^2}{R_0^2 a^2} + \frac{\Lambda c^2}{3}$$

$$\dot{\varepsilon} + 3\frac{\dot{a}}{a}(\varepsilon + p) = 0$$

$$\frac{\ddot{a}}{a} = -\frac{4\pi G}{3c^2}(\varepsilon + 3p) + \frac{\Lambda c^2}{3}$$

$$\Lambda = \frac{8\pi G}{c^4}\varepsilon_\Lambda = \frac{8\pi G}{c^2}\rho_\Lambda \quad p_\Lambda = -\rho_\Lambda c^2 \quad w_\Lambda = -1$$

$$\Omega_\Lambda = \frac{\rho_\Lambda}{\rho_c}$$

Universe driven by its contents

Non-relativistic matter : $\Omega_m = \frac{\rho_m}{\rho_c} = \frac{\varepsilon_m}{\varepsilon_c}$

« Curvature » : $\Omega_k(t) = -\frac{kc^2}{R_0^2 a^2(t) H^2(t)} = 1 - \Omega(t)$

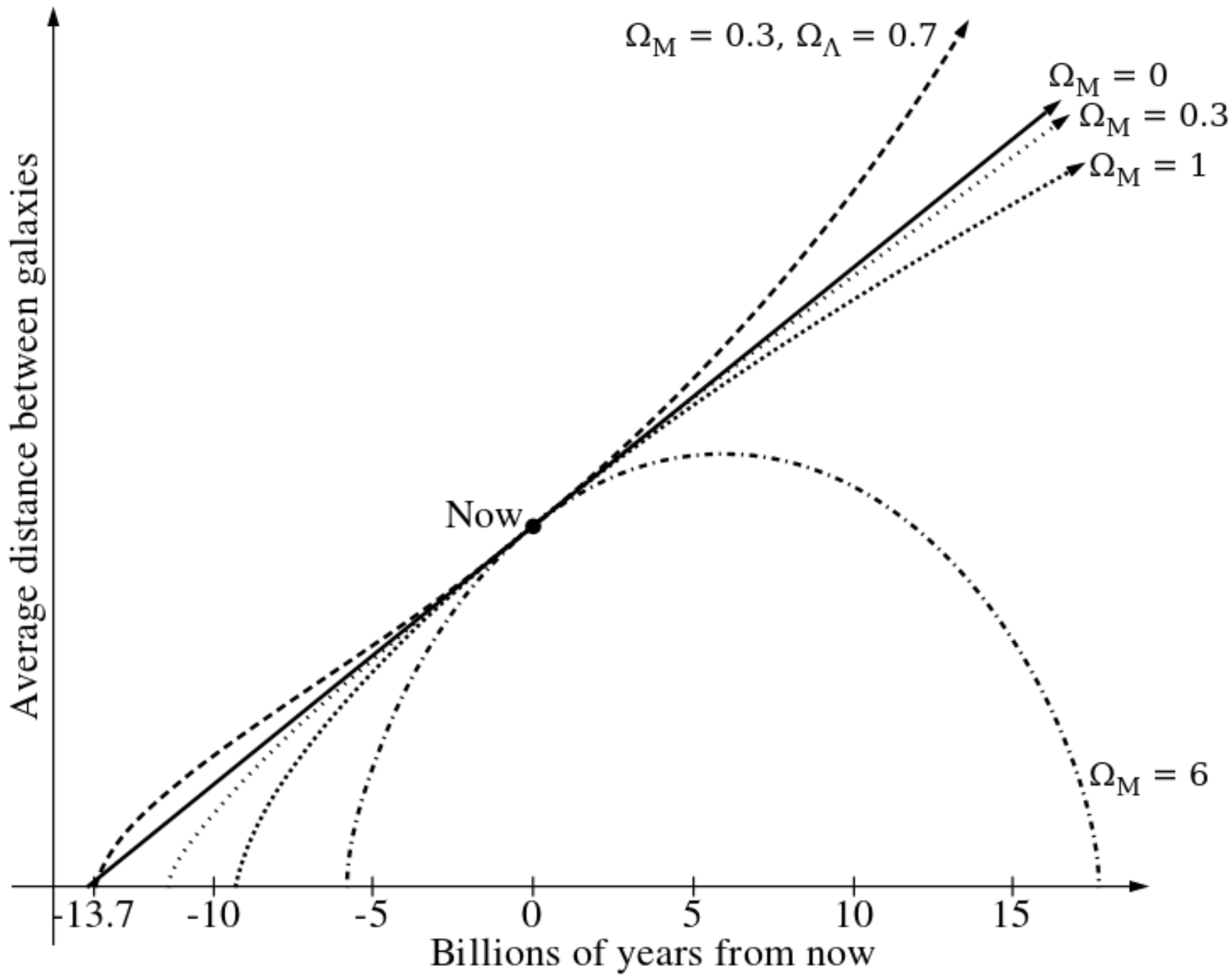
Photons : $\Omega_\gamma = \frac{\rho_\gamma}{\rho_c} = \frac{\varepsilon_\gamma}{\varepsilon_c}$

Neutrinos : $\Omega_\nu = \frac{\rho_\nu}{\rho_c} = \frac{\varepsilon_\nu}{\varepsilon_c}$

$$\Omega_r = \frac{\rho_r}{\rho_c} = \frac{\varepsilon_r}{c}$$

« Cosmological constant » : $\Lambda = \frac{8\pi G}{c^4} \varepsilon_\Lambda = \frac{8\pi G}{c^2} \rho_\Lambda$

Exotic fields...



[Universe Dynamics...]

1.4

Contents of the Universe

*The dynamics of the Universe
is driven by its contents.*

So, what is the Universe made of?

Stars

Typical luminosity in a few
hundreds Mpc :

$$j_{\star,B} \simeq 1.2 \times 10^8 L_{\odot,B} / \text{Mpc}^3$$

In the vicinity of the
Sun (1kpc) :

$$\langle M/L_B \rangle \simeq 4M_{\odot} / L_{\odot,B}$$

If typical,

$$\rho_{\star,0} = j_{\star,B} * \langle M/L_B \rangle \simeq 5 \times 10^8 M_{\odot} / \text{Mpc}^3$$

$$\Omega_{\star} \simeq 0.003 - 0.004$$

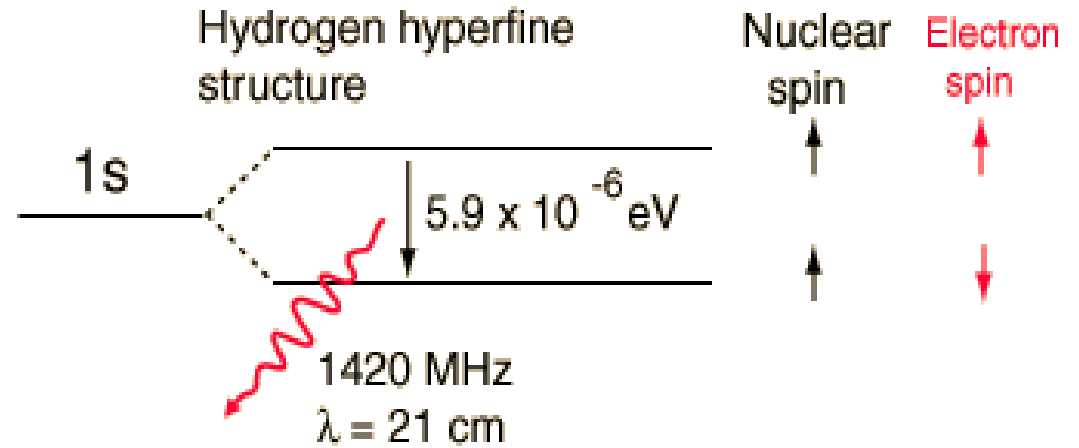


Gas (HI , H_2 , ...)

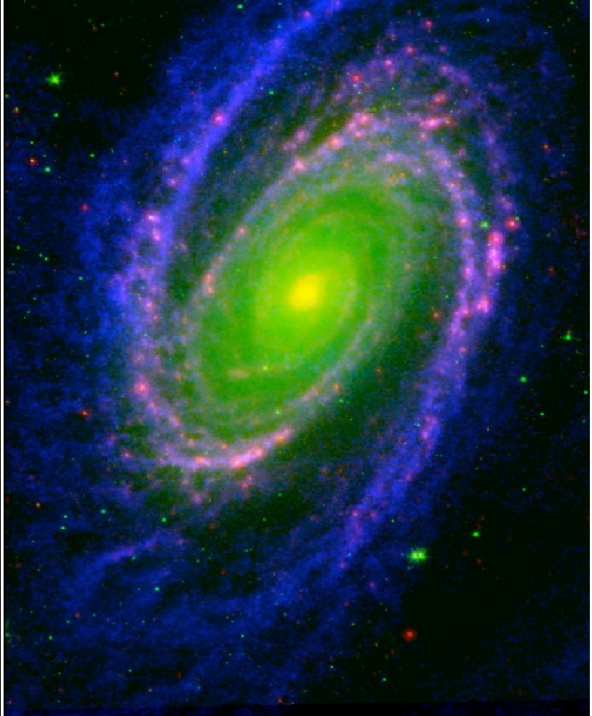
HI : 21 cm line (radio)

$$\Omega_{\text{HI}} \simeq 3 \times 10^{-4}$$

$$\Omega_{\text{H}_2} \simeq 2 \times 10^{-4}$$



atomic hydrogen (blue)
3.6 microns (green)
24 microns (red)



atomic hydrogen (blue)
3.6 microns (red)
8 microns (green)



atomic hydrogen (blue)
8 microns (green)
24 microns (red)



Spiral Galaxy Messier 81

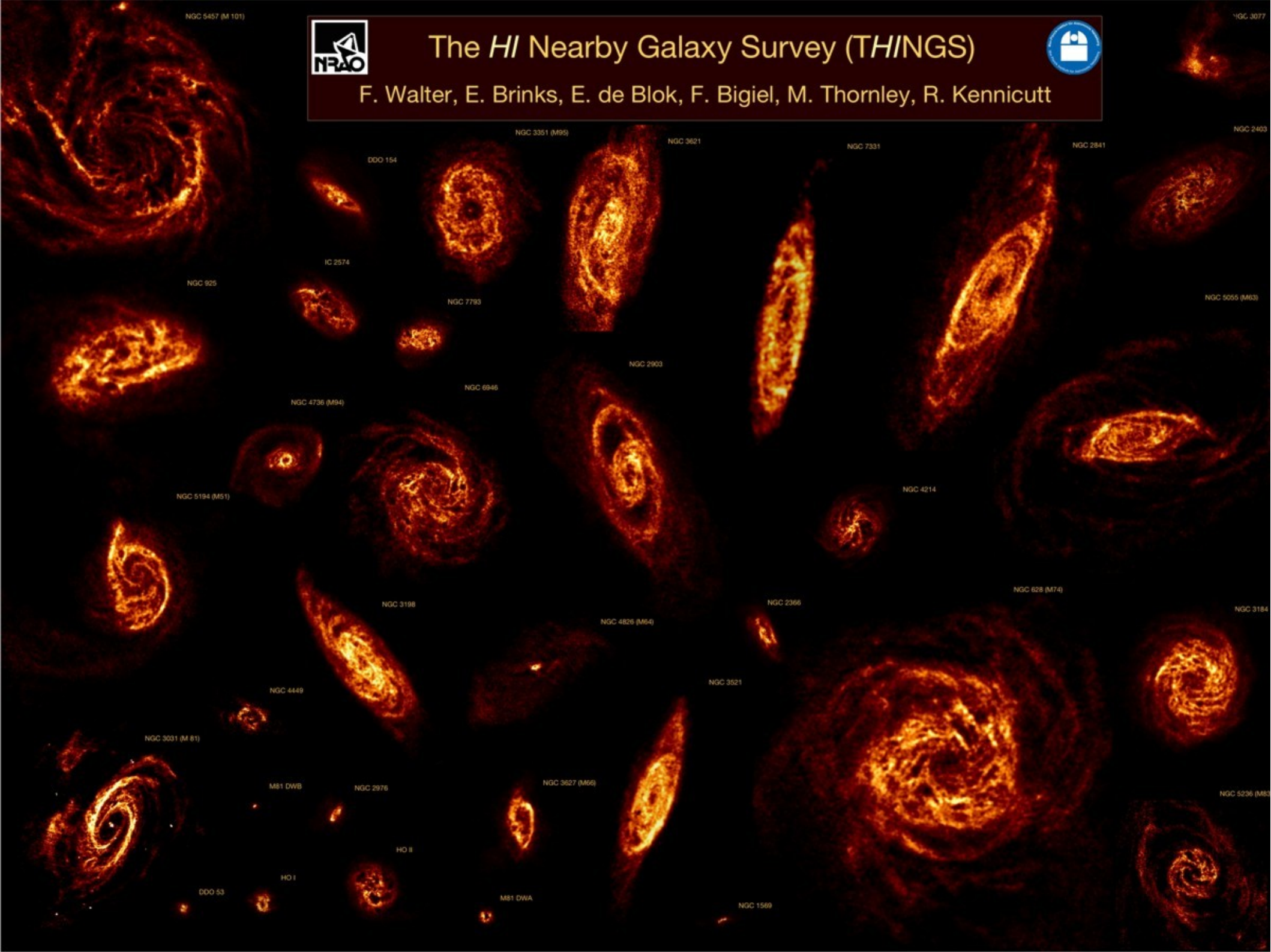
NASA Spitzer Space Telescope and NRAO VLA



The *HI* Nearby Galaxy Survey (*THINGS*)



F. Walter, E. Brinks, E. de Blok, F. Bigiel, M. Thornley, R. Kennicutt



NGC 925

NGC 5194 (M51)

NGC 3031 (M 81)

DDO 53

DDO 154

IC 2574

NGC 4736 (M94)

NGC 4449

M81 DWB

NGC 2976

NGC 7793

NGC 6946

NGC 3198

HO II

HO I

NGC 3351 (M95)

NGC 3621

NGC 2903

NGC 4826 (M64)

NGC 3627 (M66)

M81 DWA

NGC 7331

NGC 2366

NGC 3521

NGC 1569

NGC 2841

NGC 628 (M74)

NGC 4214

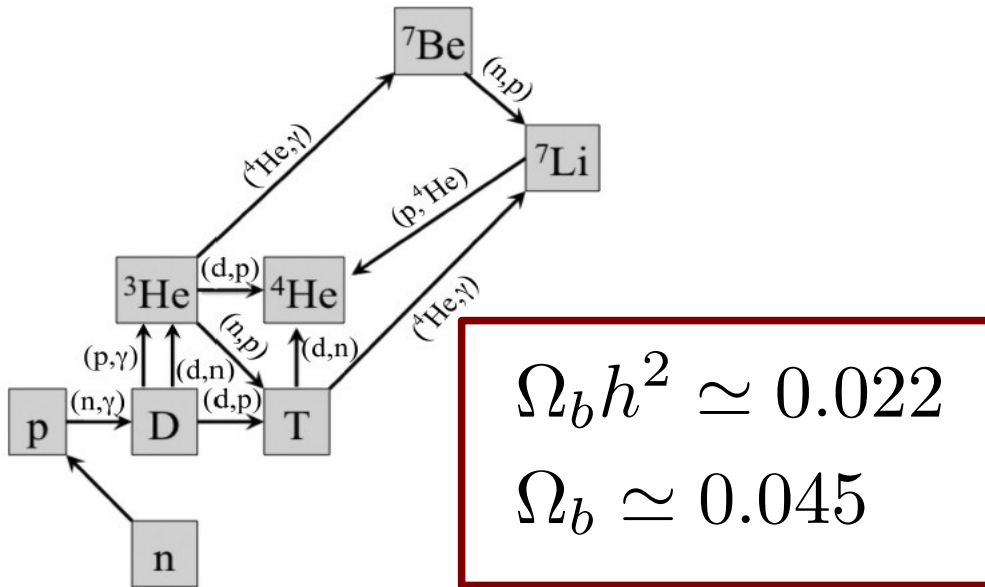
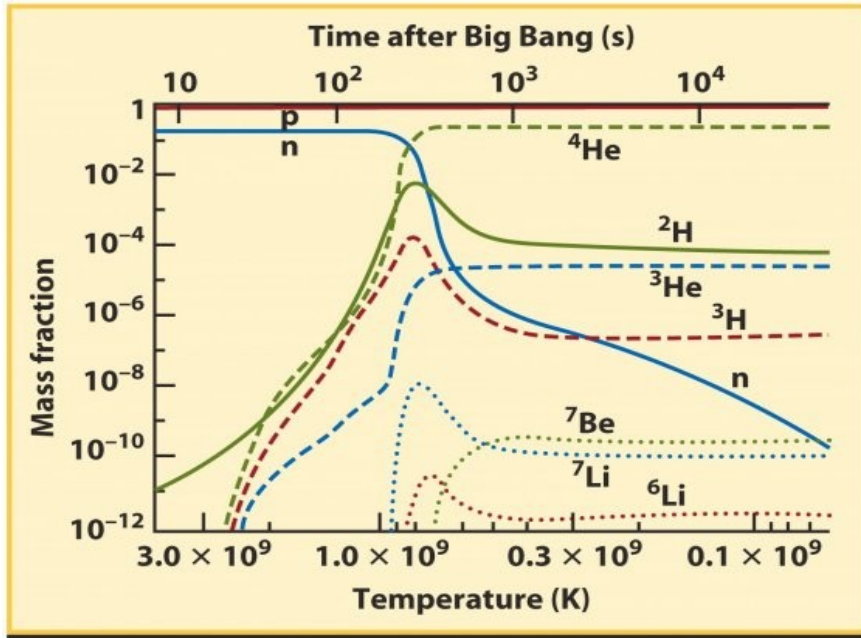
NGC 2403

NGC 5055 (M63)

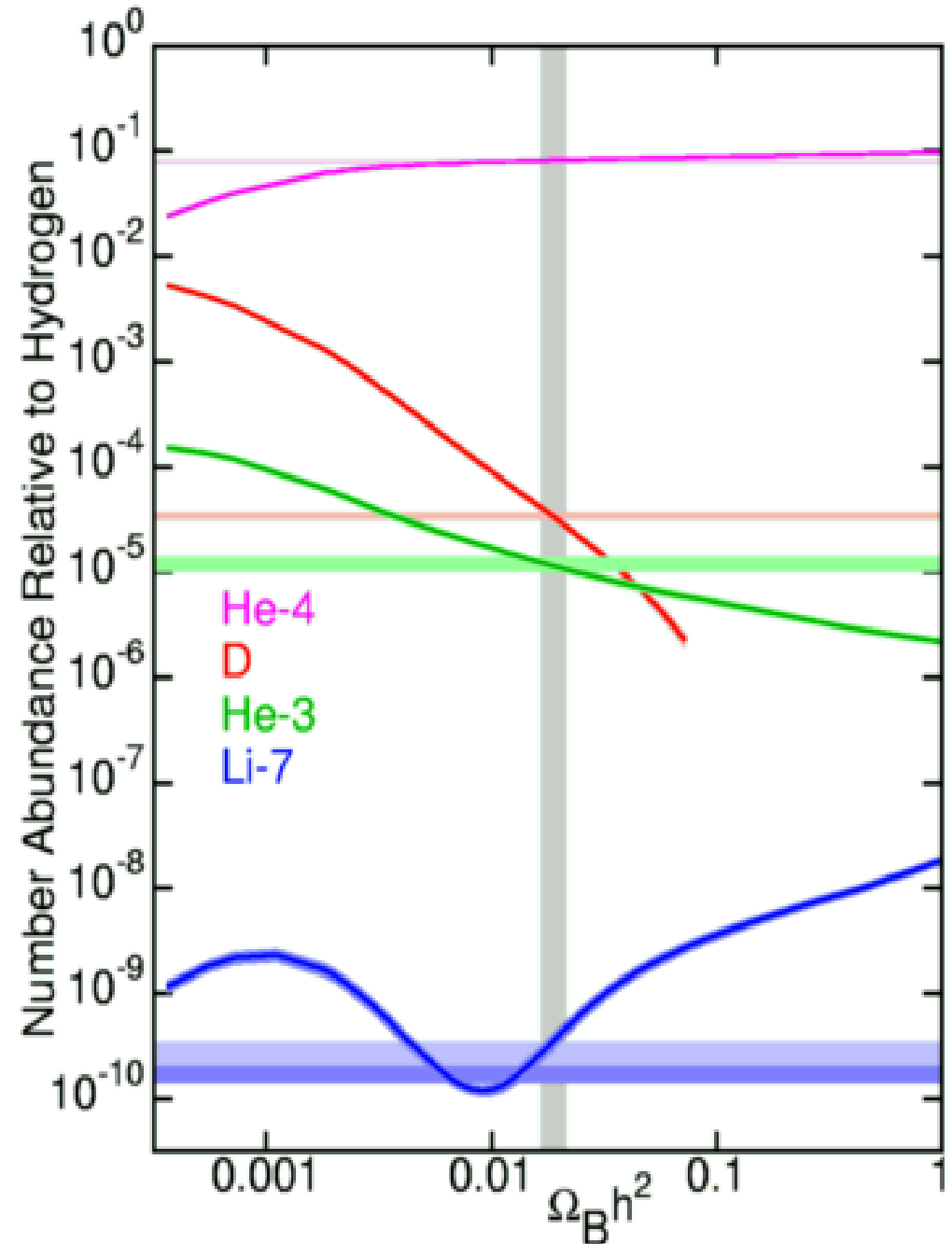
NGC 3184

NGC 5236 (M83)

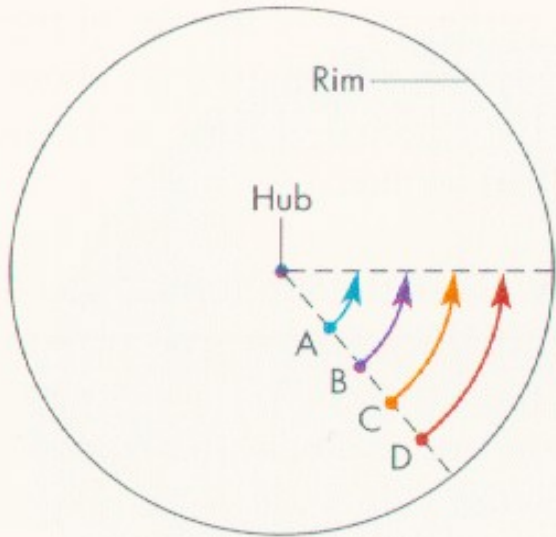
Baryonic matter : primordial nucleosynthesis



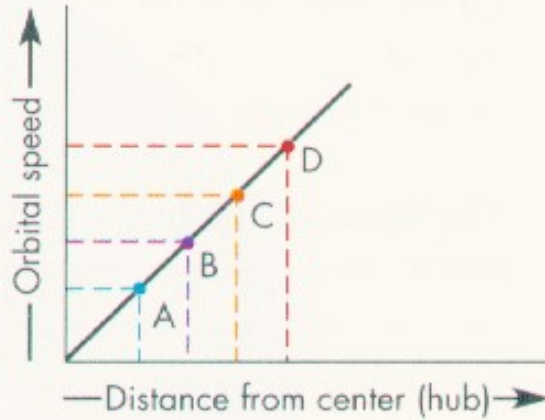
(CMB : Planck, 2014)



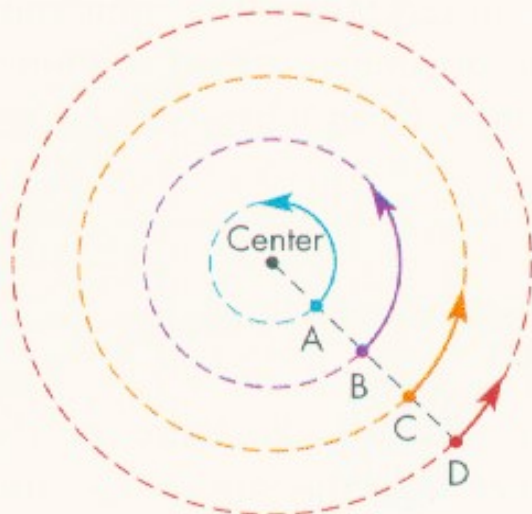
Galaxies : rotation curves



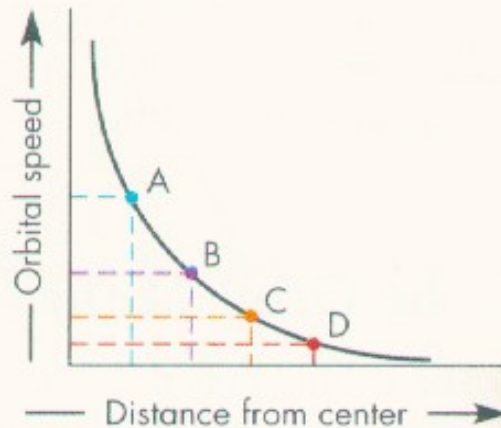
Wheel-like rotation



Rotation curve for wheel-like rotation



Planet-like rotation



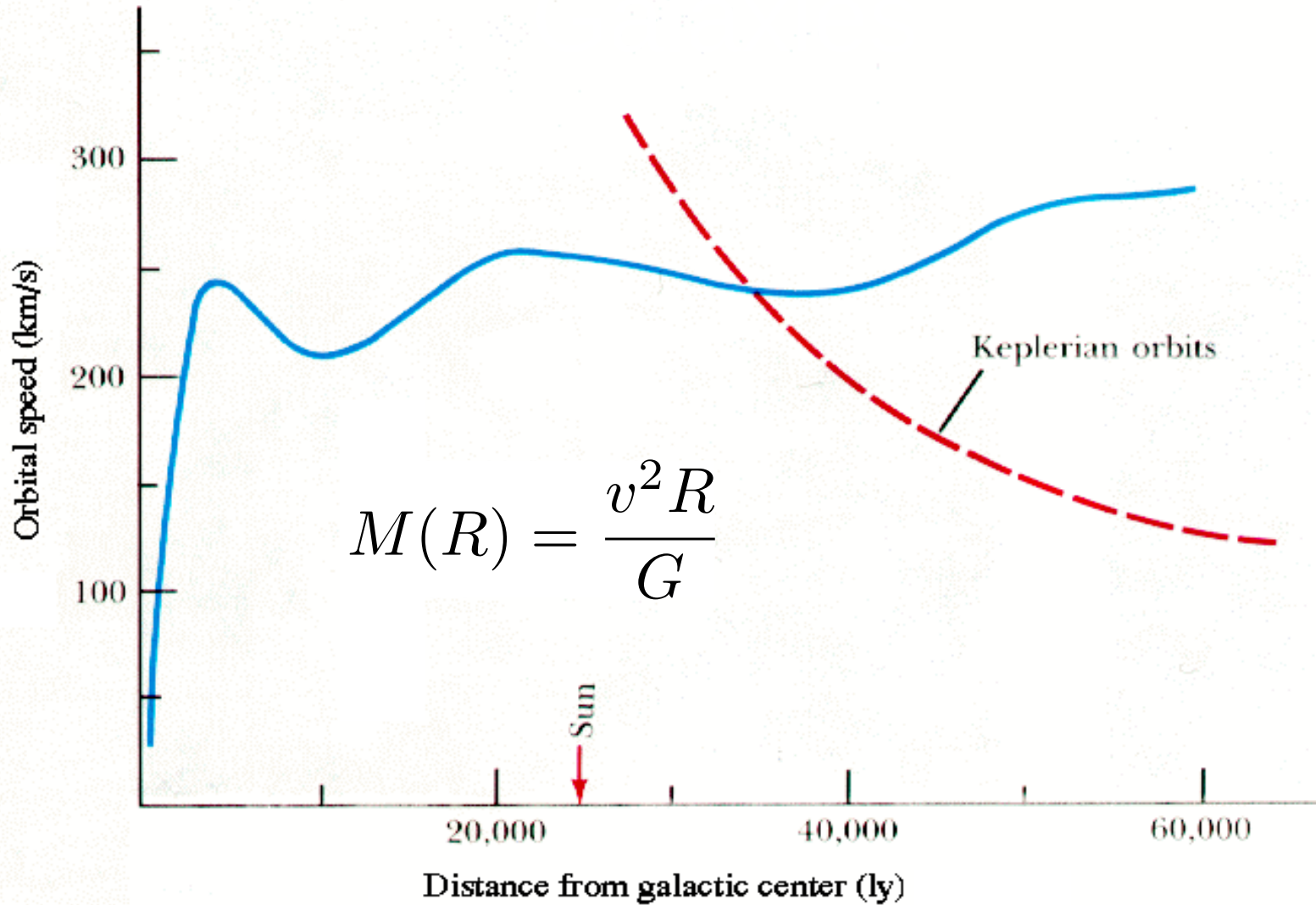
Rotation curve for planet-like rotation

$$\frac{GM(r)}{r^2} = \frac{v^2}{r}$$

$$v = \sqrt{\frac{GM(r)}{r}}$$

$$M(R) = \frac{v^2 R}{G}$$

Galaxies : rotation curves

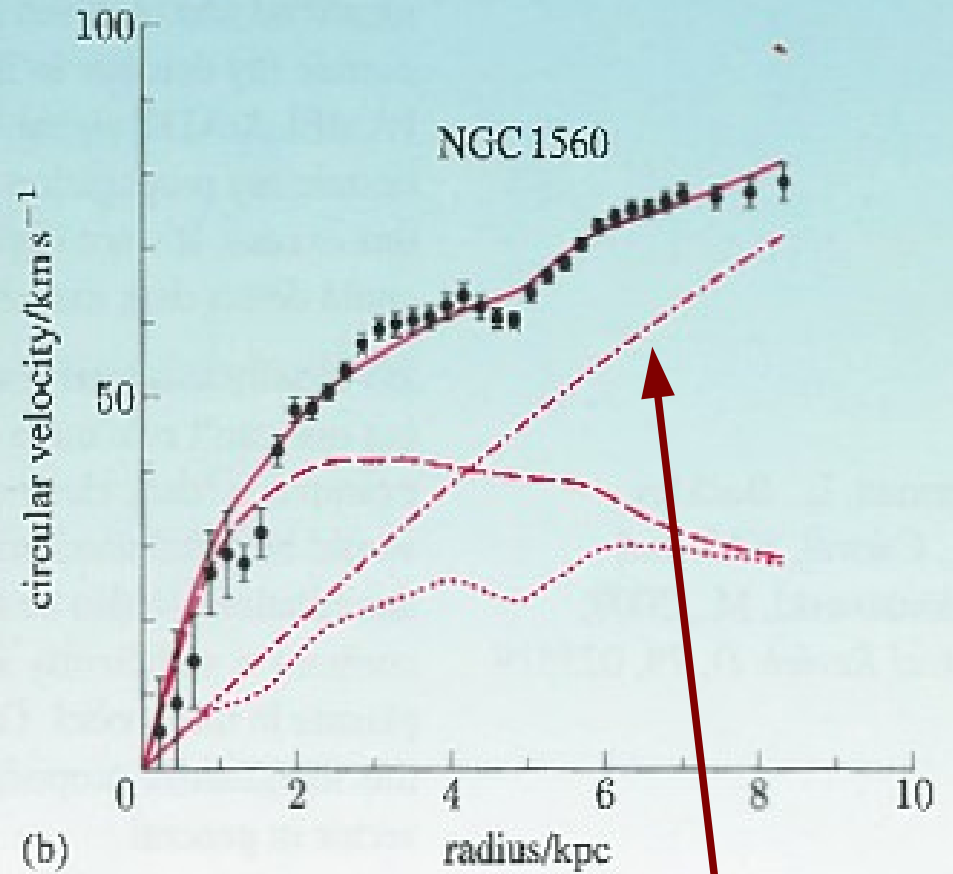


$$M(R) \simeq 10^{11} M_{\odot} \left(\frac{v}{220 \text{ km s}^{-1}} \right)^2 \left(\frac{R}{8.5 \text{ kpc}} \right)$$

Galaxies : rotation curves



(a)



(b)

$$\Omega_{\text{gal}} \simeq 0.04 \text{ to } 0.16$$

Dark matter halo !

Galaxy Clusters

Virialized objects $V + 2T = 0$

$$\frac{1}{2}M \langle v^2 \rangle = \frac{\alpha GM^2}{2 r_h}$$

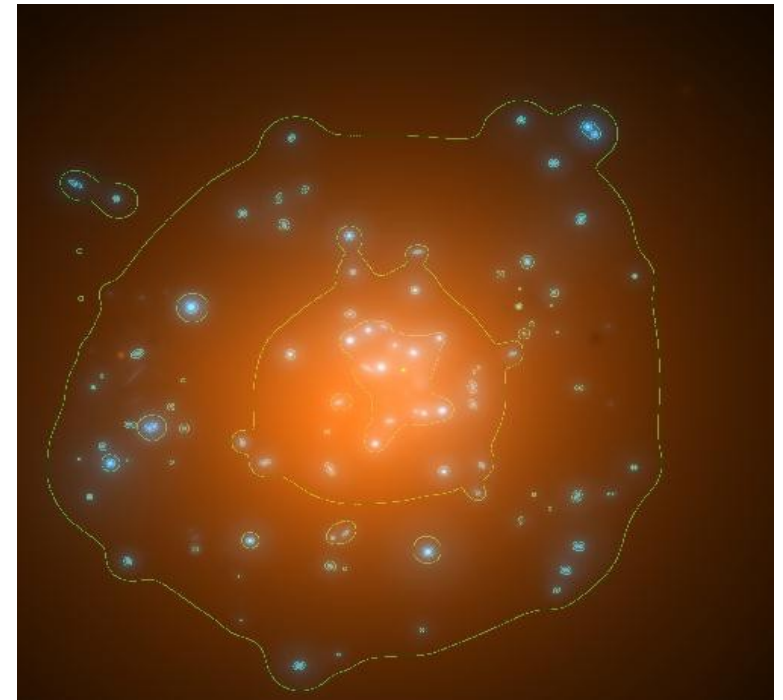
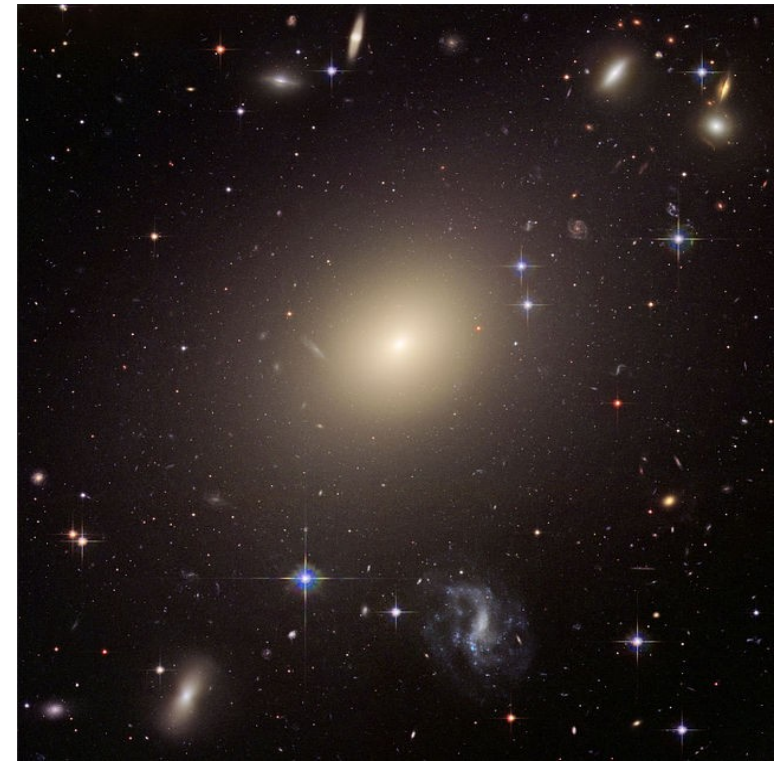
Mass may be deduced from velocity dispersion :

$$M = \frac{\langle v^2 \rangle r_h}{\alpha G}$$

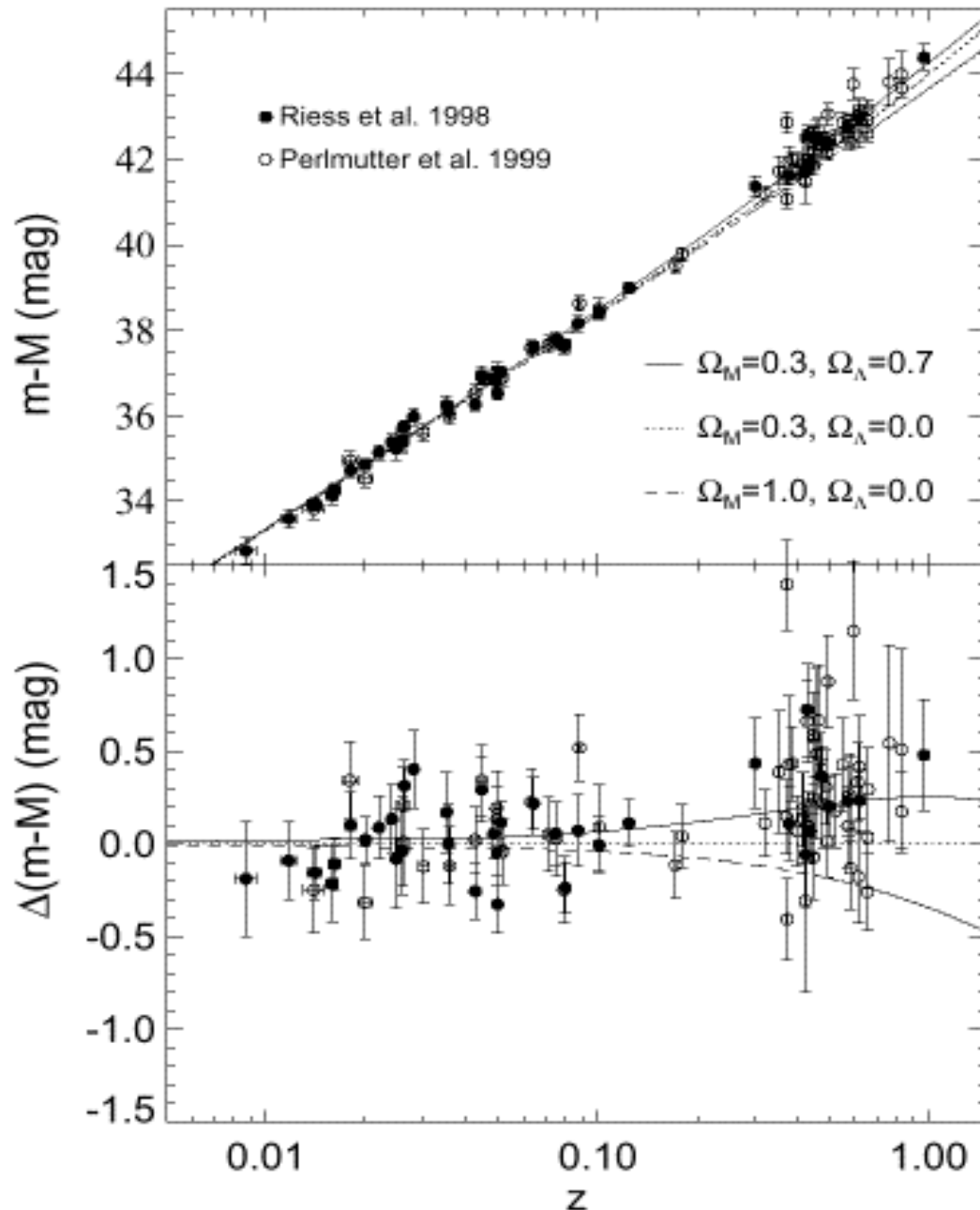
Hot X-ray emitting gas under hydrostatic equilibrium

$$\Omega_{\text{clusters}} \simeq 0.2$$

« Dark » matter again !



Type Ia Supernovae and Dark Energy



Expansion is accelerating!

Best model :

$$\Omega_m \simeq 0.3 \quad \Omega_\Lambda \simeq 0.7$$

The cosmological constant is back!

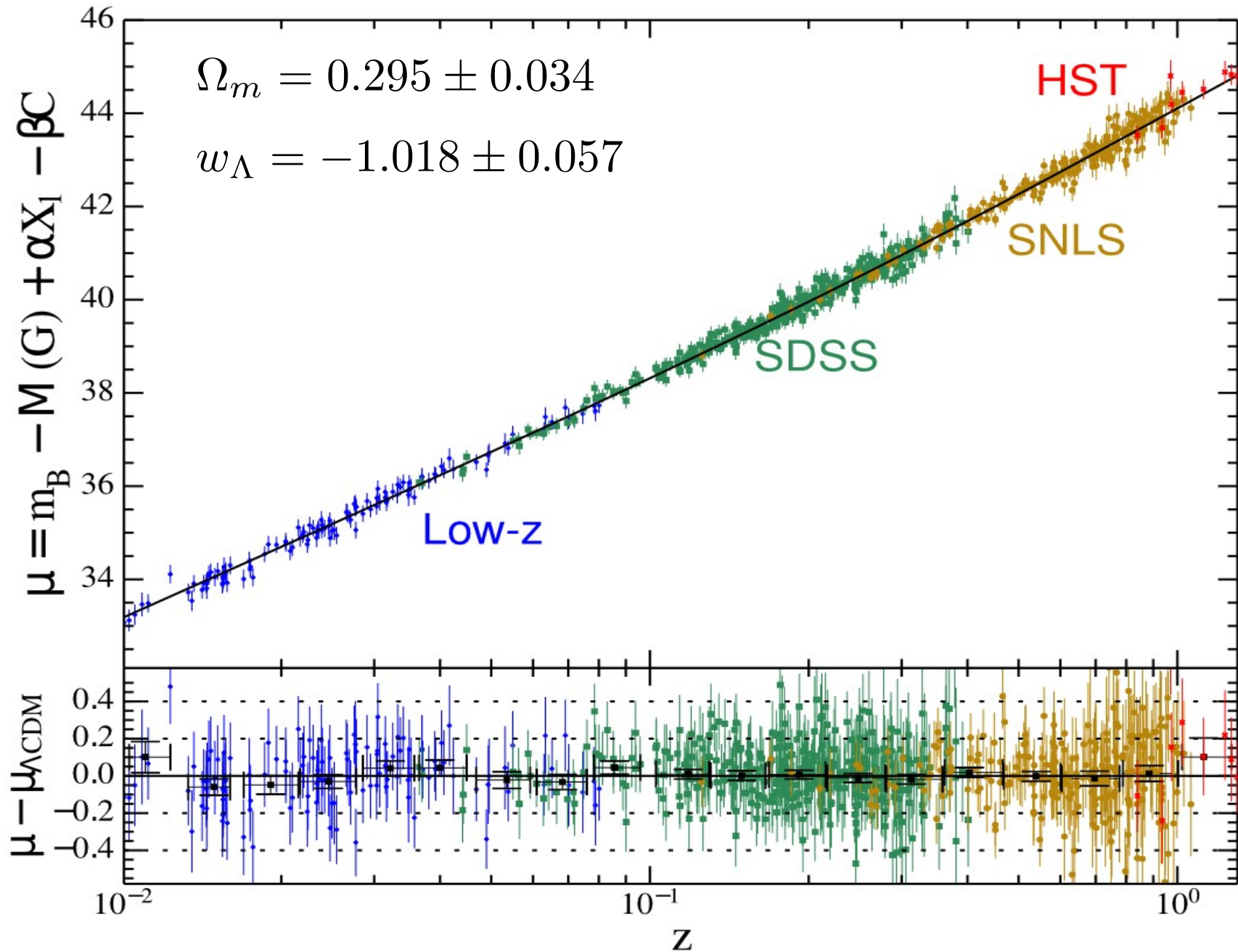
May be seen as an dark energy with

$$p = w\rho = w\epsilon/c^2$$

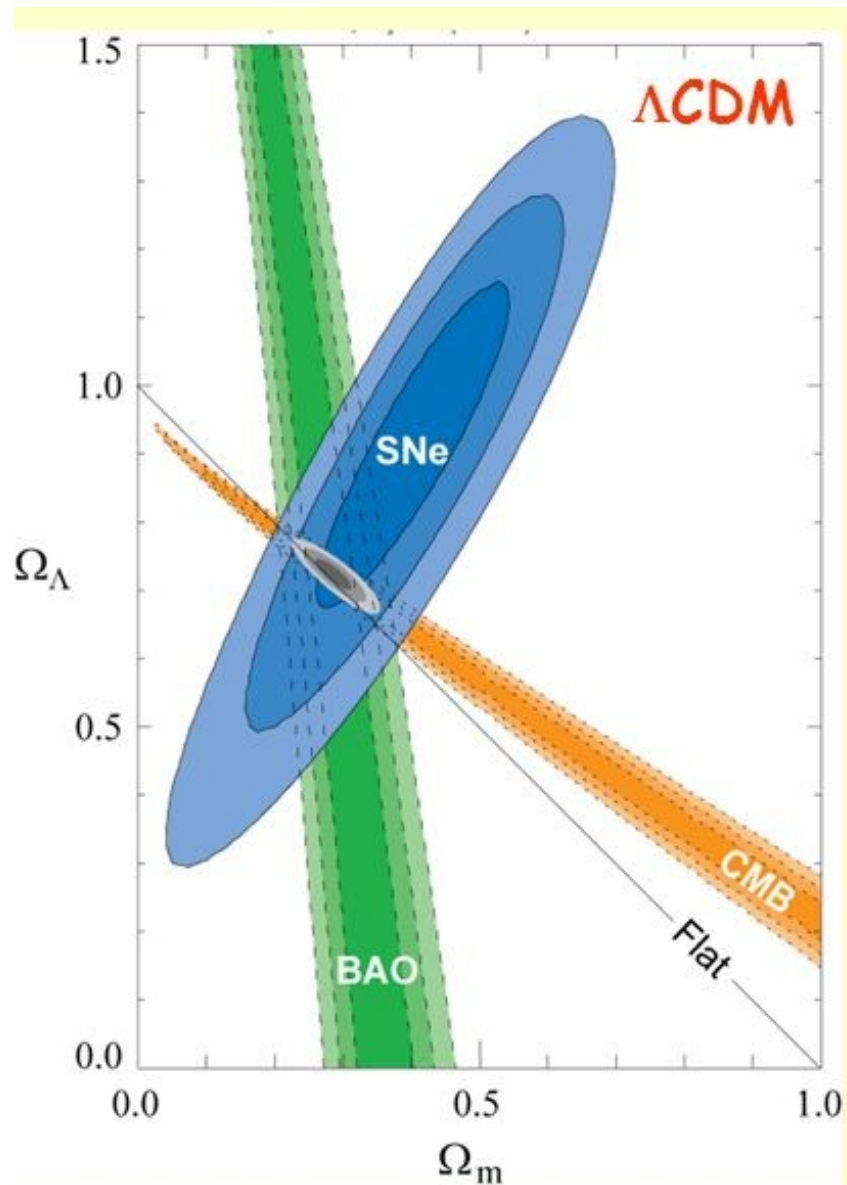
$$w = -1 \quad p = -\rho$$

Supernovae and Dark Energy

Betoule et al., 2014

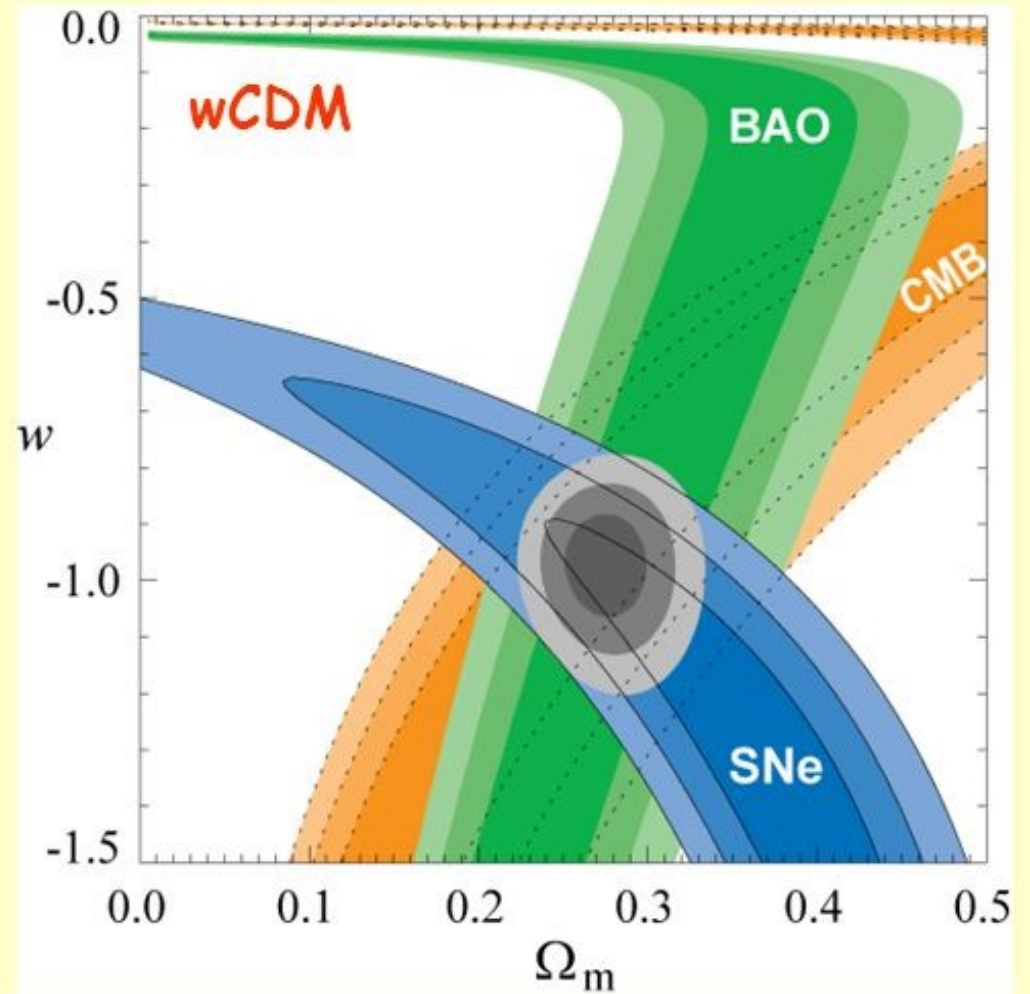


Concordance Model



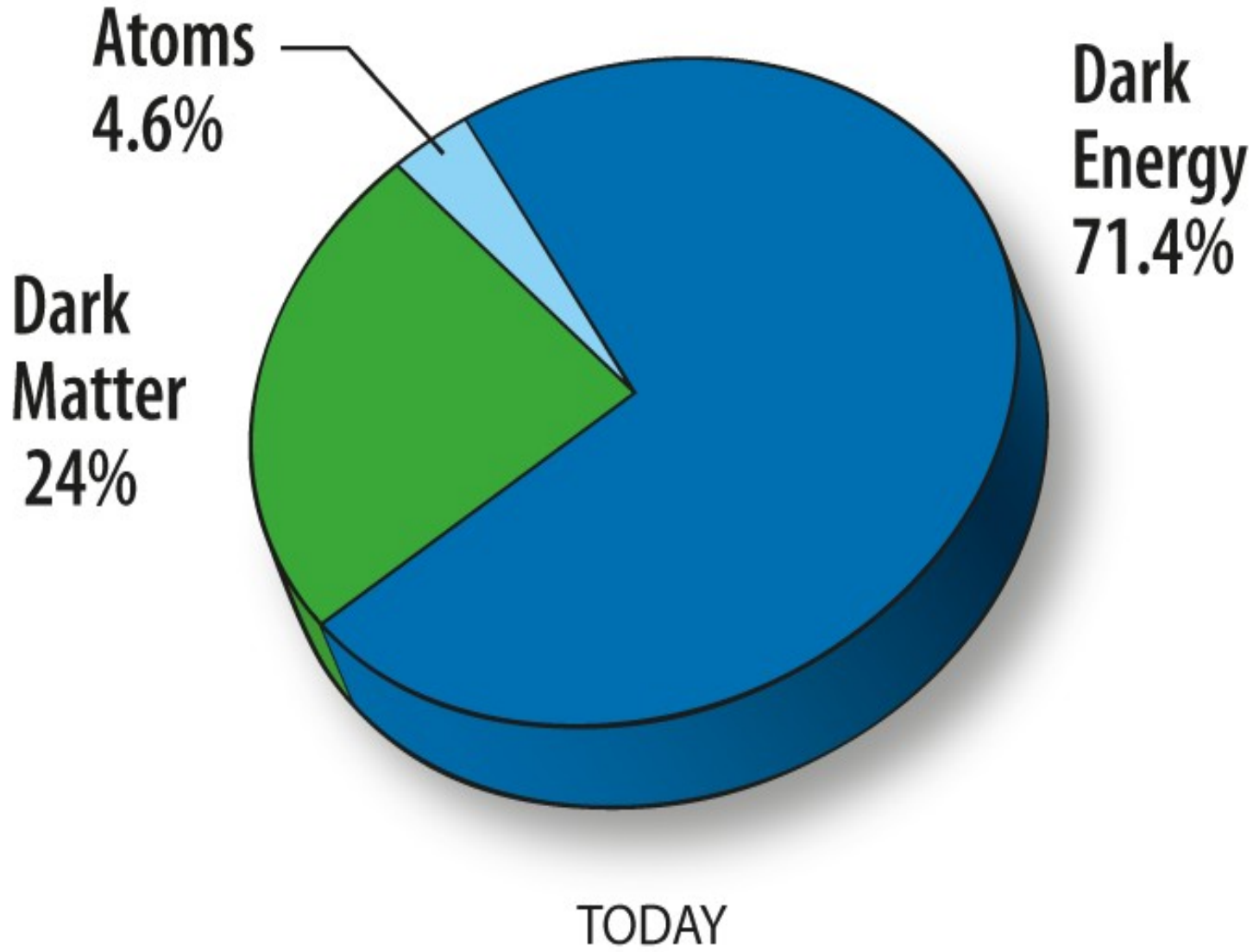
NB: CMB = WMAP-5, BAO = SDSS 2005

M. Kowalski et al., 2008, ApJ, 686, 749



$$w = -0.969^{+0.059+0.063}_{-0.063-0.066}$$

A strange Universe...



A short history of the Universe

